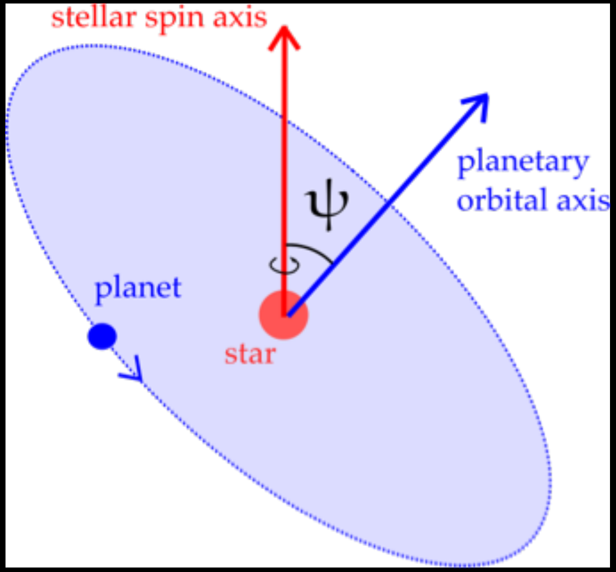
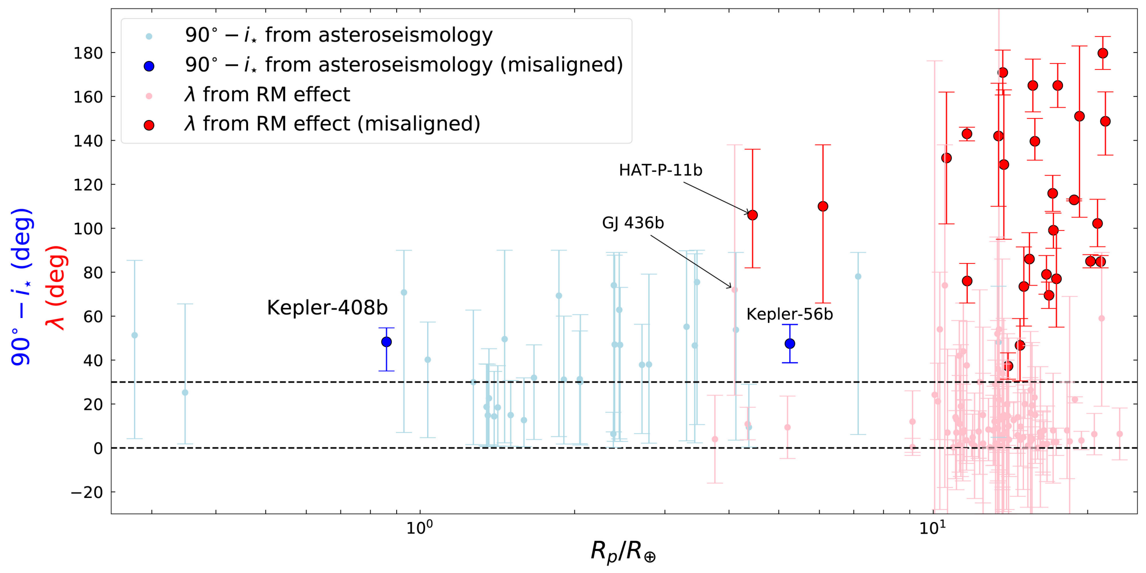
Spin-orbit architecture of transiting planetary systems probed with the Rossiter-McLaughlin effect and asteroseismology





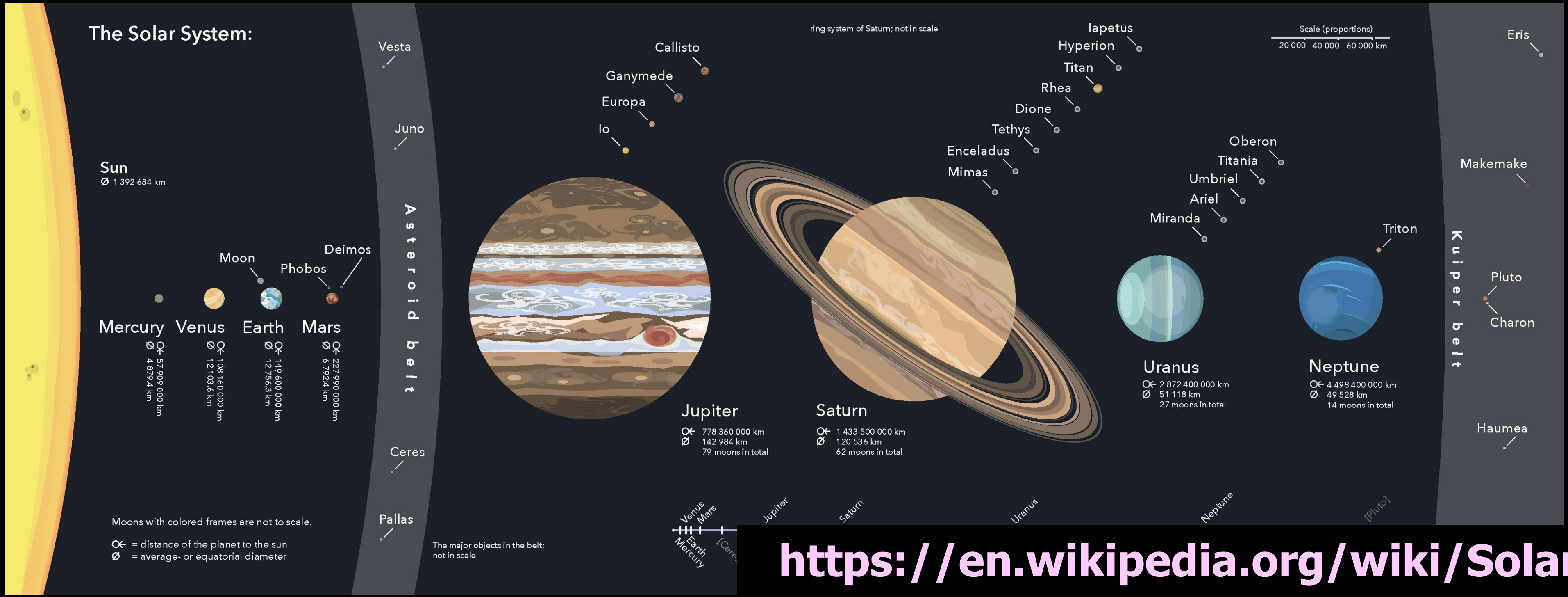
Yasushi Suto

Department of Physics and Research Center for the Early Universe
The University of Tokyo

11:00-12:00 June 14, 2019 @ Laboratoire d’astrophysique de Bordeaux, Université de Bordeaux

The Rossiter-McLaughlin effect

Architecture of the Solar system



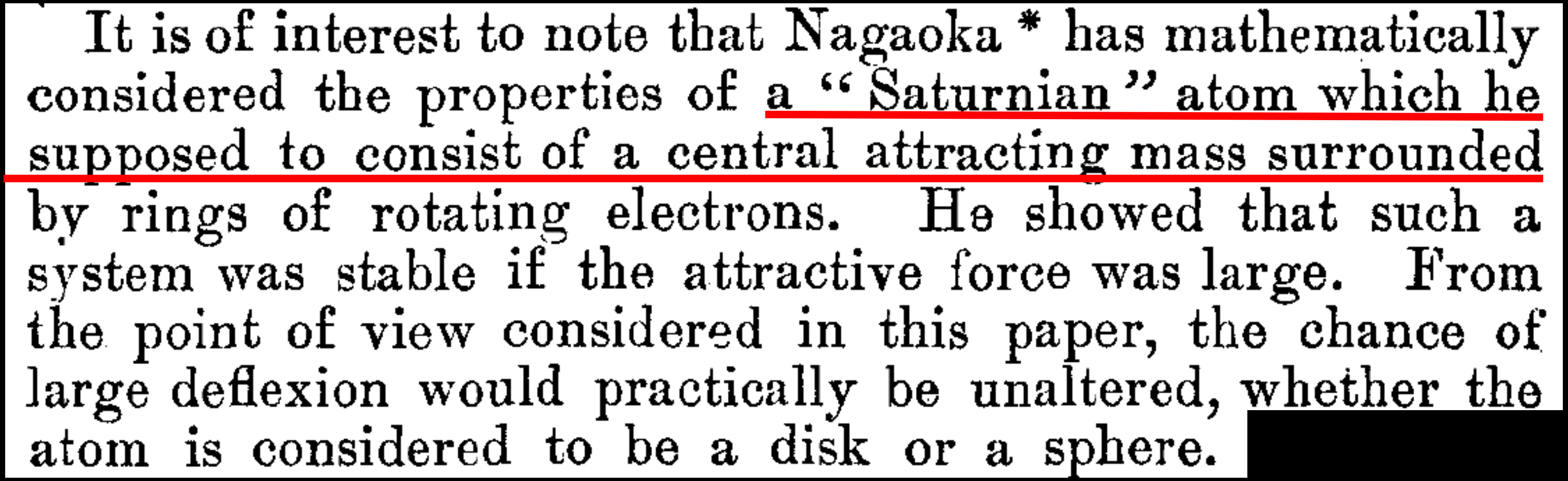
https://en.wikipedia.org/wiki/Solar_System

- Our Solar system is typical or atypical ?
  - Very stable multiplanetary systems on nearly co-planar and circular orbits
  - Rocky inner planets + Gaseous outer planets
  - satellites and rings are fairly common
  - A planet with life and (advanced) civilization

From Saturnian model to atomic model

- Saturnian architecture inspired the model of atomic structure
  - H.Nagaoka: Phil. Mag. 7(1904) 445

Ernest Rutherford: *The Scattering of α and β Particles by Matter and the Structure of the Atom* Phil. Mag. 6(1911) 669

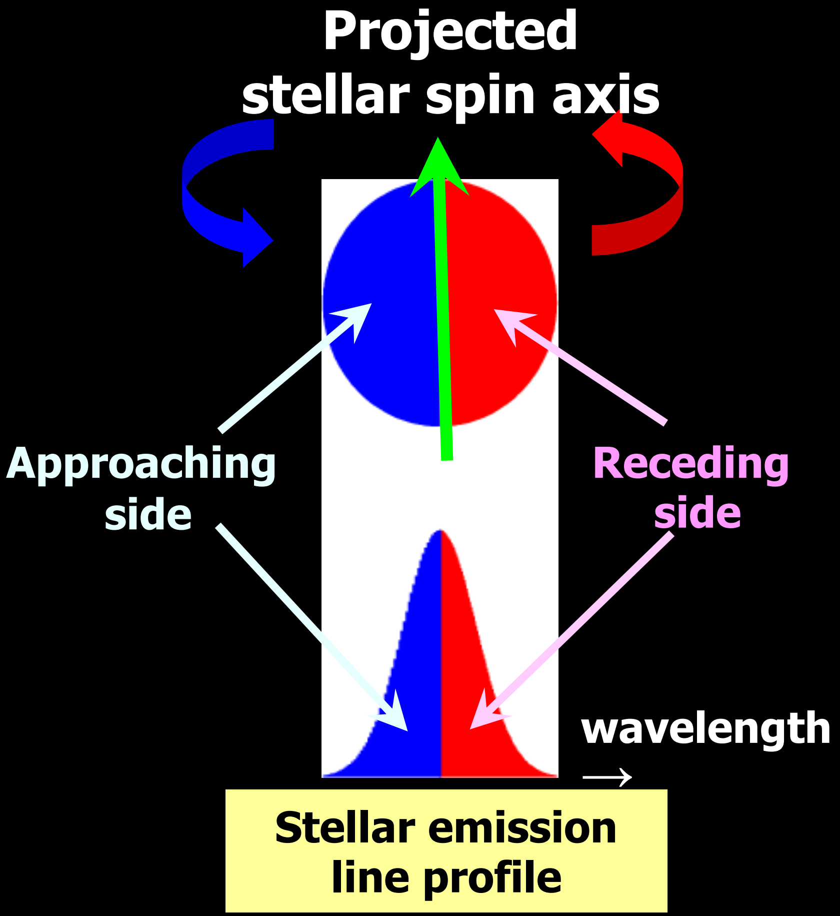


It is of interest to note that Nagaoka * has mathematically considered the properties of a "Saturnian" atom which he supposed to consist of a central attracting mass surrounded by rings of rotating electrons. He showed that such a system was stable if the attractive force was large. From the point of view considered in this paper, the chance of large deflexion would practically be unaltered, whether the atom is considered to be a disk or a sphere.

From atomic model to architecture of exoplanetary systems

Ang. Mom.	Atomic system	Exoplanetary system
L	Quantized energy levels Emission/absorption line transition	Spectroscopic radial velocity Transit photometry, Microlensing Orbital period, semi-major axis, eccentricity, planetary mass
S	Spin of nucleus Hyperfine structure splitting	Rossiter-McLaughlin effect Asteroseismology Stellar spin - planetary orbit angle Stellar spin obliquity
s	Spin of electrons Fine structure splitting	Tidal interaction between star and planet Planetary spin, planetary ring

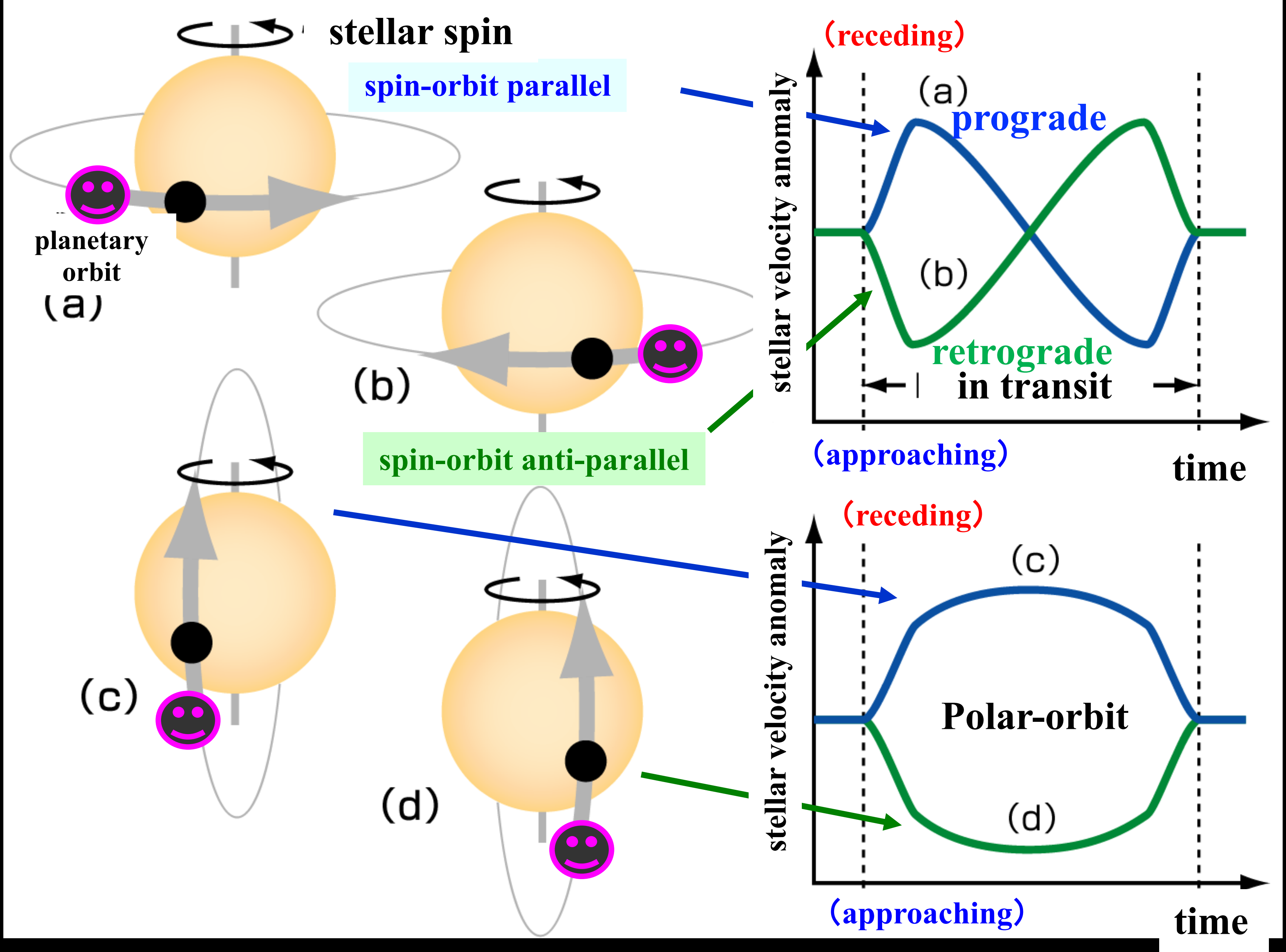
Spectroscopic transit signature: the Rossiter-McLaughlin effect



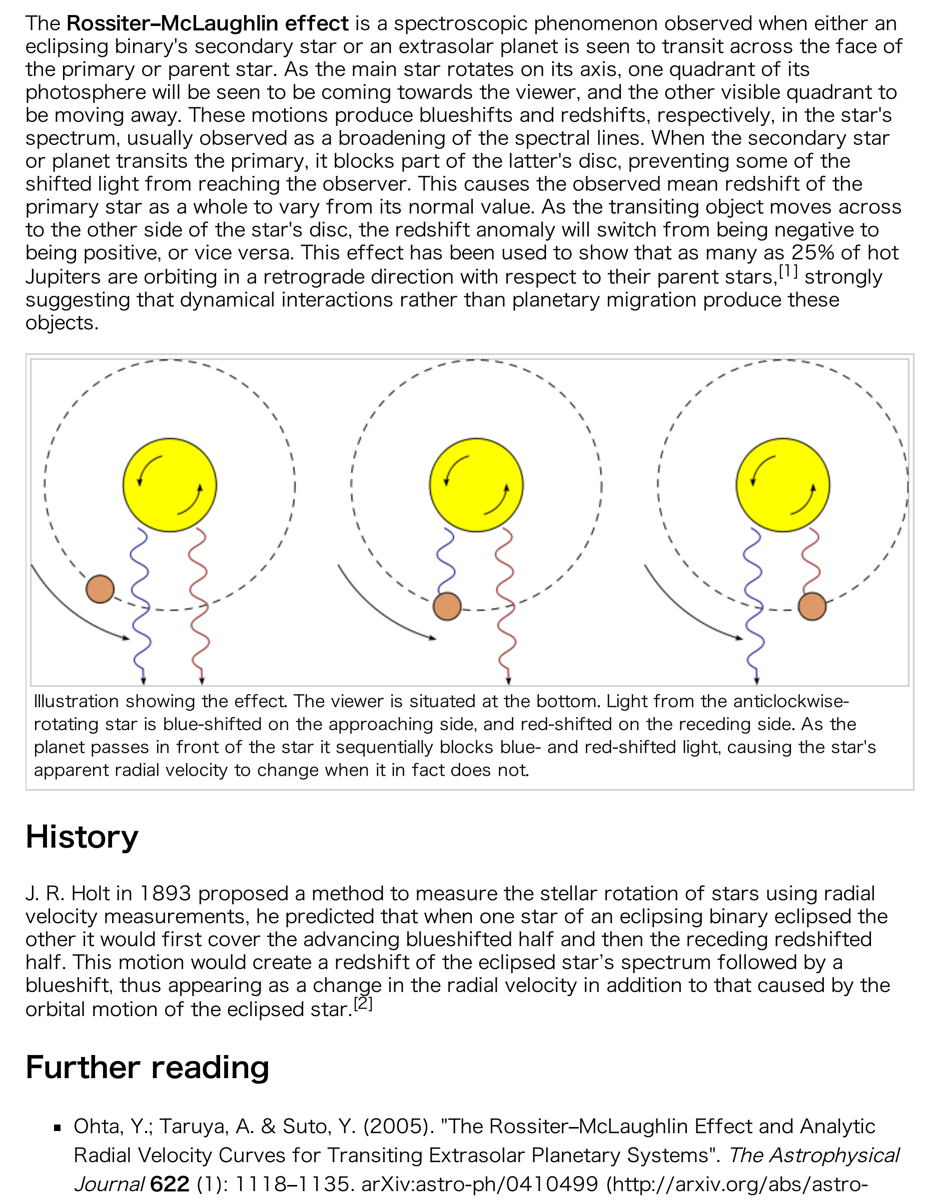
- Time-dependent asymmetry in the stellar Doppler broadened line profile
  - apparent anomaly of the stellar radial velocity
- originally proposed for eclipsing binaries

Holt, J.R. Astronomy and Astrophysics 12(1893)646
Rossiter, ApJ 60(1924)15; McLaughlin, ApJ 60 (1924)20
Hosokawa, PASJ 5(1953)88; Ohta, Taruya + YS, ApJ 622(2005)1118

Velocity anomaly due to the Rossiter-McLaughlin effect



The Rossiter-McLaughlin effect @ Wikipedia



The **Rossiter–McLaughlin effect** is a spectroscopic phenomenon observed when either an eclipsing binary's secondary star or an extrasolar planet is seen to transit across the face of the primary or parent star. As the main star rotates on its axis, one quadrant of its photosphere will be seen to be coming towards the viewer, and the other visible quadrant to be moving away. These motions produce blueshifts and redshifts, respectively, in the star's spectrum, usually observed as a broadening of the spectral lines. When the secondary star or planet transits the primary, it blocks part of the latter's disc, preventing some of the shifted light from reaching the observer. This causes the observed mean redshift of the primary star as a whole to vary from its normal value. As the transiting object moves across to the other side of the star's disc, the redshift anomaly will switch from being negative to being positive, or vice versa. This effect has been used to show that as many as 25% of hot Jupiters are orbiting in a retrograde direction with respect to their parent stars,[1] strongly suggesting that dynamical interactions rather than planetary migration produce these objects.

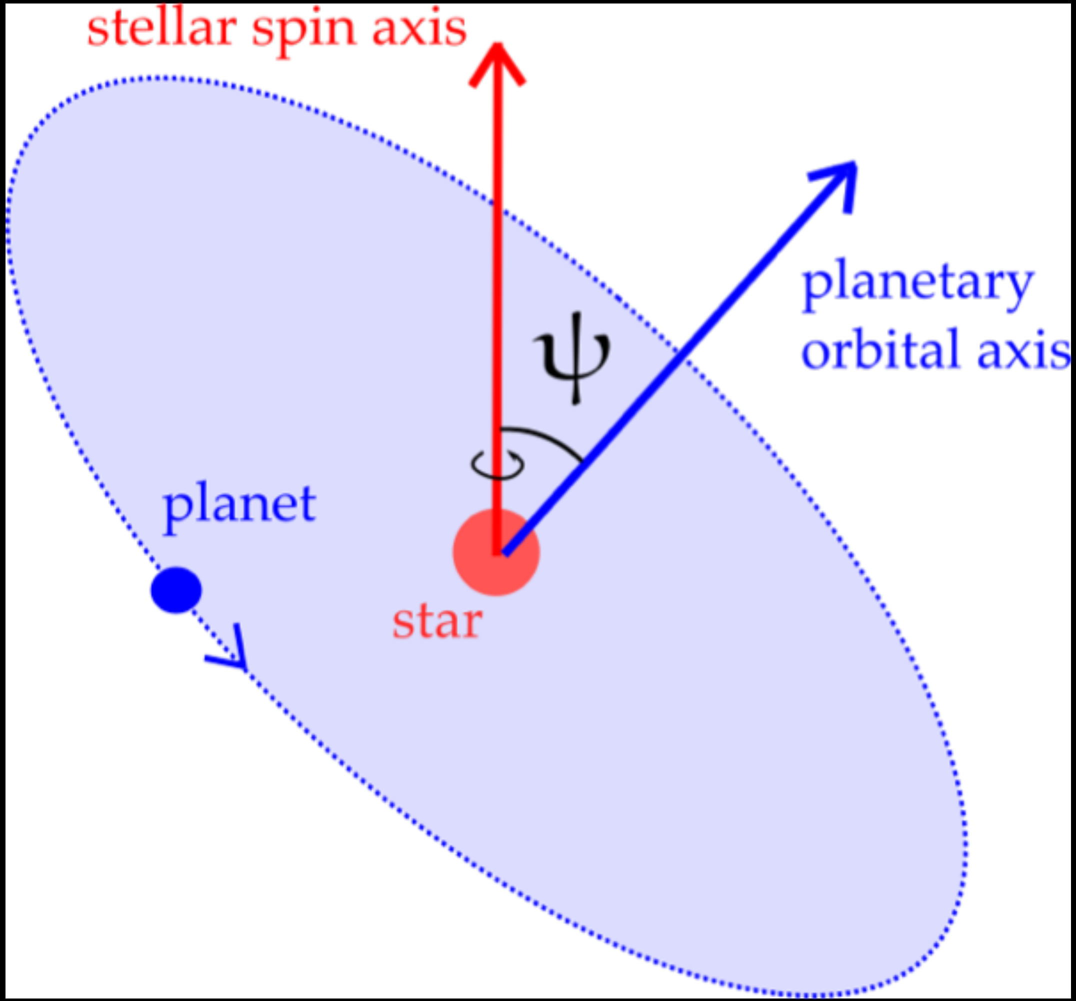
Illustration showing the effect. The viewer is situated at the bottom. Light from the anticlockwise-rotating star is blue-shifted on the approaching side, and red-shifted on the receding side. As the planet passes in front of the star it sequentially blocks blue- and red-shifted light, causing the star's apparent radial velocity to change when it in fact does not.

History

J. R. Holt in 1893 proposed a method to measure the stellar rotation of stars using radial velocity measurements, he predicted that when one star of an eclipsing binary eclipsed the other it would first cover the advancing blueshifted half and then the receding redshifted half. This motion would create a redshift of the eclipsed star's spectrum followed by a blueshift, thus appearing as a change in the radial velocity in addition to that caused by the orbital motion of the eclipsed star.[2]

Further reading

- Ohta, Y.; Taruya, A. & Suto, Y. (2005). "The Rossiter–McLaughlin Effect and Analytic Radial Velocity Curves for Transiting Extrasolar Planetary Systems". *The Astrophysical Journal* **622** (1): 1118–1135. arXiv:astro-ph/0410499 (http://arxiv.org/abs/astro-

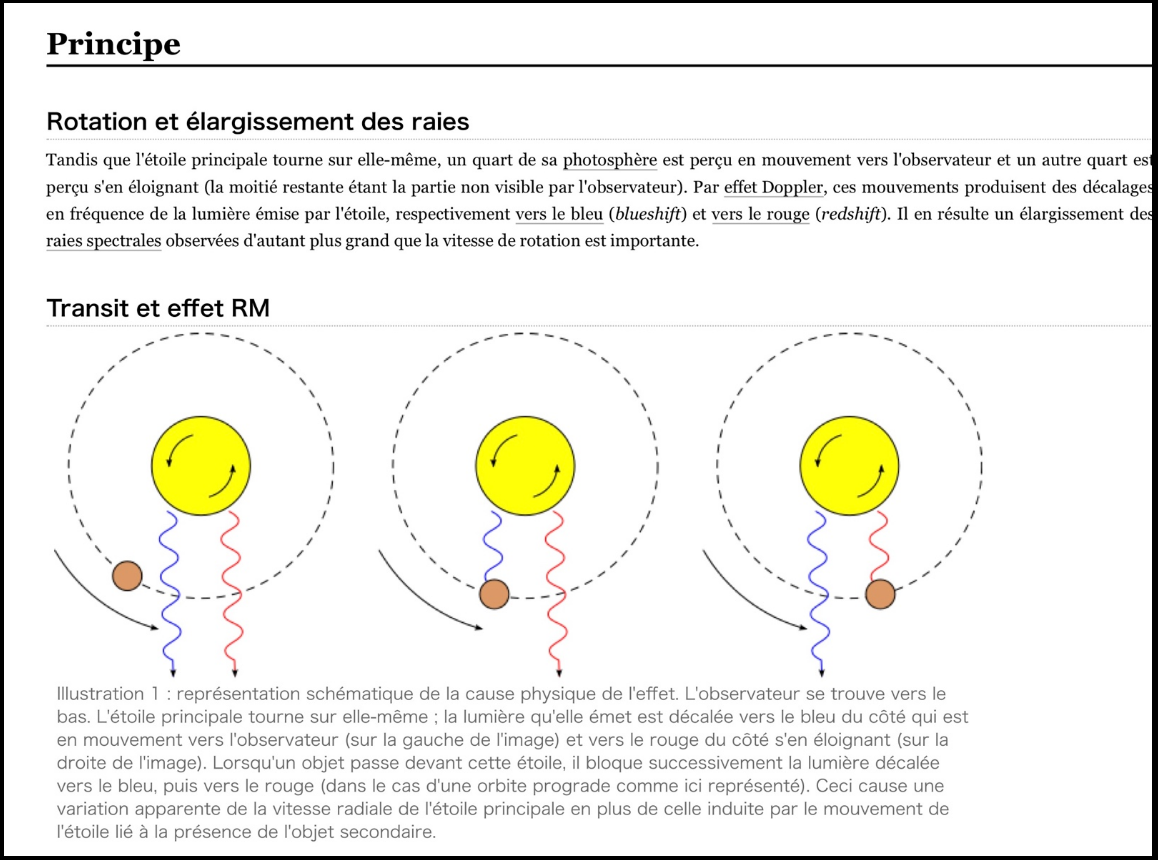


L’effet Rossiter-McLaughlin

WIKIPÉDIA

Effet Rossiter-McLaughlin

L'**effet Rossiter-McLaughlin** (en abrégé « effet RM ») est un phénomène spectroscopique observé lorsqu'un corps éclipsant (qui peut être une étoile secondaire ou une exoplanète) transite devant la surface de l'étoile primaire autour de laquelle il effectue son orbite.



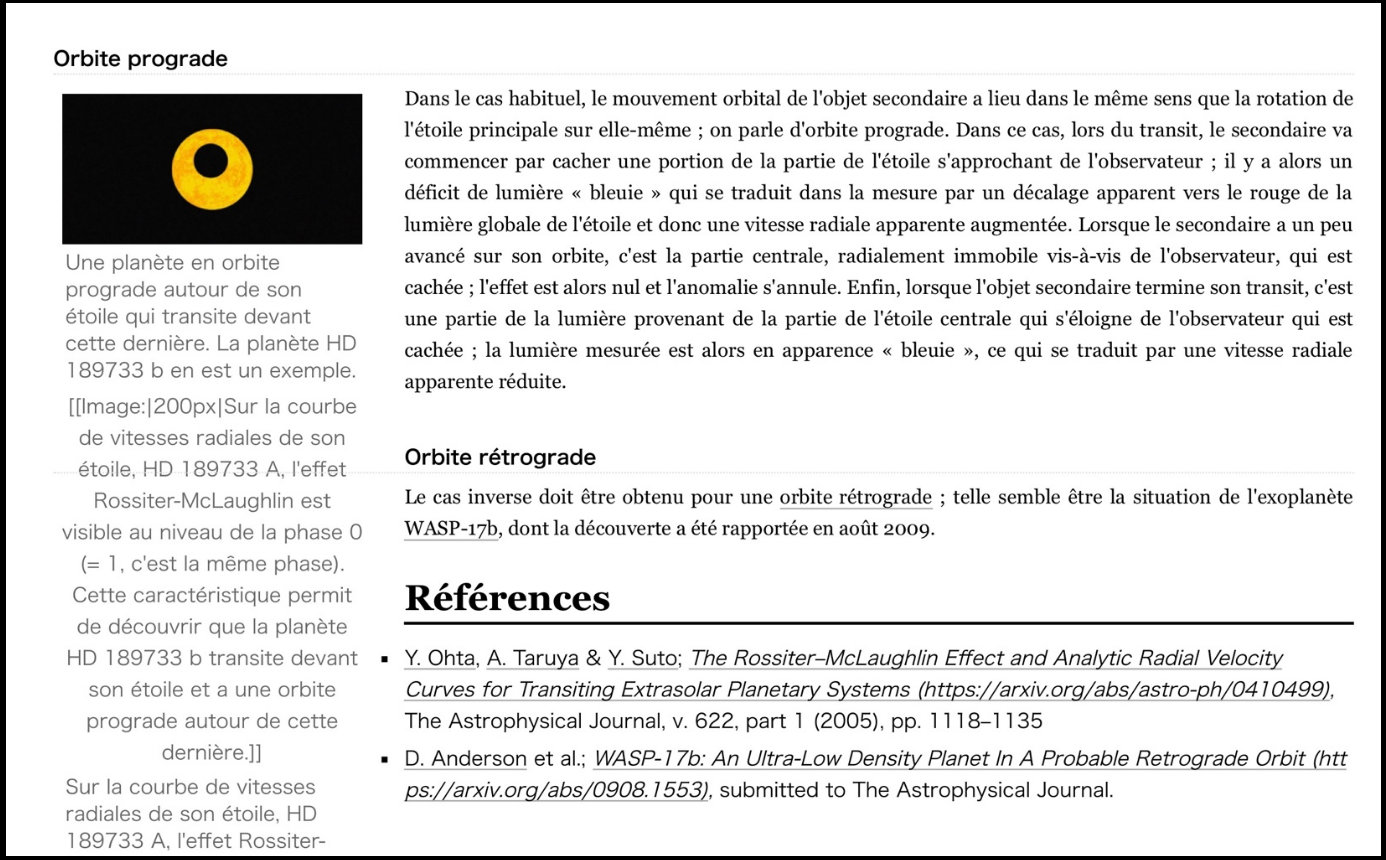
Principe

Rotation et élargissement des raies

Tandis que l'étoile principale tourne sur elle-même, un quart de sa photosphère est perçu en mouvement vers l'observateur et un autre quart est perçu s'en éloignant (la moitié restante étant la partie non visible par l'observateur). Par effet Doppler, ces mouvements produisent des décalages en fréquence de la lumière émise par l'étoile, respectivement vers le bleu (*blueshift*) et vers le rouge (*redshift*). Il en résulte un élargissement des raies spectrales observées d'autant plus grand que la vitesse de rotation est importante.

Transit et effet RM

Illustration 1 : représentation schématique de la cause physique de l'effet. L'observateur se trouve vers le bas. L'étoile principale tourne sur elle-même ; la lumière qu'elle émet est décalée vers le bleu du côté qui est en mouvement vers l'observateur (sur la gauche de l'image) et vers le rouge du côté s'en éloignant (sur la droite de l'image). Lorsqu'un objet passe devant cette étoile, il bloque successivement la lumière décalée vers le bleu, puis vers le rouge (dans le cas d'une orbite prograde comme ici représenté). Ceci cause une variation apparente de la vitesse radiale de l'étoile principale en plus de celle induite par le mouvement de l'étoile lié à la présence de l'objet secondaire.



Orbite prograde

Une planète en orbite prograde autour de son étoile qui transite devant cette dernière. La planète HD 189733 b en est un exemple.

[[Image:|200px|Sur la courbe de vitesses radiales de son étoile, HD 189733 A, l'effet Rossiter-McLaughlin est visible au niveau de la phase 0 (= 1, c'est la même phase). Cette caractéristique permit de découvrir que la planète HD 189733 b transite devant son étoile et a une orbite prograde autour de cette dernière.]]

Sur la courbe de vitesses radiales de son étoile, HD 189733 A, l'effet Rossiter-

Dans le cas habituel, le mouvement orbital de l'objet secondaire a lieu dans le même sens que la rotation de l'étoile principale sur elle-même ; on parle d'orbite prograde. Dans ce cas, lors du transit, le secondaire va commencer par cacher une portion de la partie de l'étoile s'approchant de l'observateur ; il y a alors un déficit de lumière « bleuie » qui se traduit dans la mesure par un décalage apparent vers le rouge de la lumière globale de l'étoile et donc une vitesse radiale apparente augmentée. Lorsque le secondaire a un peu avancé sur son orbite, c'est la partie centrale, radialement immobile vis-à-vis de l'observateur, qui est cachée ; l'effet est alors nul et l'anomalie s'annule. Enfin, lorsque l'objet secondaire termine son transit, c'est une partie de la lumière provenant de la partie de l'étoile centrale qui s'éloigne de l'observateur qui est cachée ; la lumière mesurée est alors en apparence « bleuie », ce qui se traduit par une vitesse radiale apparente réduite.

Orbite rétrograde

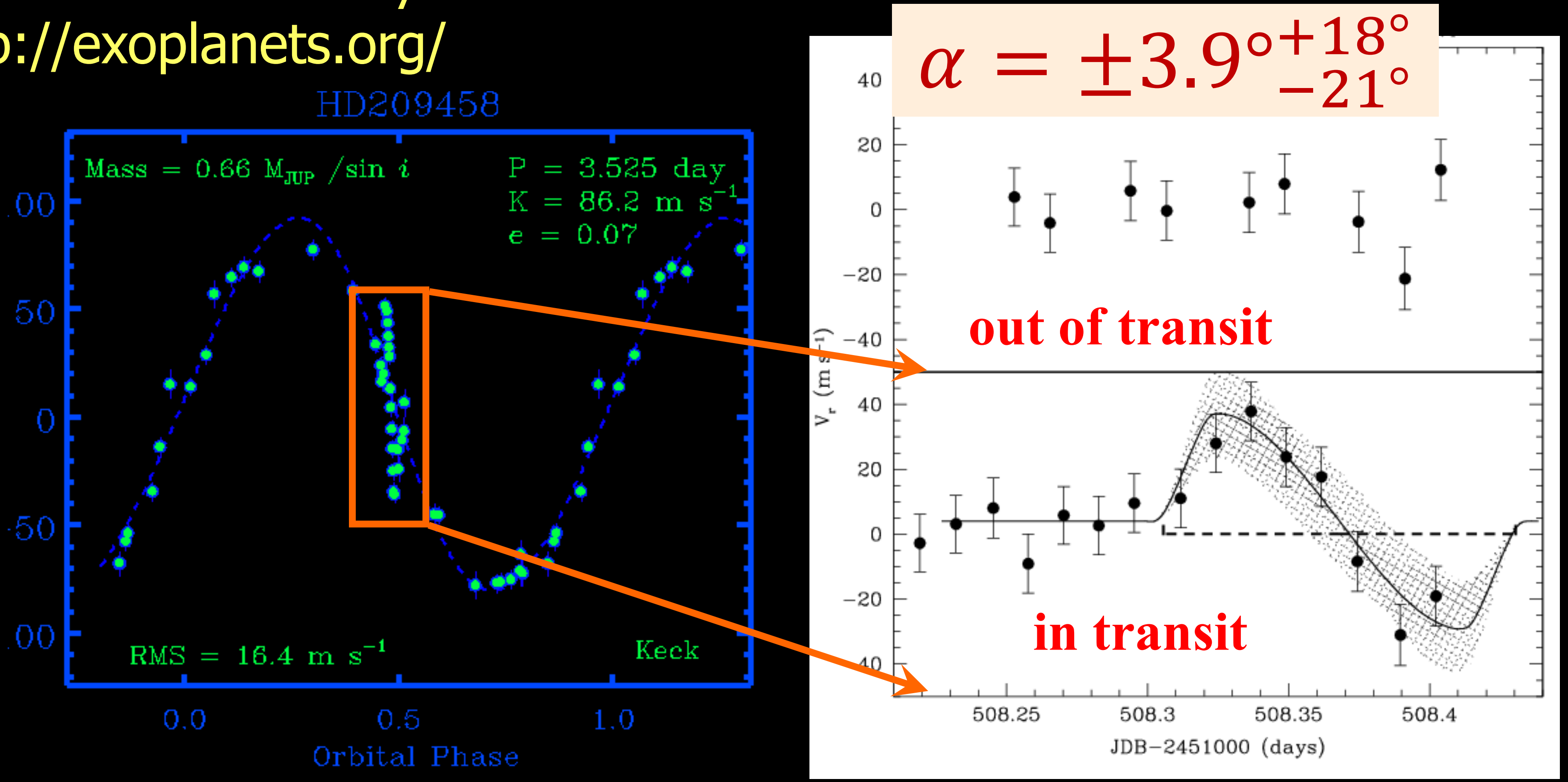
Le cas inverse doit être obtenu pour une orbite rétrograde ; telle semble être la situation de l'exoplanète WASP-17b, dont la découverte a été rapportée en août 2009.

Références

- Y. Ohta, A. Taruya & Y. Suto; *The Rossiter–McLaughlin Effect and Analytic Radial Velocity Curves for Transiting Extrasolar Planetary Systems (https://arxiv.org/abs/astro-ph/0410499)*, The Astrophysical Journal, v. 622, part 1 (2005), pp. 1118–1135
- D. Anderson et al.; *WASP-17b: An Ultra-Low Density Planet In A Probable Retrograde Orbit (https://arxiv.org/abs/0908.1553)*, submitted to The Astrophysical Journal.

The first detection of the RM effect: HD209458

HD209458 radial velocity data
http://exoplanets.org/

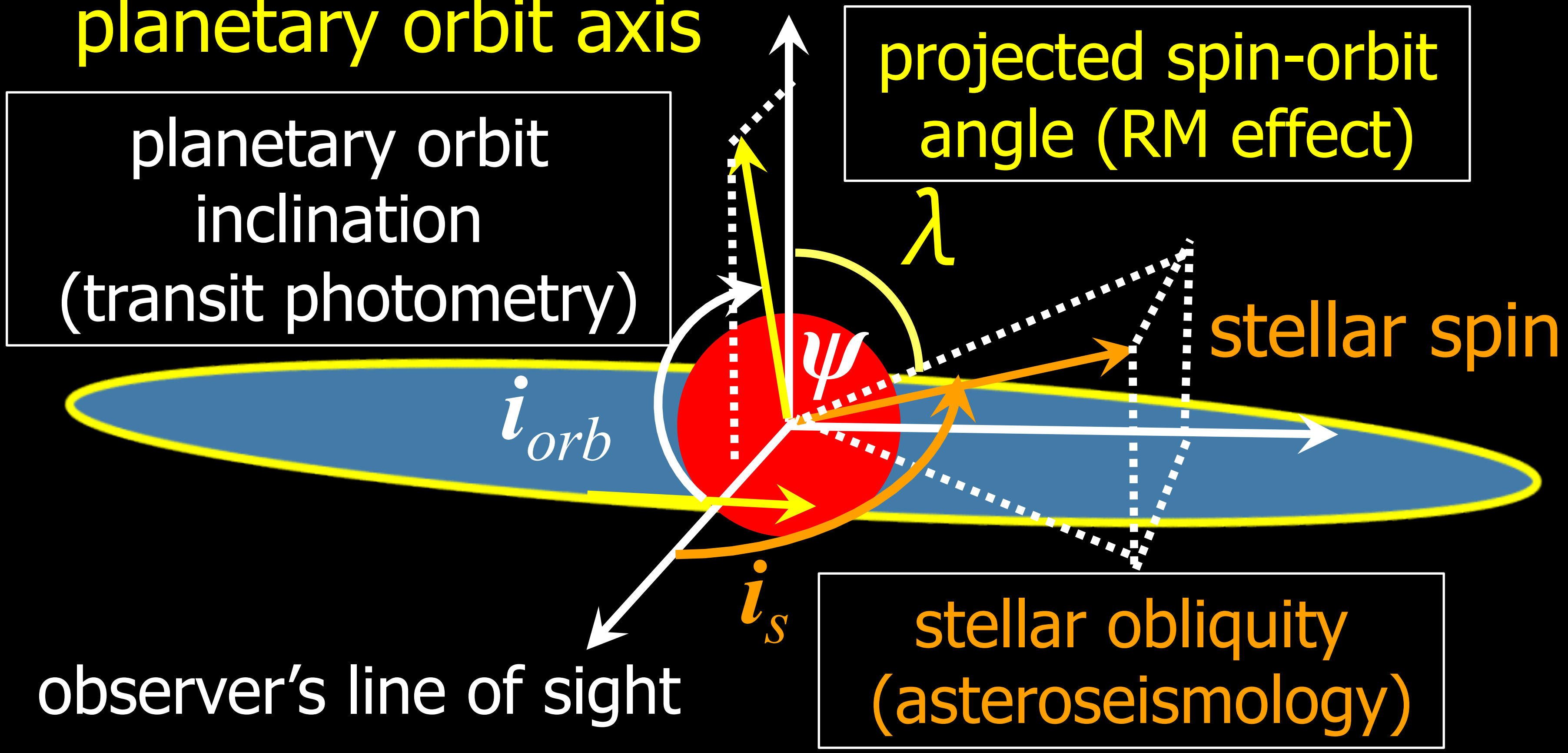


$$\alpha = \pm 3.9^{\circ}{}^{+18^{\circ}}_{-21^{\circ}}$$

(This is not their original data in 2000)

Stellar rotation and planetary orbit
Queloz et al. (2000) A&A 359, L13
ELODIE on 193cm telescope

Spin-orbit angles of a transiting planet



planetary orbit axis

planetary orbit inclination (transit photometry)

projected spin-orbit angle (RM effect)

stellar spin

stellar obliquity (asteroseismology)

observer's line of sight

$$\cos\Psi = \sin i_s \sin i_{orb} \cos\lambda + \cos i_s \cos i_{orb}$$
$$\approx \sin i_s \cos\lambda$$

Origin of the spin-orbit misalignment ?

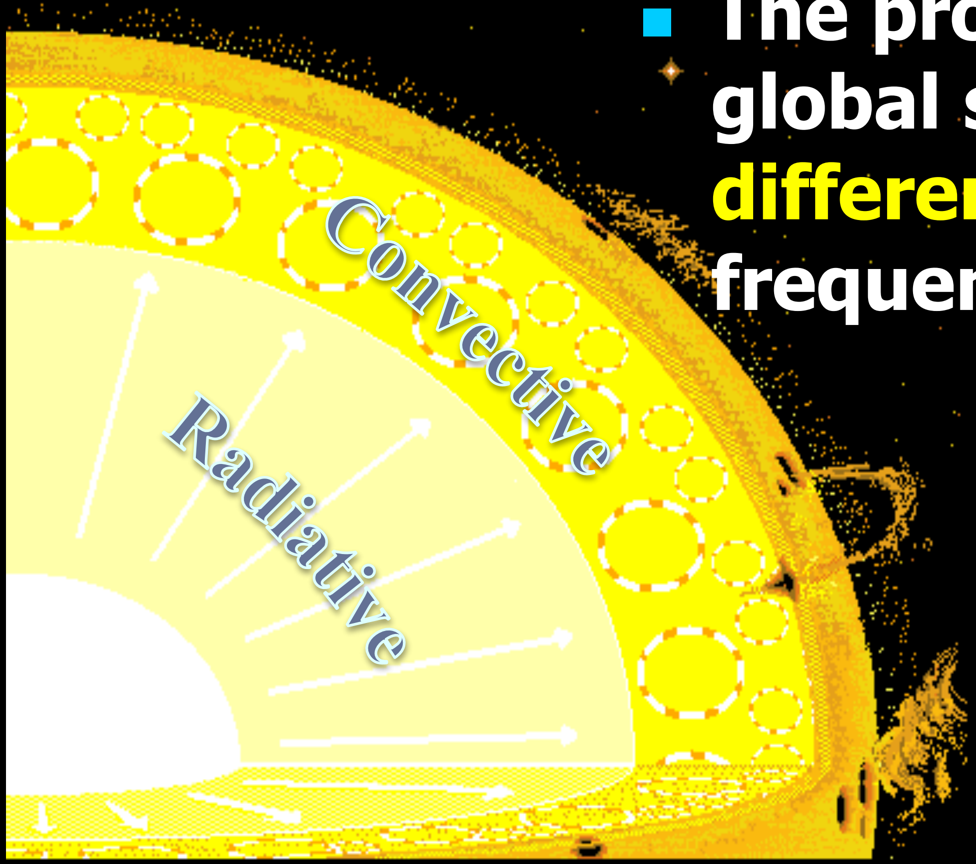
- Occurrence rate of misalignment from numerical simulations ? (large uncertainty of the initial configuration of planets)
- Efficiency of tidal realignment by convective zone of stars with $T_{eff}<6100K$?
- Complementary statistics from stellar obliquity with/without planets → asteroseismology
- Difference between single- and multi- transiting planetary systems → asteroseismology

Asteroseismology

Oscillations of Sun-like stars

$(0.8M_{\odot} < M < 2.5\ M_{\odot})$

- **Convection triggers oscillation waves inside stars**
- **The propagating waves form global standing waves with different eigenmode frequencies**



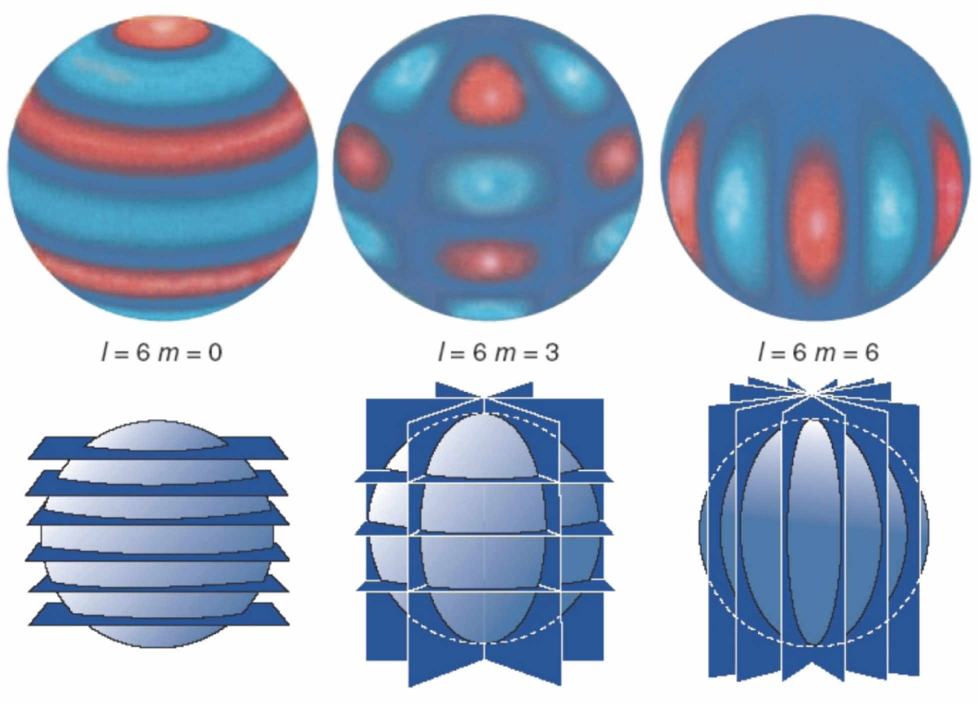
- The induced temperature perturbations are measured through the stellar photometric pulsation

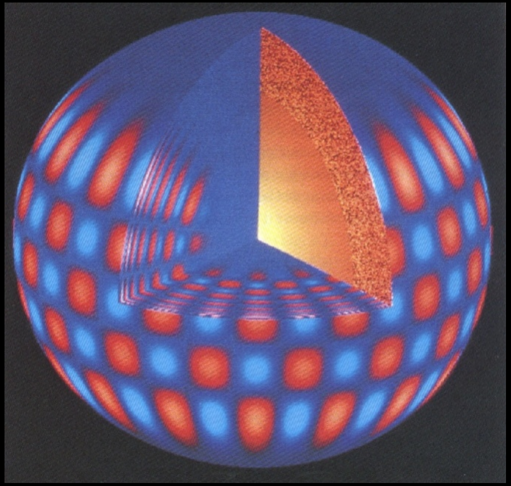
Characterizing the stellar pulsations

- Expansion in terms of spherical harmonics

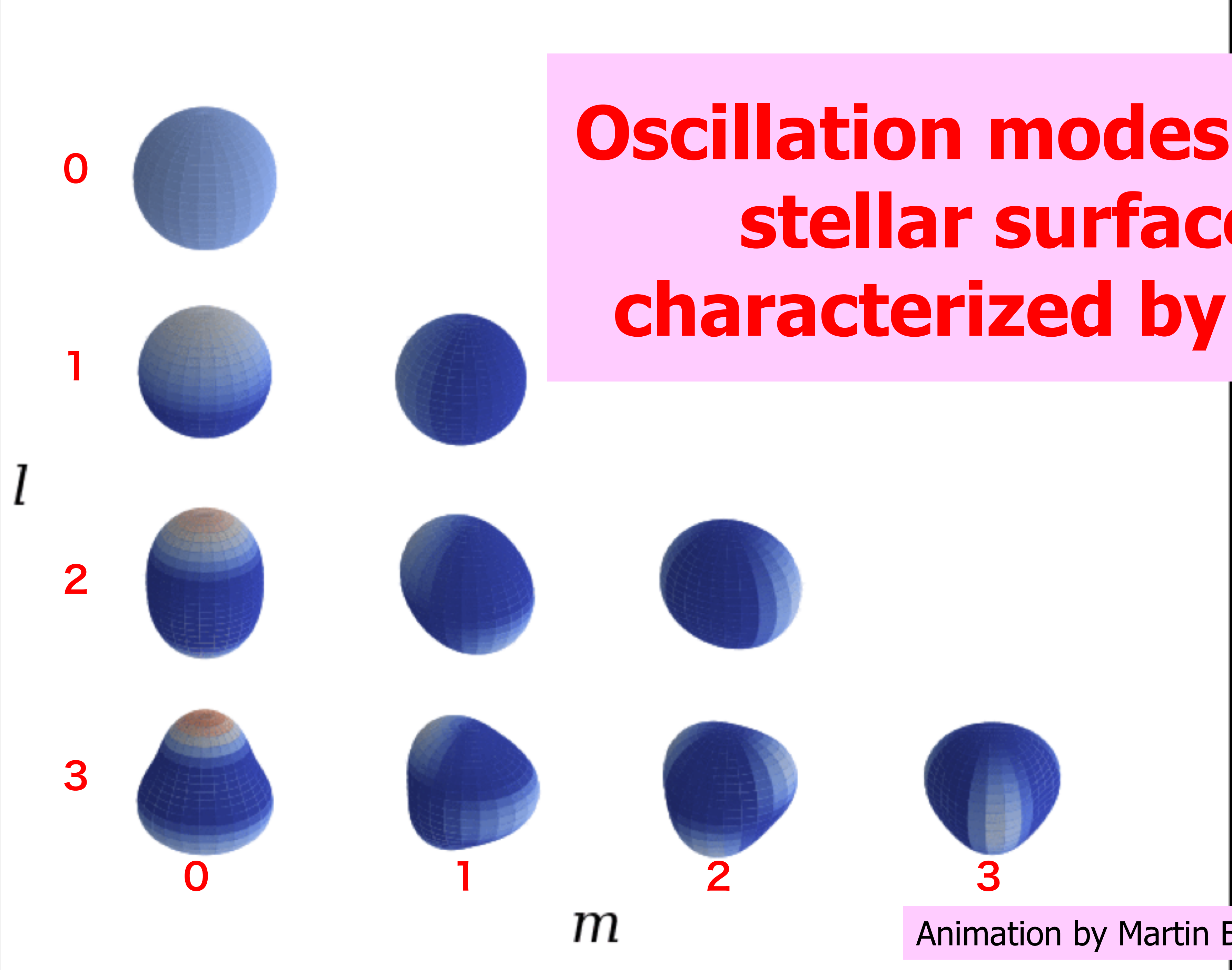
$$Y_{lm}(\theta,\varphi) \propto P_l^{|m|}(\cos\theta)\, e^{im\varphi}$$

- Three integers to characterize the mode
  - *n* radial order
  - *l* angular degree
  - *m* azimuthal order



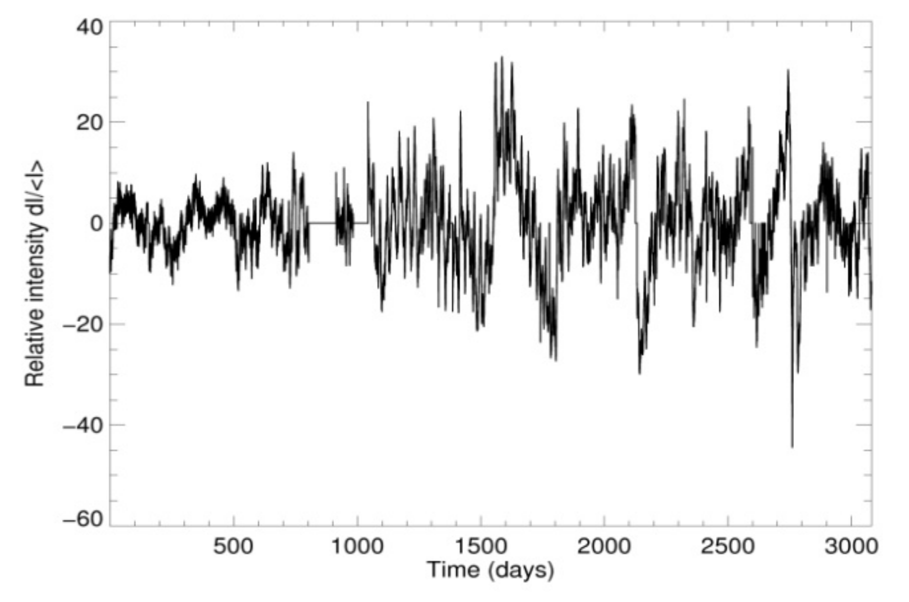


Oscillation modes of the stellar surfaces characterized by (l,m)



Animation by Martin B. Nielsen

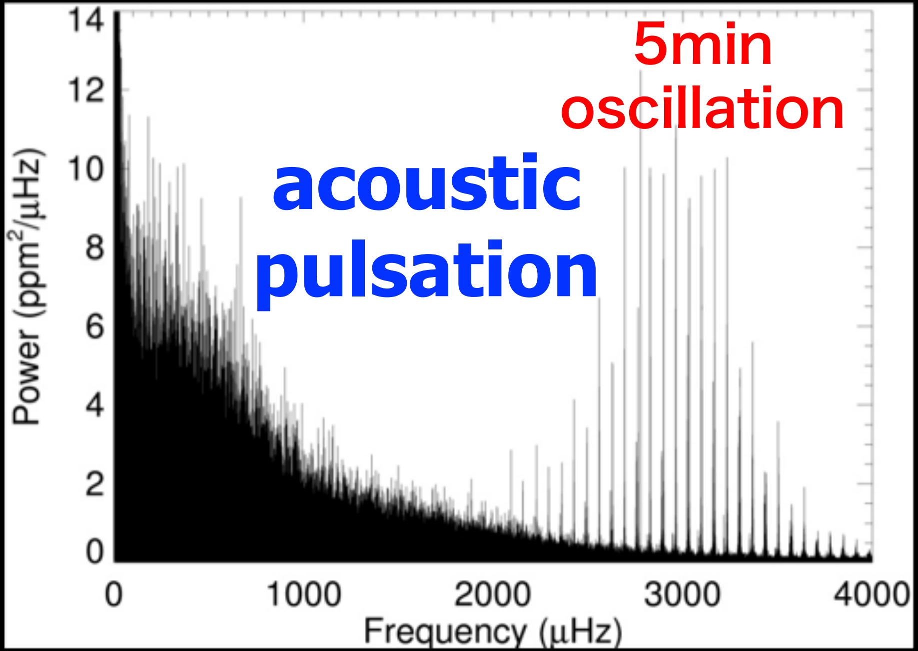
From lightcurve to power spectrum



Lightcurve of the Sun
in time domain

Fourier Transform

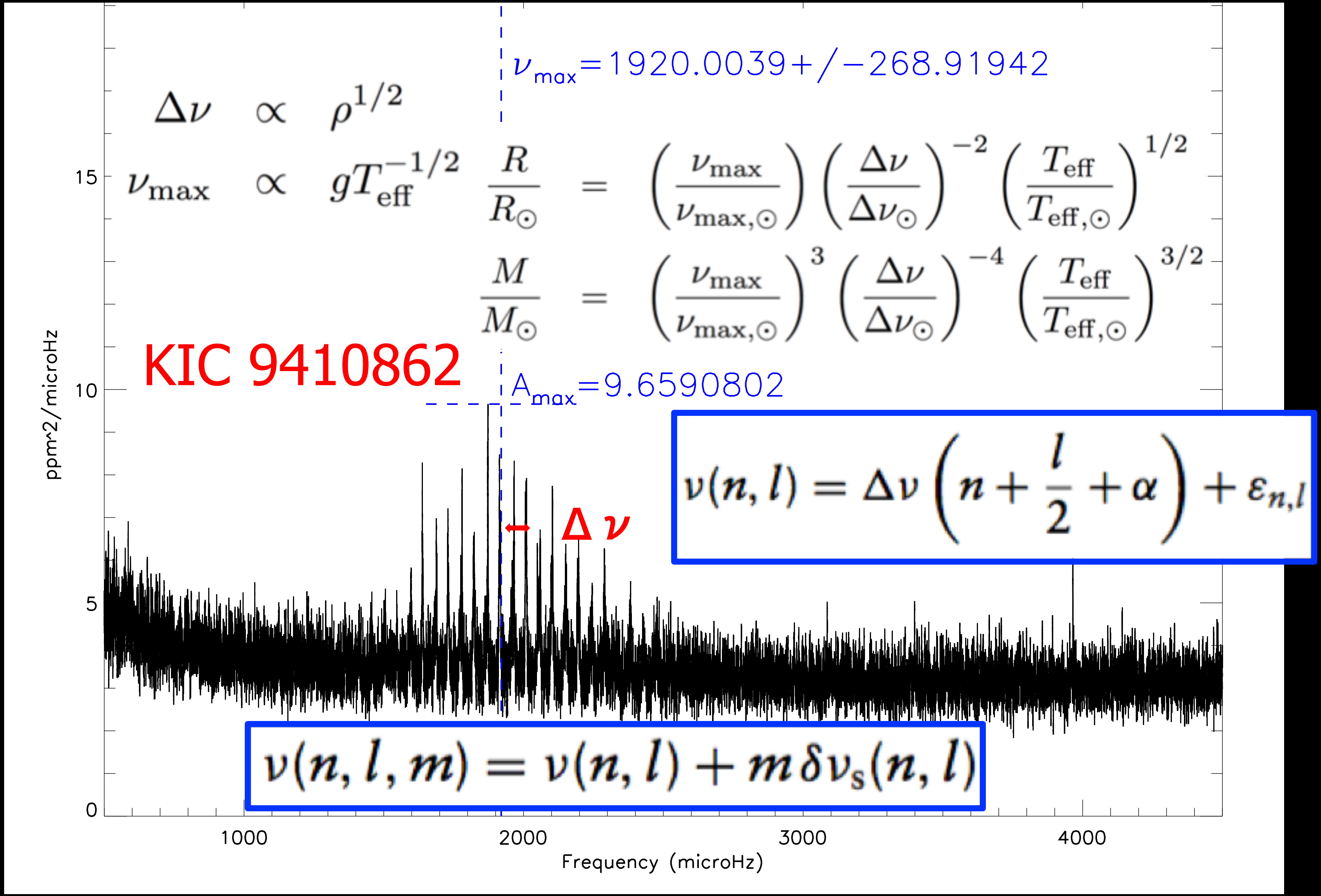
Power spectrum
in frequency domain



5min
oscillation

acoustic
pulsation

From oscillations to mass and radius



$$\Delta\nu \propto \rho^{1/2}$$

$$\nu_{\mathrm{max}} \propto g T_{\mathrm{eff}}^{-1/2}$$

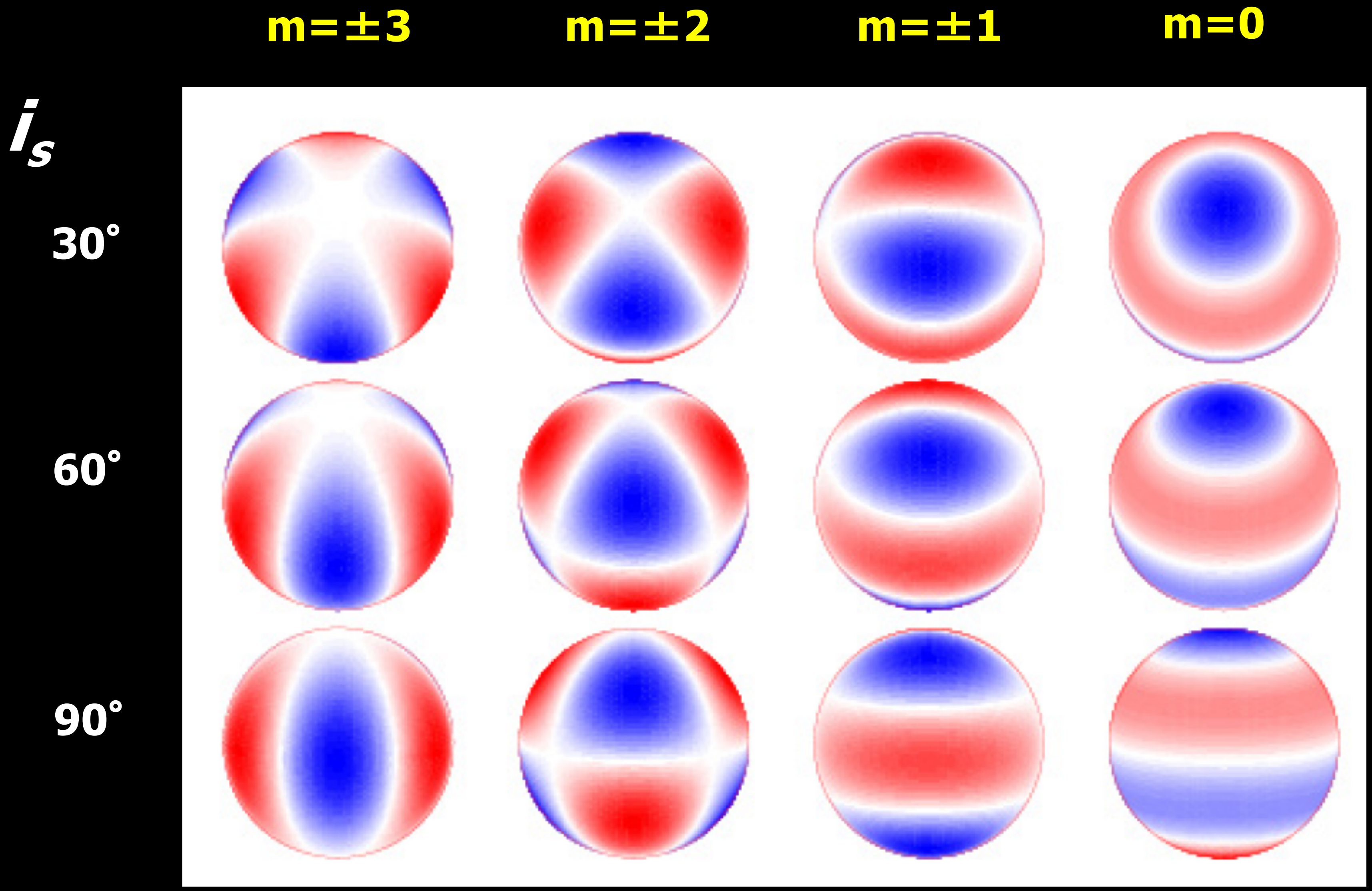
$$\frac{R}{R_\odot} = \left(\frac{\nu_{\mathrm{max}}}{\nu_{\mathrm{max},\odot}}\right)\left(\frac{\Delta\nu}{\Delta\nu_\odot}\right)^{-2}\left(\frac{T_{\mathrm{eff}}}{T_{\mathrm{eff},\odot}}\right)^{1/2}$$

$$\frac{M}{M_\odot} = \left(\frac{\nu_{\mathrm{max}}}{\nu_{\mathrm{max},\odot}}\right)^{3}\left(\frac{\Delta\nu}{\Delta\nu_\odot}\right)^{-4}\left(\frac{T_{\mathrm{eff}}}{T_{\mathrm{eff},\odot}}\right)^{3/2}$$

$$\boxed{\nu(n,l) = \Delta\nu\left(n + \frac{l}{2} + \alpha\right) + \varepsilon_{n,l}}$$

$$\boxed{\nu(n,l,m) = \nu(n,l) + m\,\delta\nu_s(n,l)}$$

Dependence on the stellar obliquity (l=3)



T.L. Campante, arXiv:1405.3145

Stellar obliquity and power spectrum

- **Oscillation in the corotating frame of the star**

$$\Psi_{nlm}(r,\theta,\varphi,t) = R_n(r)Y_{lm}(\theta,\varphi)e^{-iw_{nl}t} \propto e^{i(m\varphi - w_{nl}t)}$$

- **Oscillation frequency in the observer's frame**

$$\Psi_{nlm'}(r,\theta+i_*,\varphi-\Omega_*t,t) \propto e^{i(m'\varphi - m'\Omega_*t - w_{nl}t)}$$

- **Obliquity changes the amplitude of modes**

$$P(w) = \sum_{n,l}\sum_{m=-l}^{l} \frac{\mathcal{E}_{lm}(i_s)H_{nl}(w)}{1+4(\nu-\nu_{nlm})^2/\Gamma^2_{nlm}}$$

Toutain & Gouttebroze, (1993)
Gizon & Solanki (2003)
Kamiaka, Benomar & Suto (2018)

m-dependence of the mode amplitude

$$\mathcal{E}_{lm}(i_s) = \frac{(l-|m|)!}{(l+|m|)!}\left[P_l^{|m|}(\cos i_s)\right]^2$$

m-dependence of the mode frequency

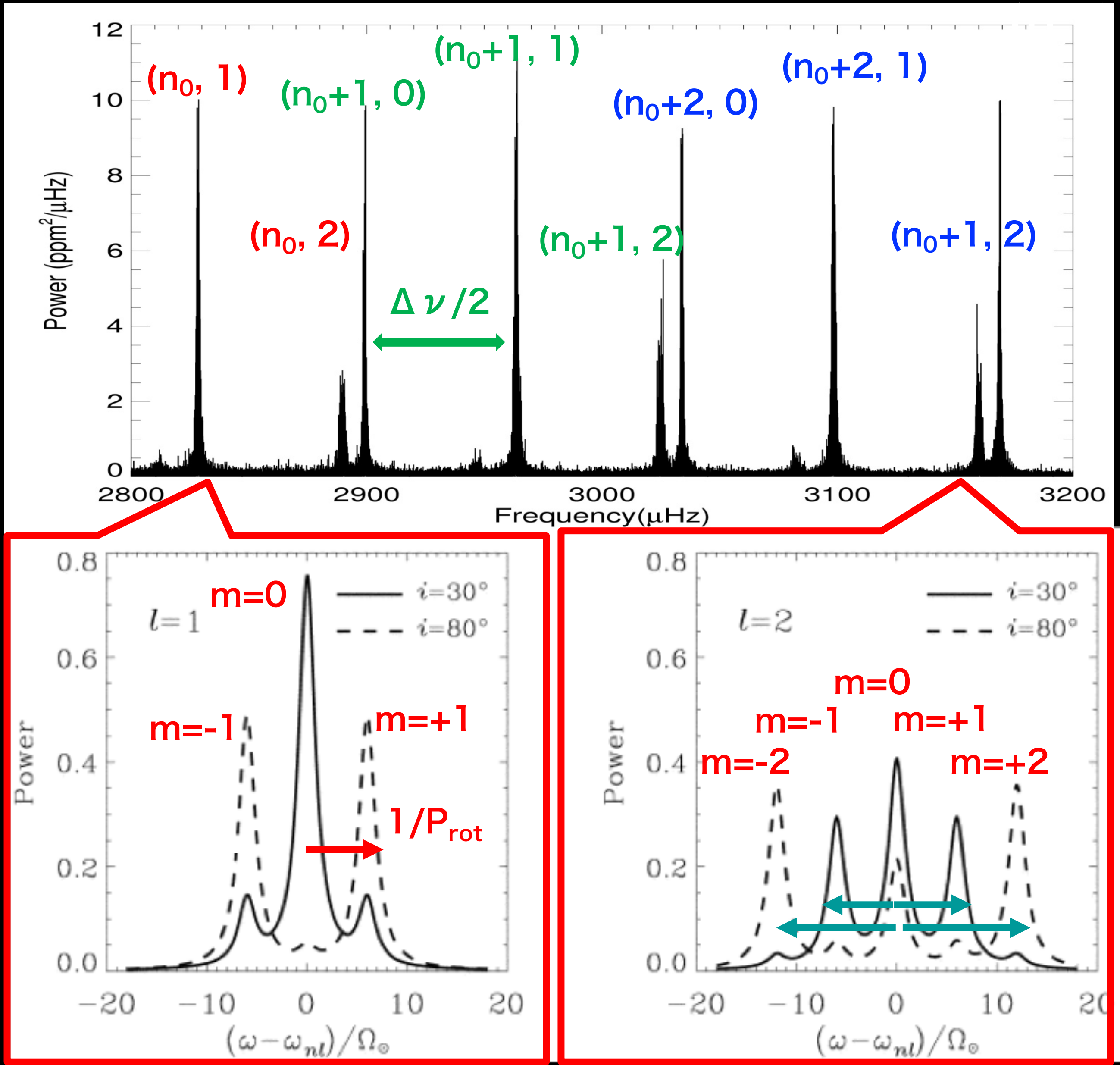
$$\nu_{nlm'} = \nu_{nl} + m\boxed{\delta\nu_*}(1-\boxed{C_{nl}})$$

stellar rotation small correction factor

Stellar rotation breaks the m-degeneracy

$$\nu(n,l) = \Delta\nu\left(n + \frac{l}{2} + \alpha\right) + \varepsilon_{n,l}$$

$$\nu(n,l,m) = \nu(n,l) + m\,\delta\nu_{\rm s}(n,l)$$



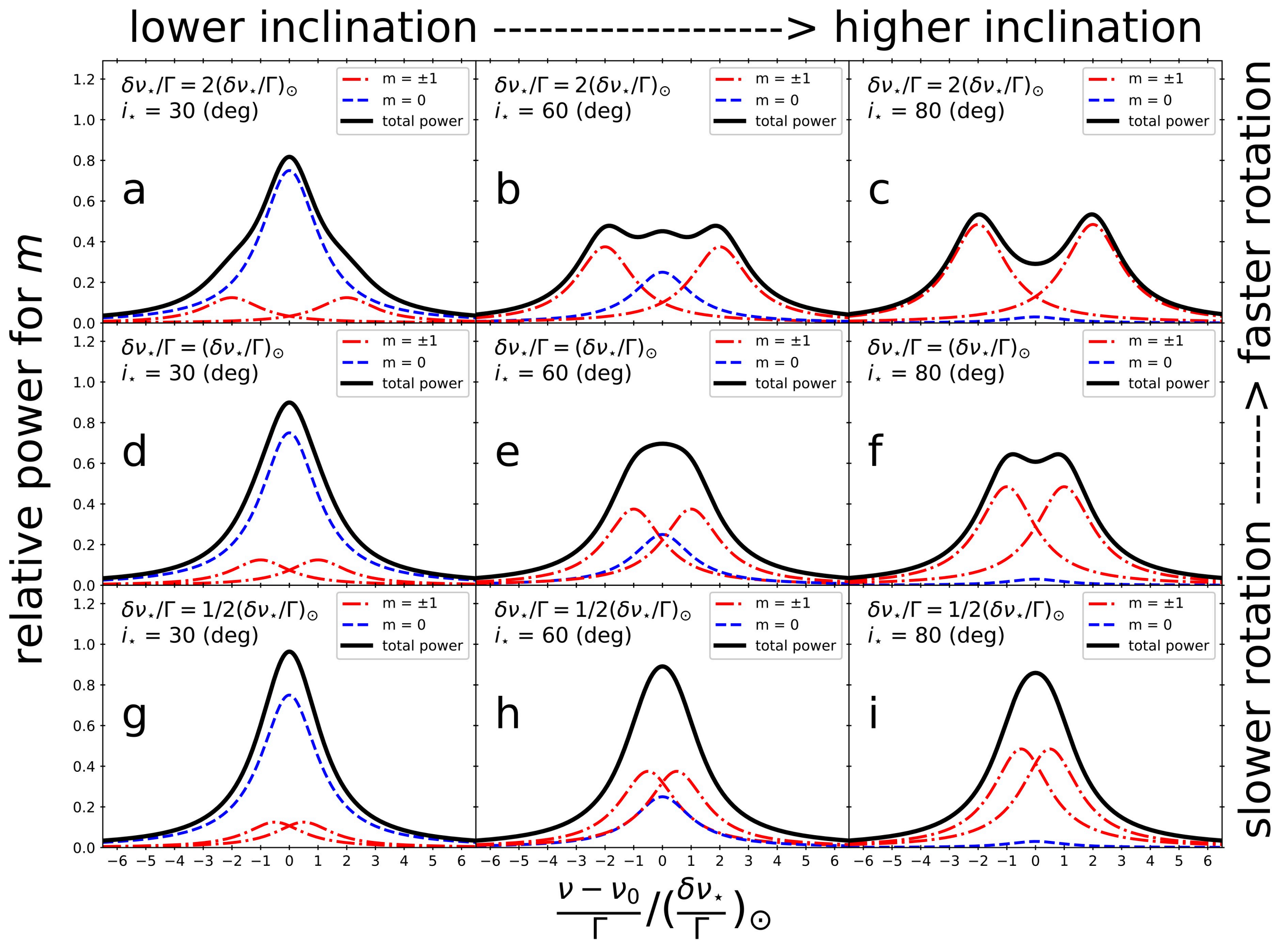
Stellar obliquity from asteroseismology

- Oscillation line mode profile: complementary probe of spin-orbit angles of exoplanetary systems

$\delta\nu_\star$: stellar rotation frequency

Γ: line width of the oscillation mode

Toutain & Gouttebroze, (1993)
Gizon & Solanki (2003)
Kamiaka, Benomar & Suto (2018)



c.f., Larmor's theorem vs. the Zeeman effect

- Lagrangian for a particle of mass m and charge q under scalar potential φ

$$L = \frac{1}{2}mv^2 - q\varphi(r, z)$$

- frame rotation around z-axis with frequency Ω

$$L = \frac{1}{2}mv^2 - q\varphi(r, z) + m\Omega(xv_y - yv_x) + \frac{1}{2}m\Omega^2 r^2$$

- Homogeneous magnetic field B along z-axis

$$\begin{aligned} L &= \frac{1}{2}mv^2 - q\varphi(r, z) + q\boldsymbol{v}\cdot\boldsymbol{A} \\ &= \frac{1}{2}mv^2 - q\varphi(r, z) + \frac{qB}{2}(xv_y - yv_z) \end{aligned}$$

- Frame rotation is equivalent to magnetic field (B=2m Ω/q)
  - B breaks the degeneracy of m-level (Zeeman effect)
  - Classical asteroseimology ⇔ quantum Zeeman effect

History of my personal prejudices on the spin-orbit architecture of planetary systems

Evolution of my own prejudice 1

Spin-orbit misalignment for exoplanets is unlikely

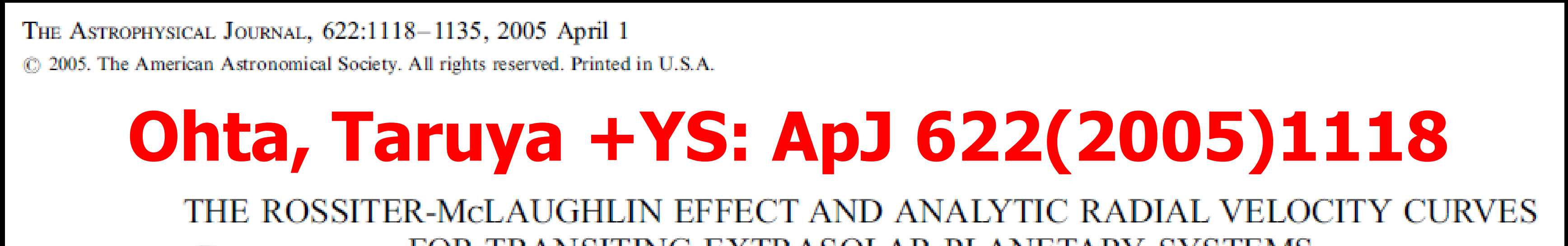
- **Queloz et al. (2000)**
  - First RM result for HD209458

$$\alpha = \pm 3.9^{\circ}{}^{+18^{\circ}}_{-21^{\circ}}$$

- **Ohta, Taruya + YS (2005)**
  - **Perturbative analytic formula for the RM effect**
  - spin-orbit angle should be small according the standard planet formation (Hayashi) model
  - If not, it indicates a new non-standard formation channel for exoplanets
- **Winn et al. (2005)**

$$\lambda = -4.4^{\circ} \pm 1.4^{\circ}$$

  - Significantly improved the RM measurement accuracy for HD209458 on the basis of OTS approach



The Astrophysical Journal, 622:1118–1135, 2005 April 1

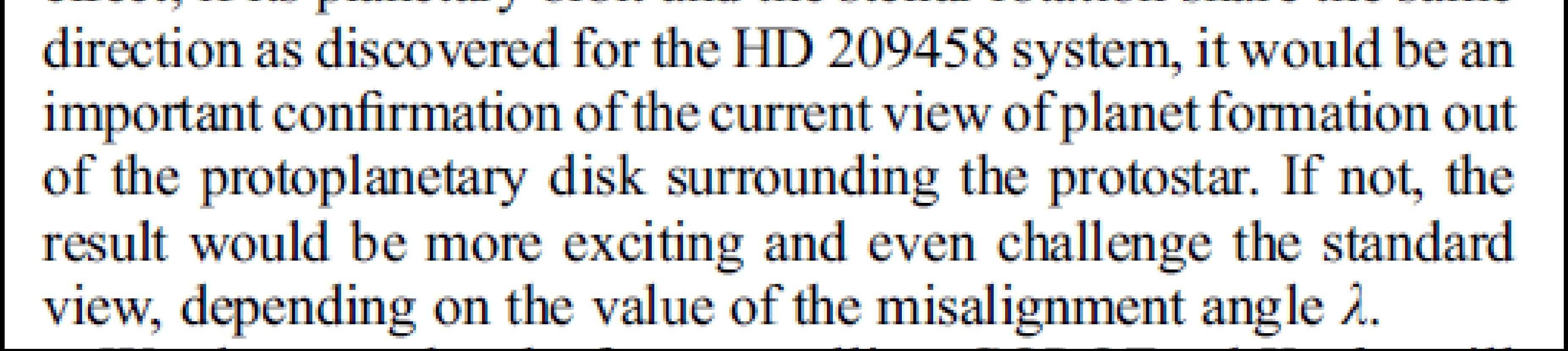
© 2005. The American Astronomical Society. All rights reserved. Printed in U.S.A.

Ohta, Taruya +YS: ApJ 622(2005)1118

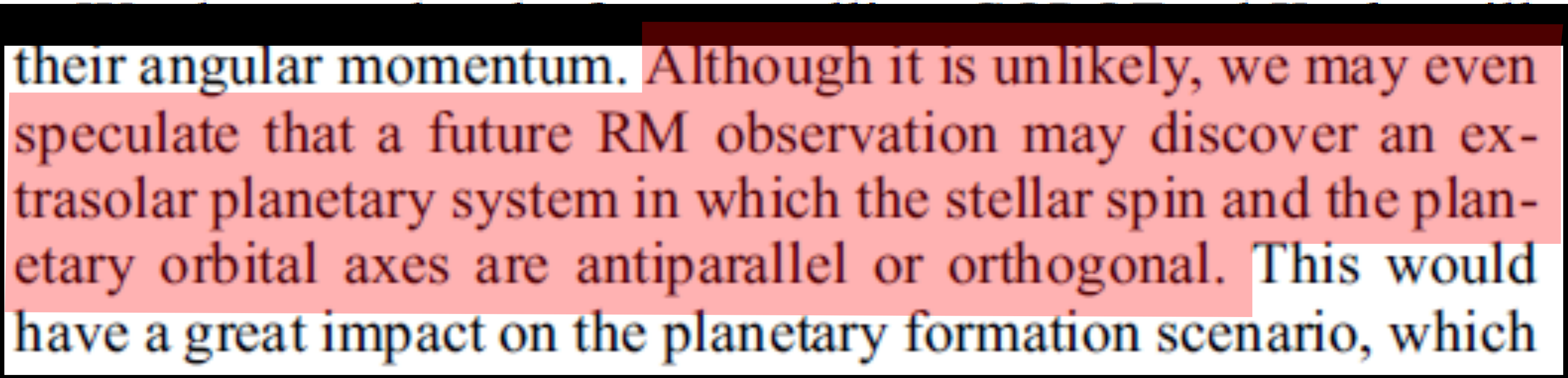
THE ROSSITER-McLAUGHLIN EFFECT AND ANALYTIC RADIAL VELOCITY CURVES FOR TRANSITING EXTRASOLAR PLANETARY SYSTEMS

Yasuhiro Ohta, Atsushi Taruya,[1] and Yasushi Suto[1]
Department of Physics, The University of Tokyo, Tokyo 113-0033, Japan; ohta@utap.phys.s.u-tokyo.ac.jp,
ataruya@utap.phys.s.u-tokyo.ac.jp, suto@phys.s.u-tokyo.ac.jp

Received 2004 October 13; accepted 2004 December 10



ect; ts planetary orbit and the stellar rotation share the same direction as discovered for the HD 209458 system, it would be an important confirmation of the current view of planet formation out of the protoplanetary disk surrounding the protostar. If not, the result would be more exciting and even challenge the standard view, depending on the value of the misalignment angle λ.

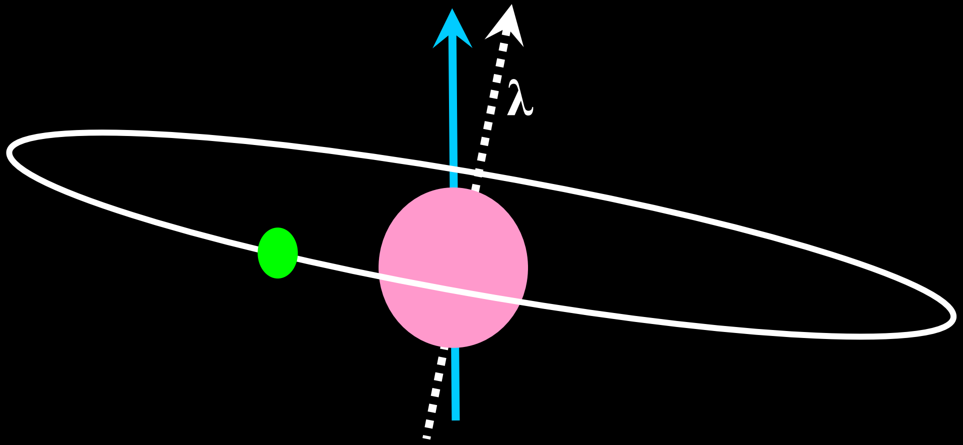


their angular momentum. Although it is unlikely, we may even speculate that a future RM observation may discover an extrasolar planetary system in which the stellar spin and the planetary orbital axes are antiparallel or orthogonal. This would have a great impact on the planetary formation scenario, which

Measurement of spin-orbit alignment in an extrasolar planetary system

- Joshua N. Winn, R.W. Noyes, M.J. Holman, D.B. Charbonneau, Y. Ohta, A. Taruya, Y. Suto, N. Narita, E.L. Turner, J.A. Johnson, G.W. Marcy, R.P. Butler, & S.S. Vogt
  - ApJ 631(2005)1215 (astro-ph/0504555)

$$\boxed{\lambda = -4.4^\circ \pm 1.4^\circ}$$



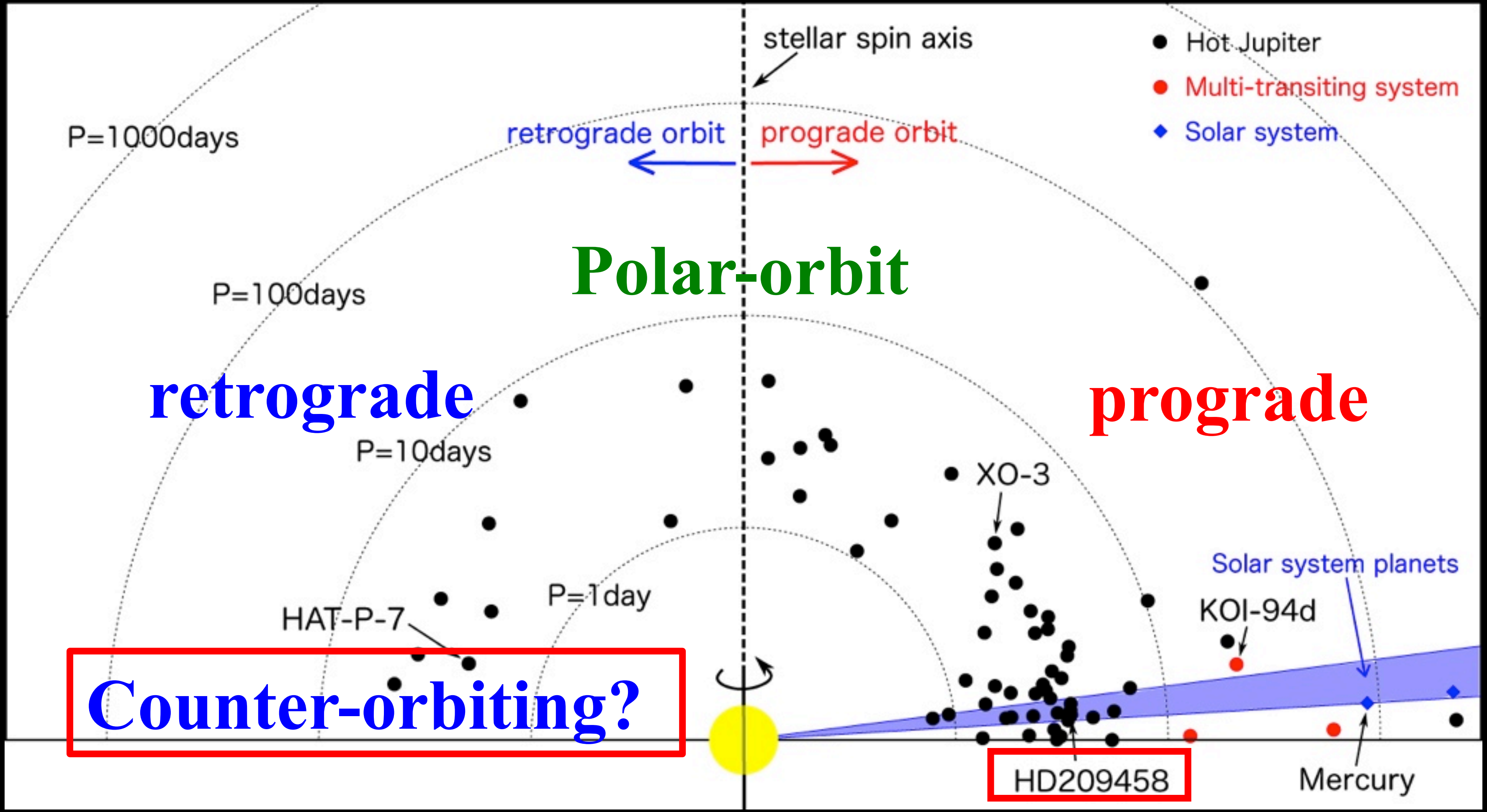
HD209458: Keck data + velocity anomaly template based on the perturbation formula by Ohta, Taruya & YS (2005)

Evolution of my own prejudice 2

Spin-orbit misalignment may be common for Hot Jupiters, but should not for transiting multi-planetary systems

- Around 2010, it became clear that a fair fraction of the observed Hot-Jupiters exhibits large spin-orbit misalignment
- This cannot happen, however, in transiting multi-planetary systems, which is unlikely to have suffered from significant dynamical disturbance, and thus should keep the initial condition (e.g., our Solar system aligned within several degrees)
- Let us test this prediction with the RM measurement for a transiting multi-planet system !

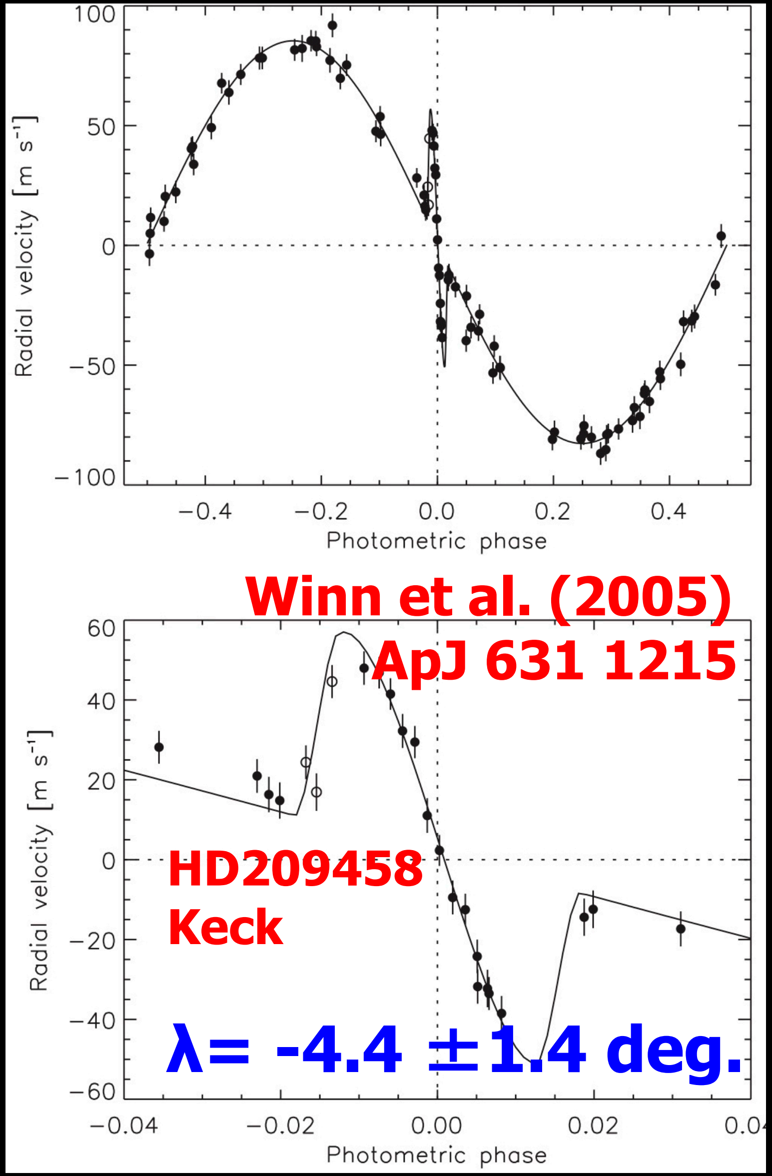
Projected spin-orbit angle distribution



As of June 2013, 29 out of 70 planets are known to have $\lambda > \pi/8$

Xue, Y.S., Tayura, Hirano, Fujii, and Masuda, ApJ 784(2014)66

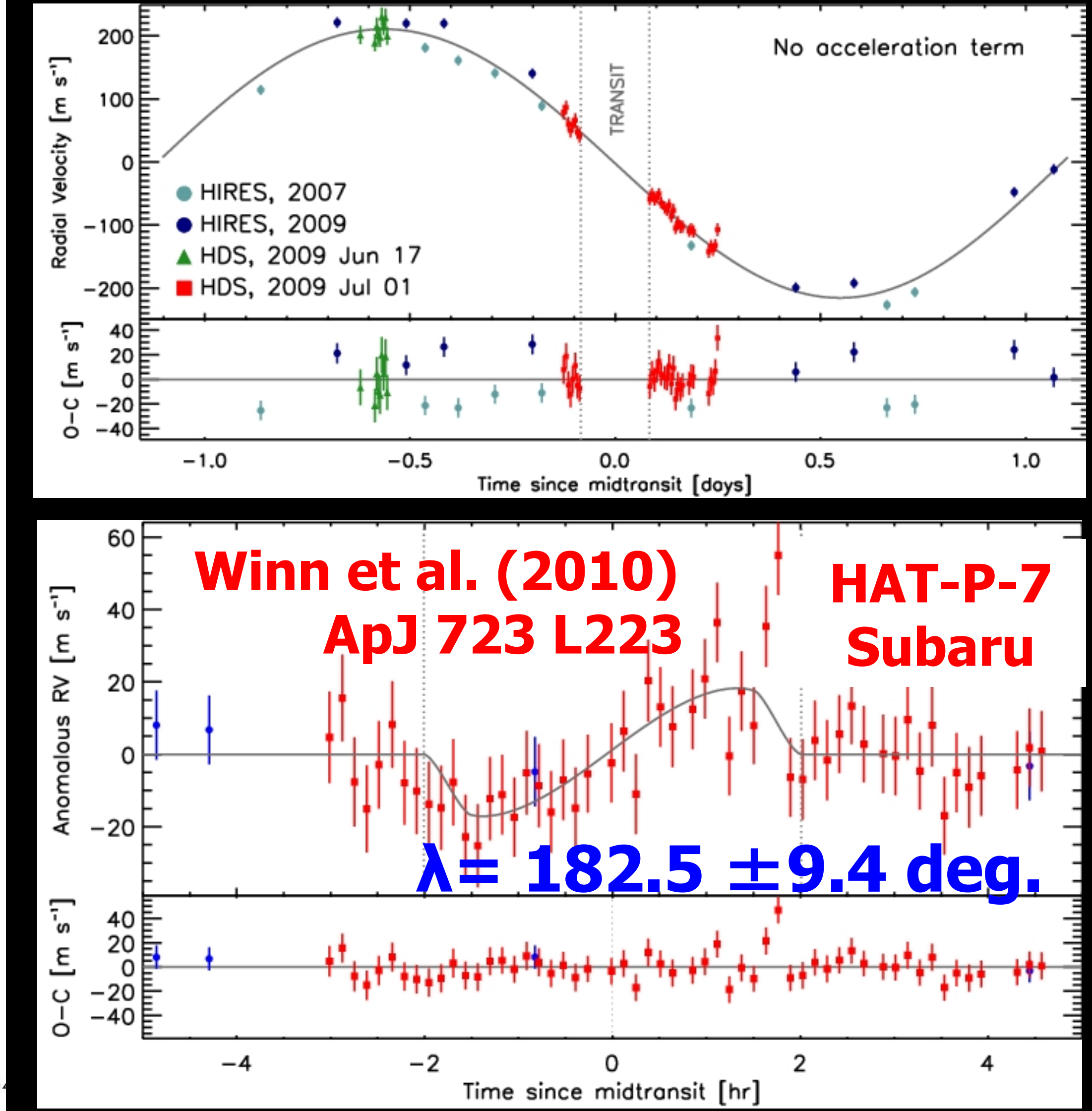
Prograde and retrograde orbits



Winn et al. (2005)
ApJ 631 1215

HD209458
Keck

λ= -4.4 ±1.4 deg.

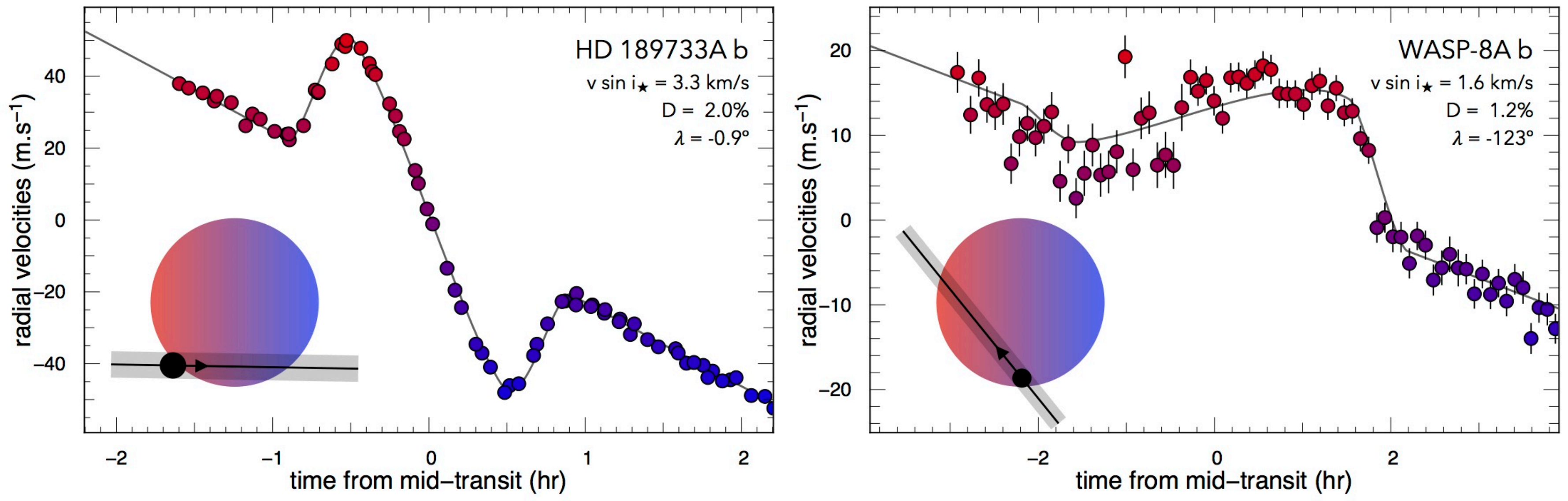


Winn et al. (2010)
ApJ 723 L223

HAT-P-7
Subaru

λ= 182.5 ±9.4 deg.

Examples of RM velocity anomaly



Aligned case

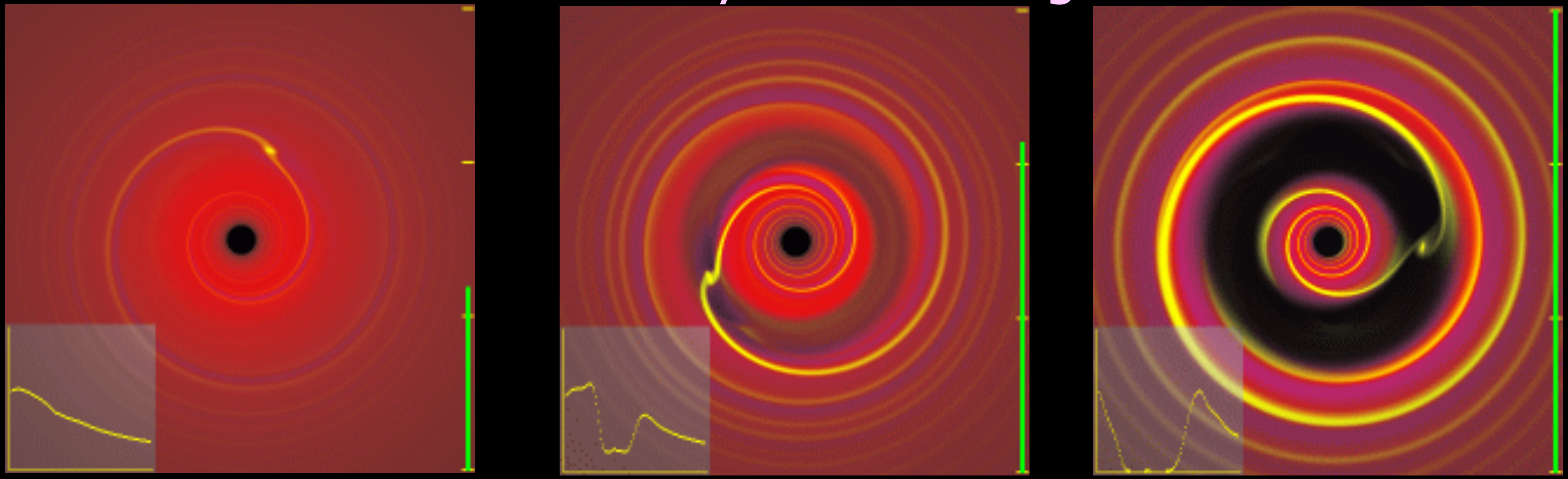
Misaligned case

Ohta, Taruya, & YS, ApJ 622(2005)1118
Winn et al. ApJ 631(2005)1215
Fabrycky & Winn, ApJ 696(2009)1230
Winn & Fabrycky, ARA&A 53(2015)409
Triaud arXiv:1709.06376

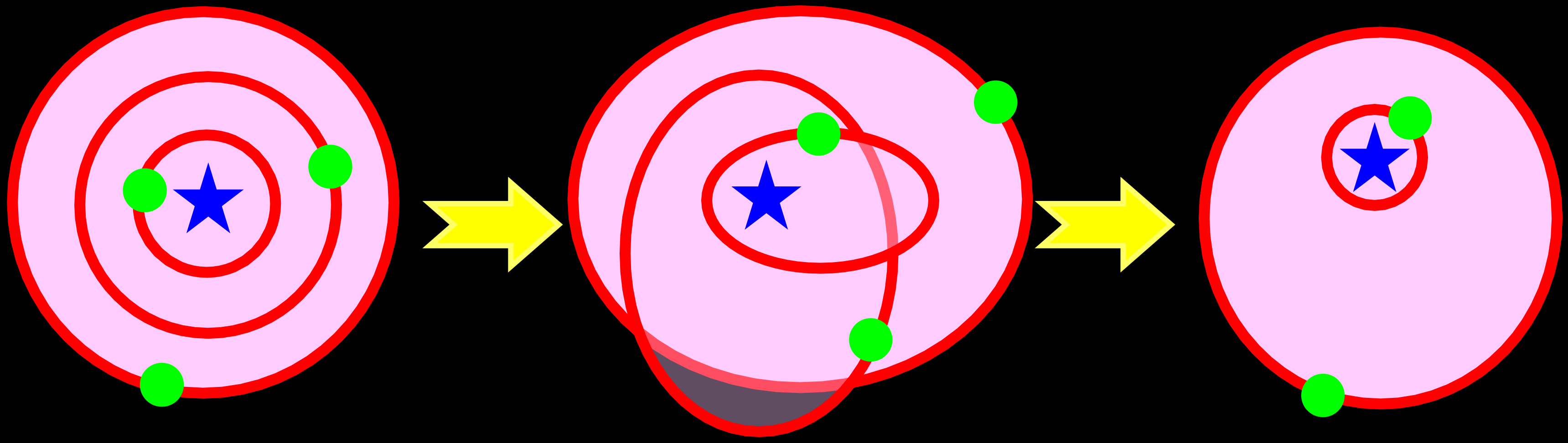
Planet migration channels

- Gravitational scattering
  - Planet - planet
- Type I migration
  - Low-mass planet - spiral wave in the gas disk
- Type II migration
  - High-mass planet - gap in the disk

Simulation by Phil Armitage

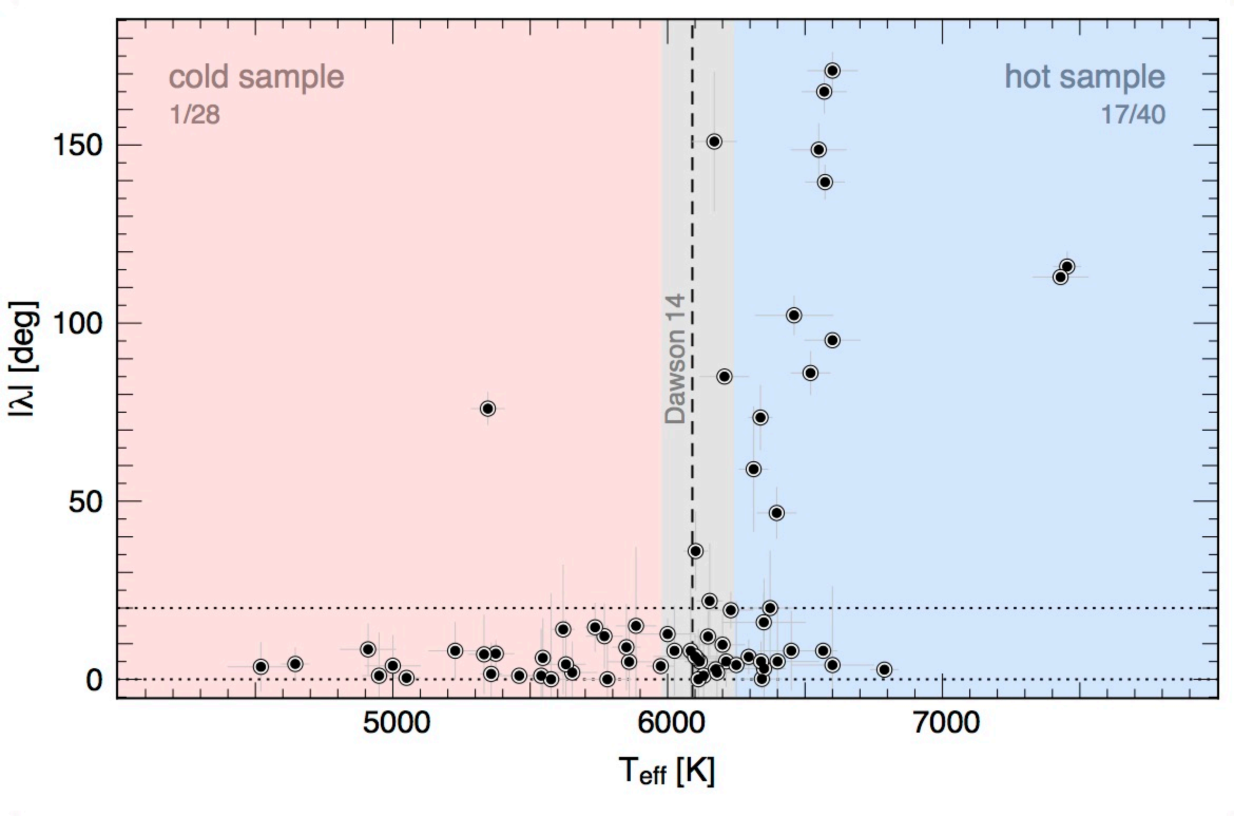


Planet-planet gravitation scattering + star-planet tidal interaction = circularized but misaligned Hot Jupiters



- Broad distribution of spin-orbit angles is generated due to planet scattering, tidal circularization, and the Lidov-Kozai effect (e.g., Nagasawa, Ida + Bessho 2008)

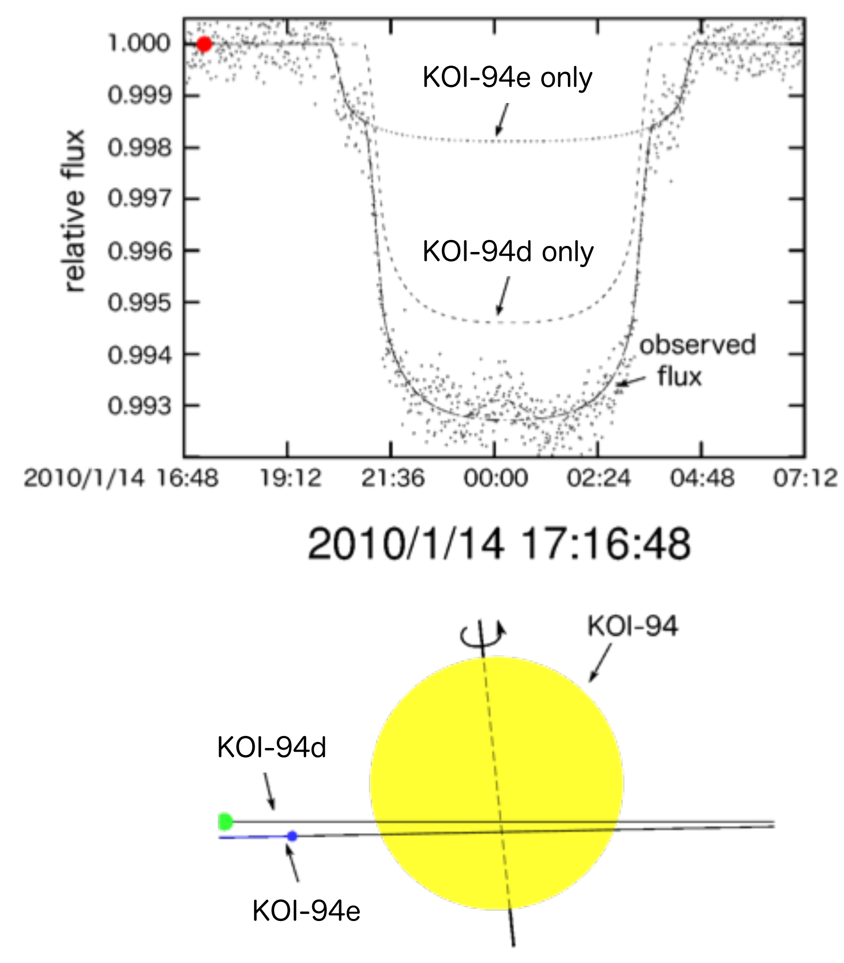
Projected misalignment vs. stellar effective temperature

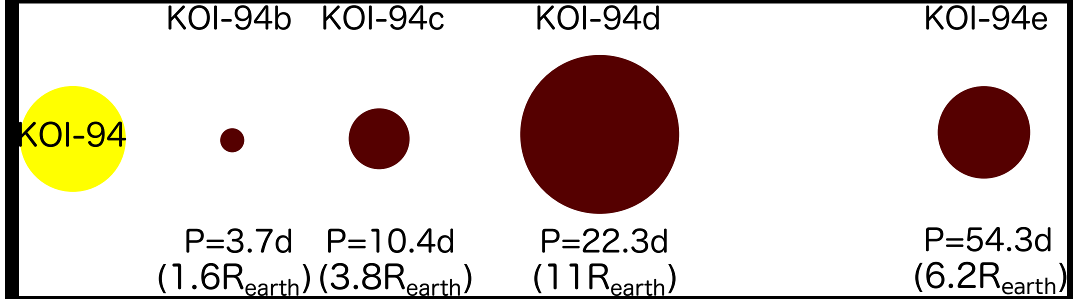


Triaud arXiv:1709.06376

More efficient spin-orbit "realignment" through star-planet tidal interaction due to the thicker convective zones of cool stars with T_{eff}<6100K ? (Winn et al. 2010)

RM observation of KOI-94 with Subaru: a system with 4 transiting planets





- First detection of planet-planet eclipse !
  - Even before we conduct the RM measurement in August 2012, we found an anomalous transit signature from Kepler archive on January 14, 2010
  - The orbital planes of those planets are well-aligned

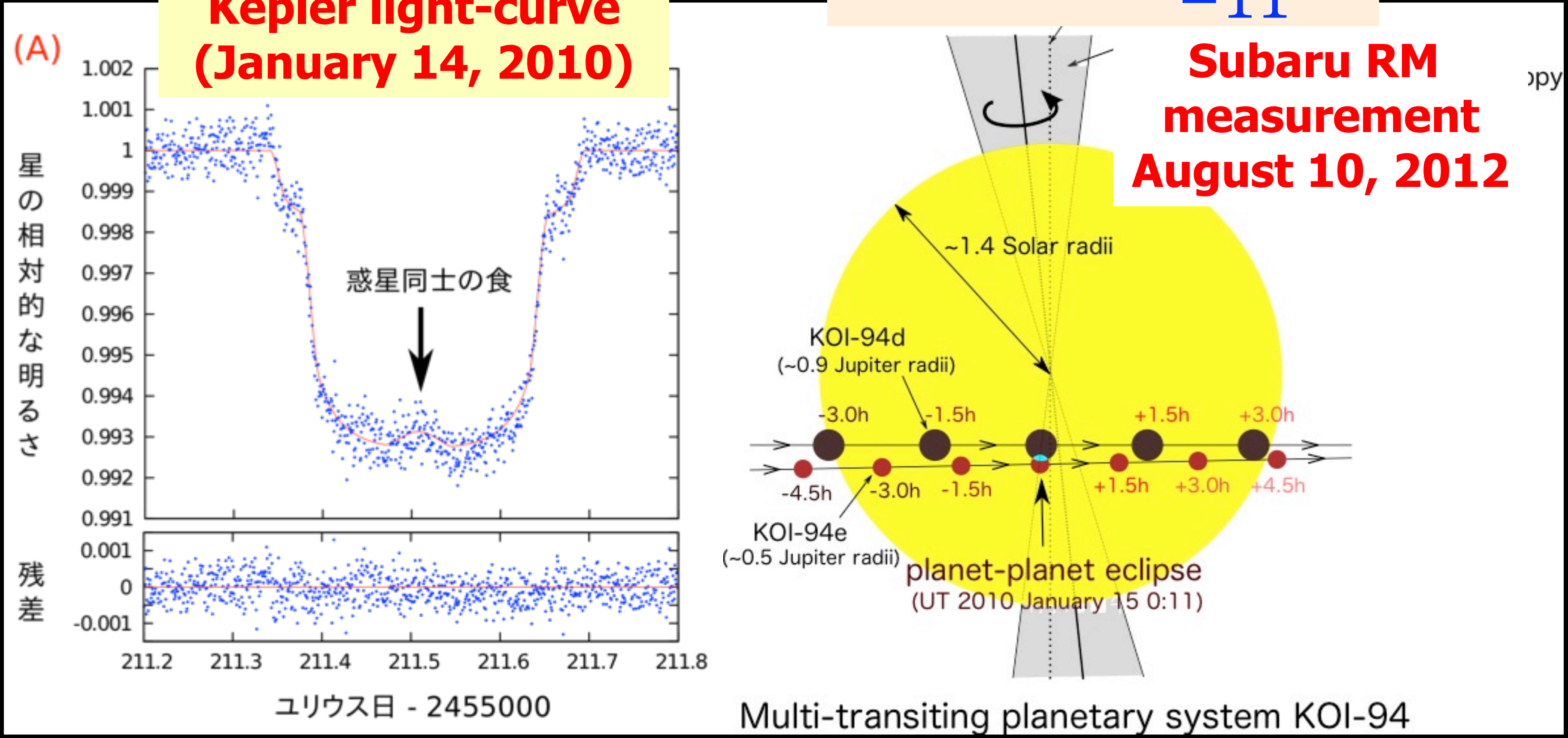
Hirano et al. ApJL 759 (2012)L36

Spin-orbit alignment of KOI-94

$$\lambda = -6^{\circ}{}^{+13^{\circ}}_{-11^{\circ}}$$

Kepler light-curve (January 14, 2010)

Subaru RM measurement August 10, 2012



Multi-transiting planetary system KOI-94

Hirano et al. ApJL 759 (2012) L36
Masuda et al. ApJ 778 (2013) 185

Evolution of my own prejudice 3

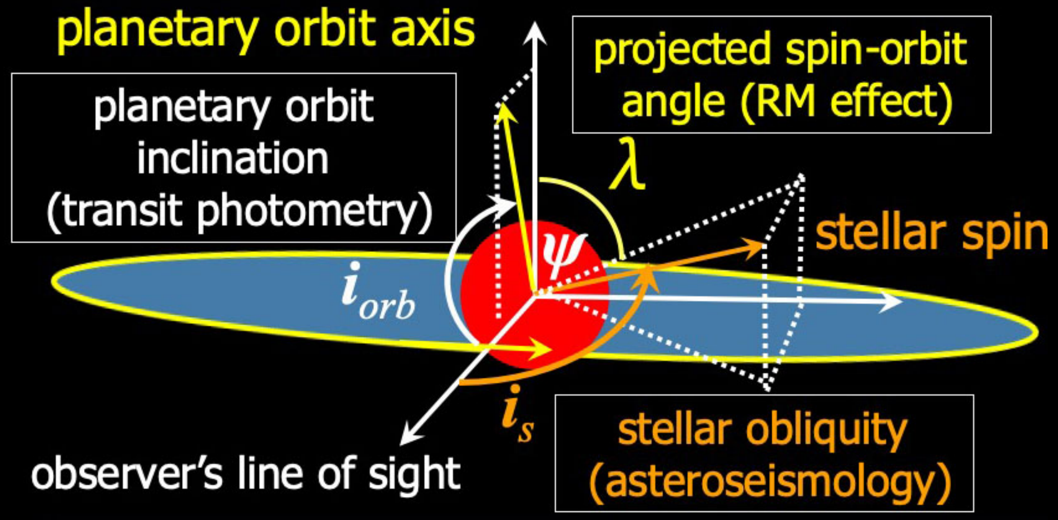
Spin-orbit misalignment should not exist for transiting multi-planetary systems

- Subaru spectroscopy + Kepler photometry of a transiting 4 planet system KOI 94 (Hirano et al. 2012, Masuda et al. 2013)
  - First measurement of RM effect for transiting multi-planet system
  - First discovery of planet-planet eclipse
  - KOI-94 was approved as Kepler-89
- Finally a reasonable picture established (?)

Evolution of my own prejudice 4

Stellar obliquity is another key

- Asteroseismology indicated the stellar obliquity of 47 ± 6 degree for Kepler-56
  - Kepler-56: red giant ($1.3M_s$, $4.3R_s$) + two transiting planets (10.5day, 20.4day)
  - Huber et al. Science 342(2013) 331
- RM effect measures the *projected spin-orbit angle*
  - Is this also the case for other multi-planet systems, especially with a main-sequence host star ?



planetary orbit axis

projected spin-orbit angle (RM effect)

planetary orbit inclination (transit photometry)

λ

stellar spin

ψ

i_{orb}

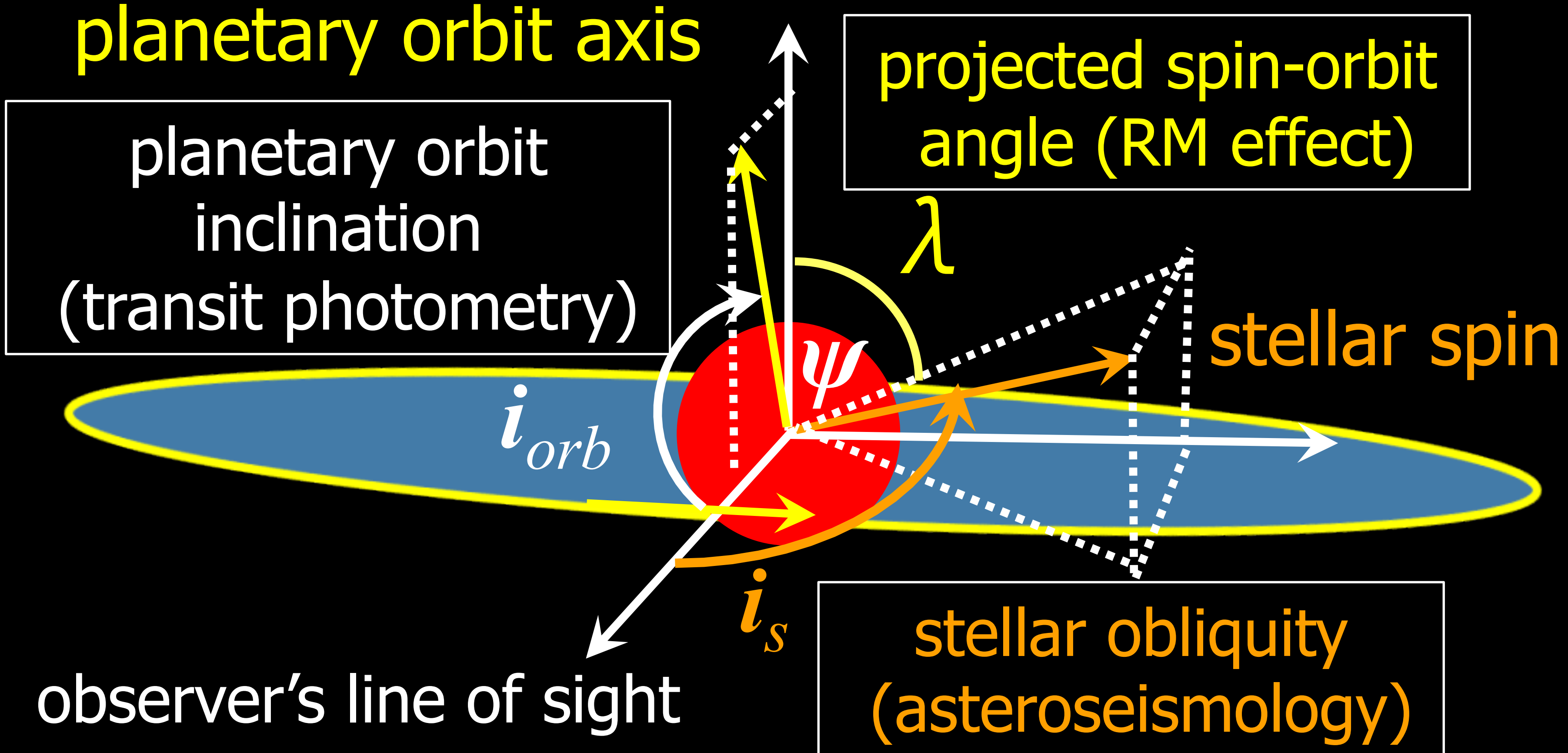
i_s

observer's line of sight

stellar obliquity (asteroseismology)

$$\cos\Psi = \sin i_s \sin i_{orb} \cos\lambda + \cos i_s \cos i_{orb}$$

Spin-orbit angles of a transiting planet



planetary orbit inclination (transit photometry)

projected spin-orbit angle (RM effect)

stellar obliquity (asteroseismology)

observer's line of sight

$$\cos\Psi = \sin i_s \sin i_{orb} \cos\lambda + \cos i_s \cos i_{orb}$$
$$\approx \sin i_s \cos\lambda$$

True spin-orbit angles from RM effect + asteroseismology

- Only two systems have both measurements of λ (RM) and i_s (asteroseismology)
  - Kepler-25 (F-star+ planets with 6 and 13days)

$$\lambda = 9.4^\circ \pm 7.1^\circ \quad i_s = 65.4^\circ{}^{+12.1^\circ}_{-7.4^\circ} \quad \Psi = 26.9^\circ{}^{+7.0^\circ}_{-9.2^\circ}$$

  - see Campante et al. (2016) $i_s = 80.6^\circ{}^{+6.5^\circ}_{-9.3^\circ}$ $\Psi = 12.6^\circ{}^{+6.7^\circ}_{-11.0^\circ}$
  - HAT-P-7 (F-star + a single planet with 2.2 days)

$$\lambda = 186^\circ{}^{+10^\circ}_{-11^\circ} \quad i_s = 27^\circ{}^{+35^\circ}_{-18^\circ} \quad \Psi = 122^\circ{}^{+30^\circ}_{-18^\circ}$$

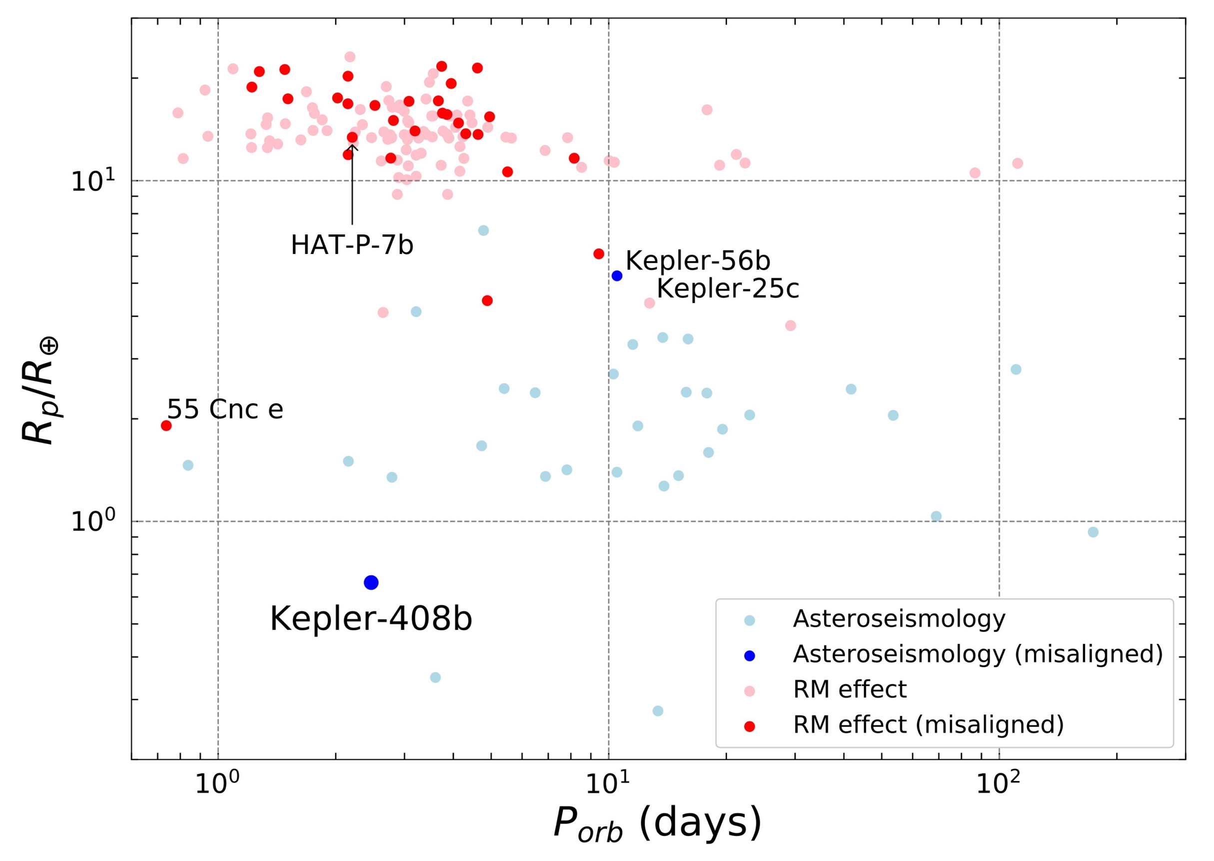
Not a counter-orbiting planet

Benomar, Masuda, Shibahashi + YS, PASJ 66(2014) 9421

see also Huber et al. (2013) , Campante et al.(2016)

Evolution of my own prejudice 5

asteroseismology is really reliable ?



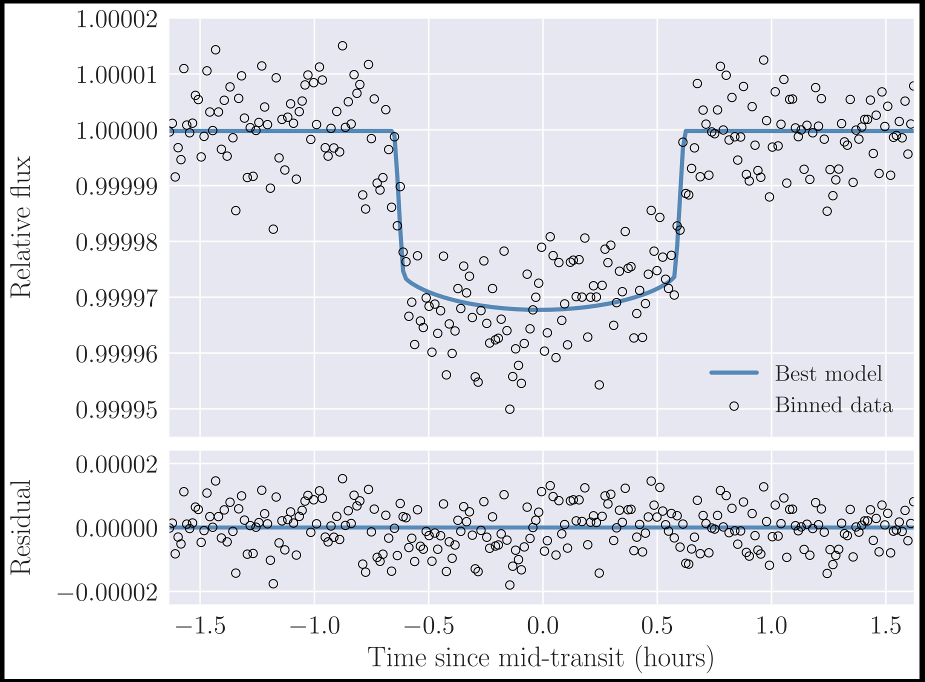
- RM effect
  - short-period and large planets
- Asteroseismology
  - independent of the properties of planets

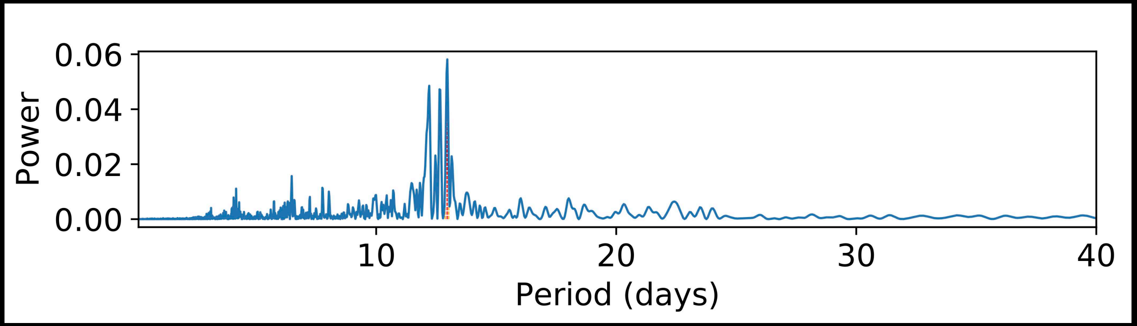
Kamiaka, Benomar & YS
MNRAS(2018)
Kamiaka, Benomar, YS, Dai,
Masuda, & Winn
AJ 157(2019)137
Suto, Kamiaka & Benomar
AJ 157(2019)172

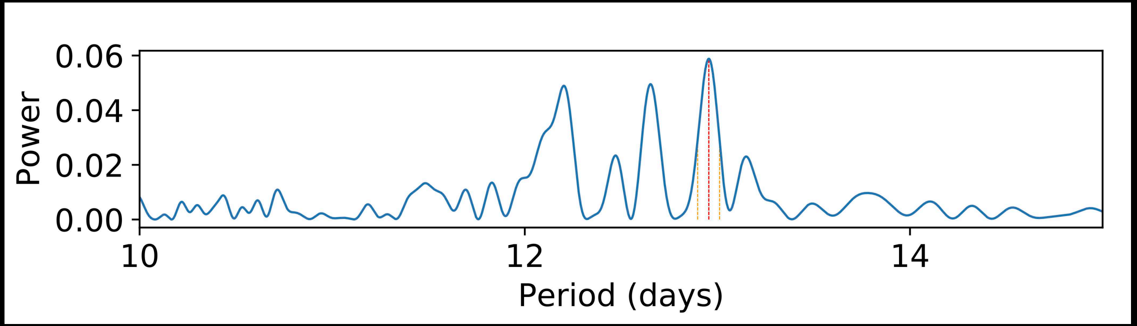
- Asteroseismology is based on various (non-trivial) assumptions, and required complicated and careful modeling

Transiting planetary system Kepler-408

- Kepler-408
  - Star: 6100K, 1.05M$_{sun}$, 1.25R$_{sun}$
  - Planet: sub-Earth size 0.86R$_E$, 2.5day orbital period



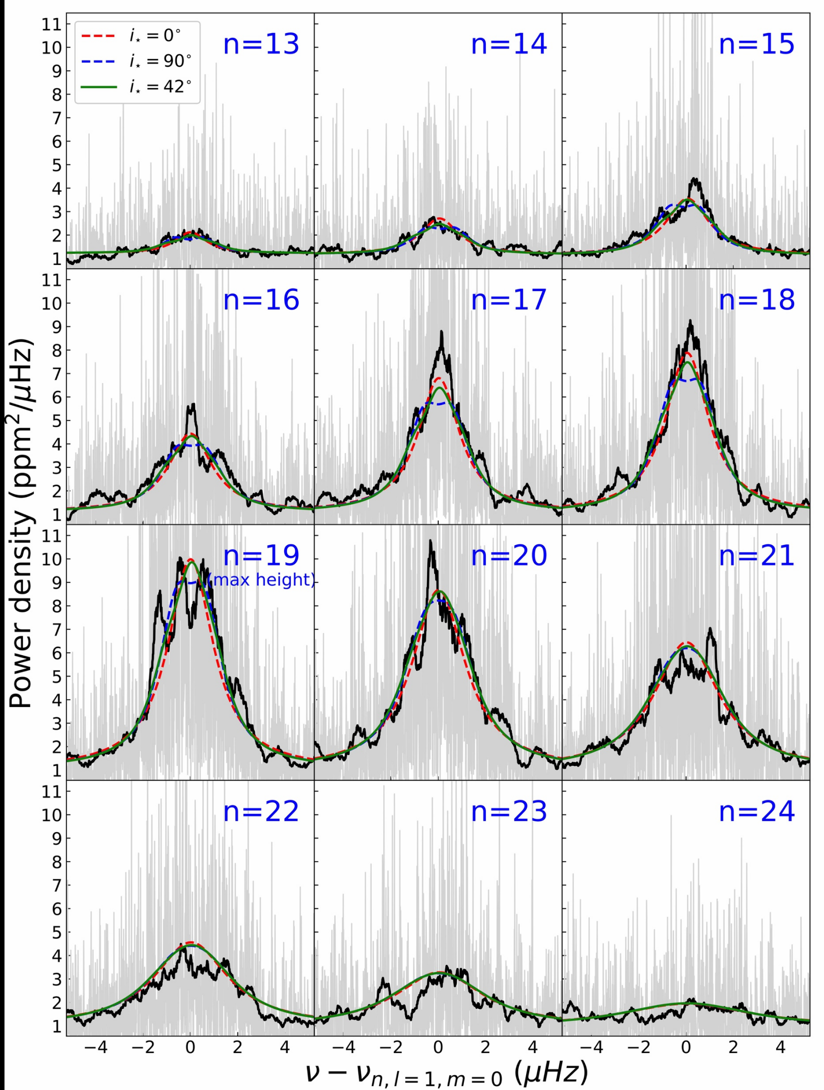


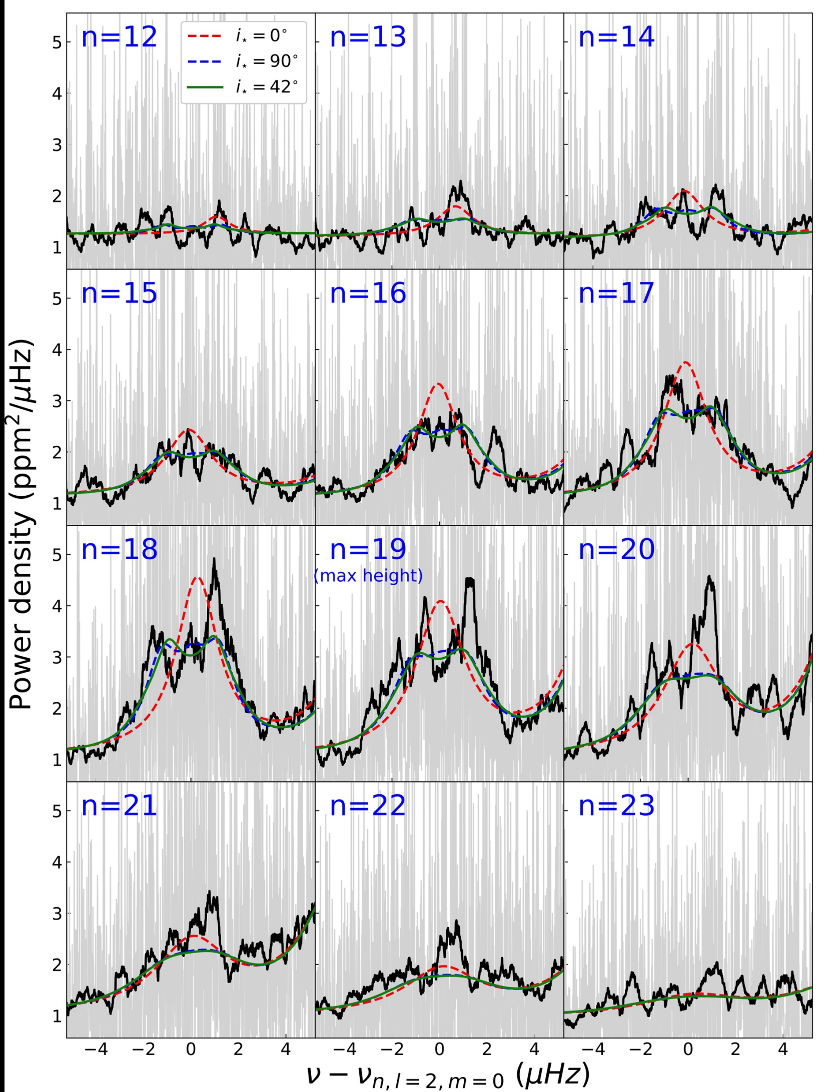


Kamiaka, Benomar, YS, Dai, Masuda & Winn, AJ 157(2019)137

Oscillation profiles (n,l) of Kepler-408

l=1

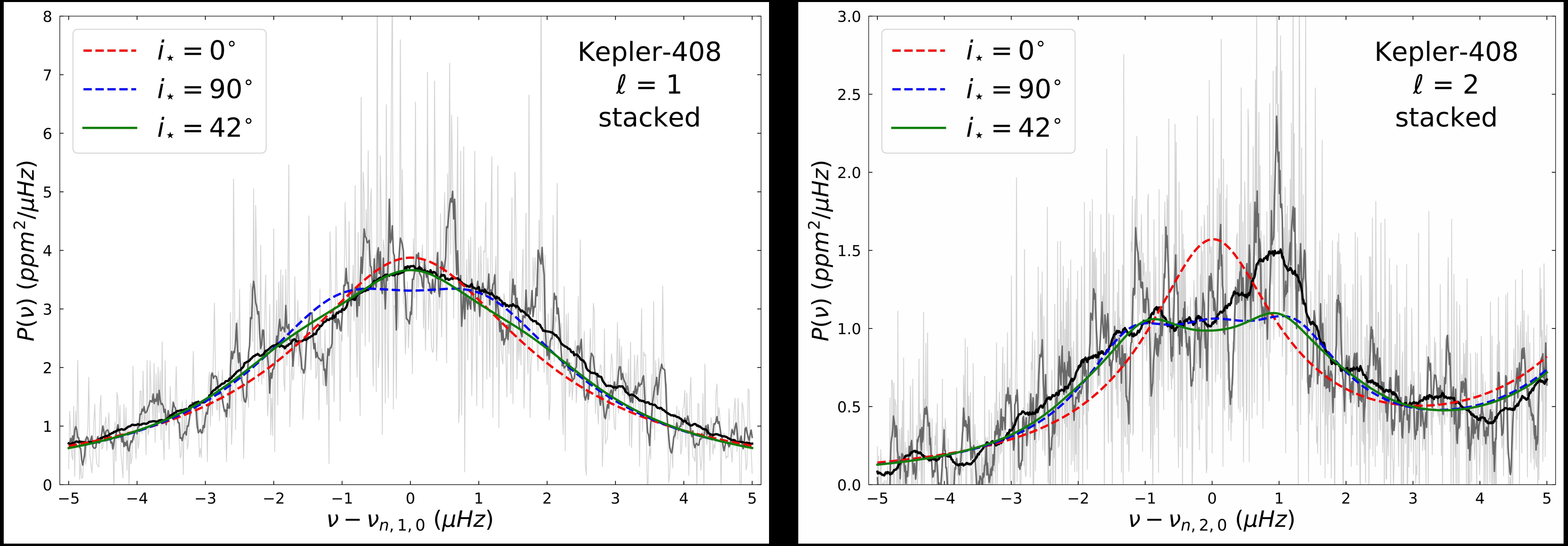




l=2

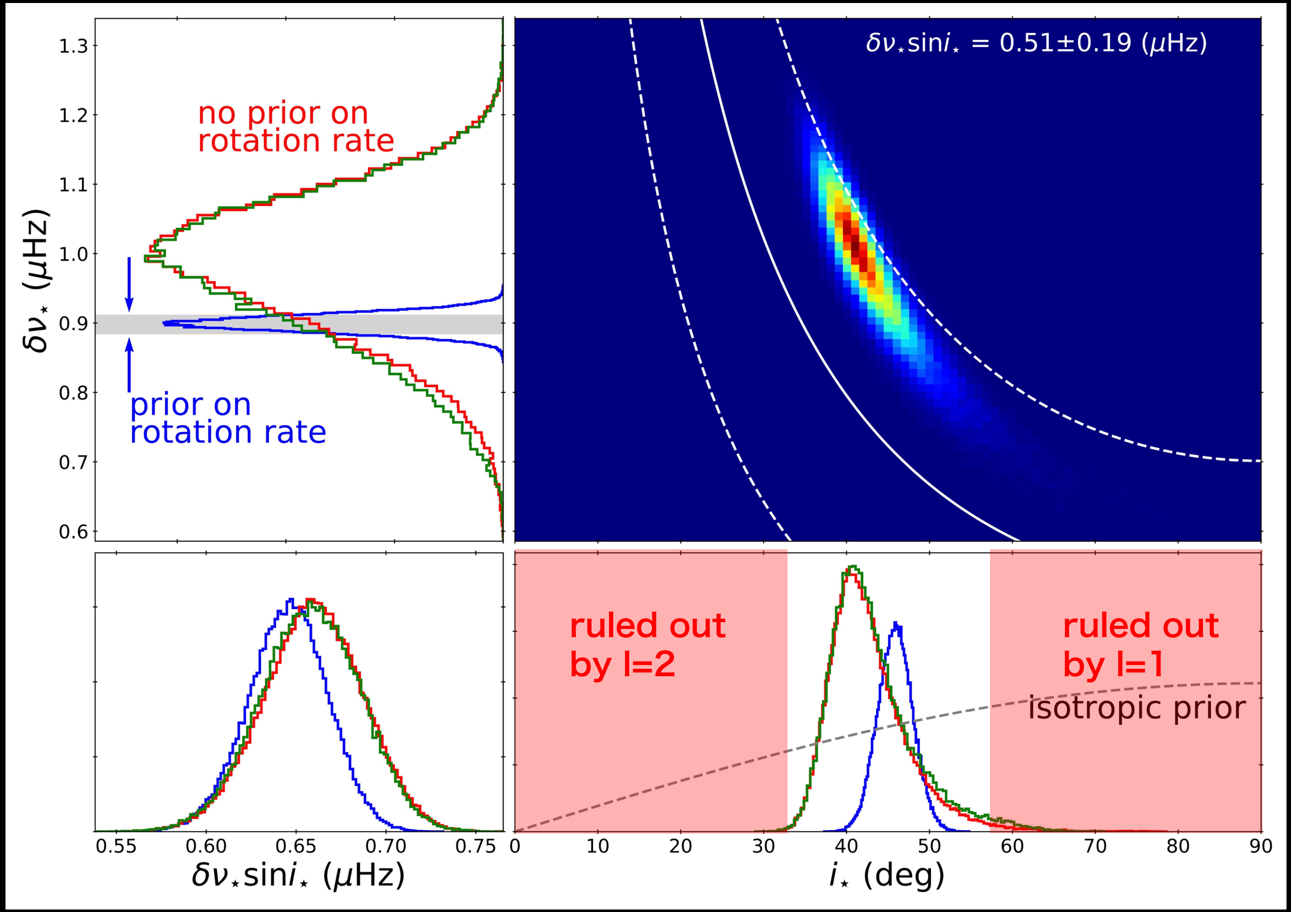
Kamiaka, Benomar, YS, Dai, Masuda, & Winn, AJ 157(2019)137

Stacked oscillation spectra of Kepler-408



Kamiaka, Benomar, YS, Dai, Masuda, & Winn, AJ 157(2019)137

Asteroseismic constraints on Kepler-408



Kamiaka, Benomar, YS, Dai, Masuda, & Winn
AJ 157(2019)137

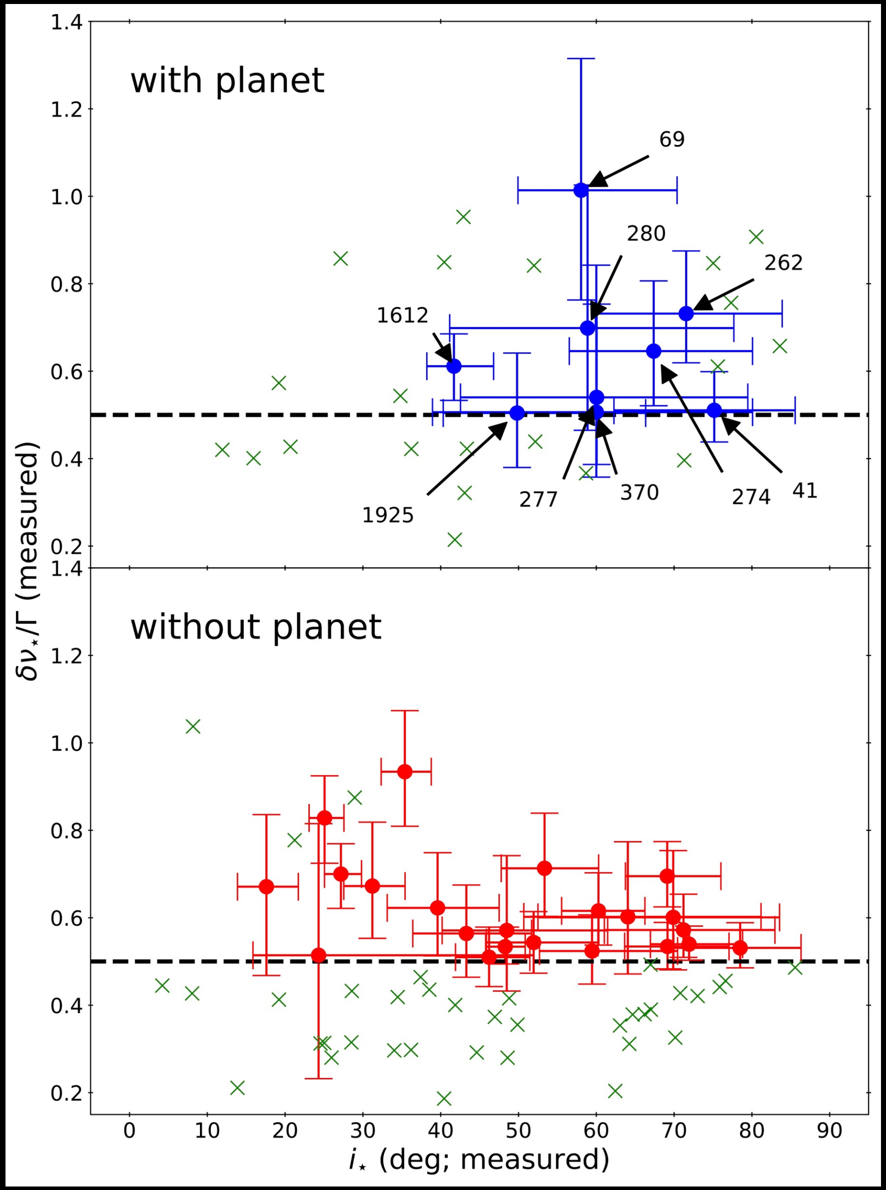
- Consistent with the other estimate
  - Photometric rotation period : P_{rot}
  - Doppler line broadening : $v_{rot} \sin i_\star$
- The smallest size planet in an oblique orbit

$$i_\star = \sin^{-1}\left(\frac{v_{\rm rot}\sin i_\star}{2\pi R_\star/P_{\rm rot}}\right) = 44^{+20}_{-15}\,(\mathrm{deg})$$

i_s of Kepler stars from asteroseismology: with/without planets

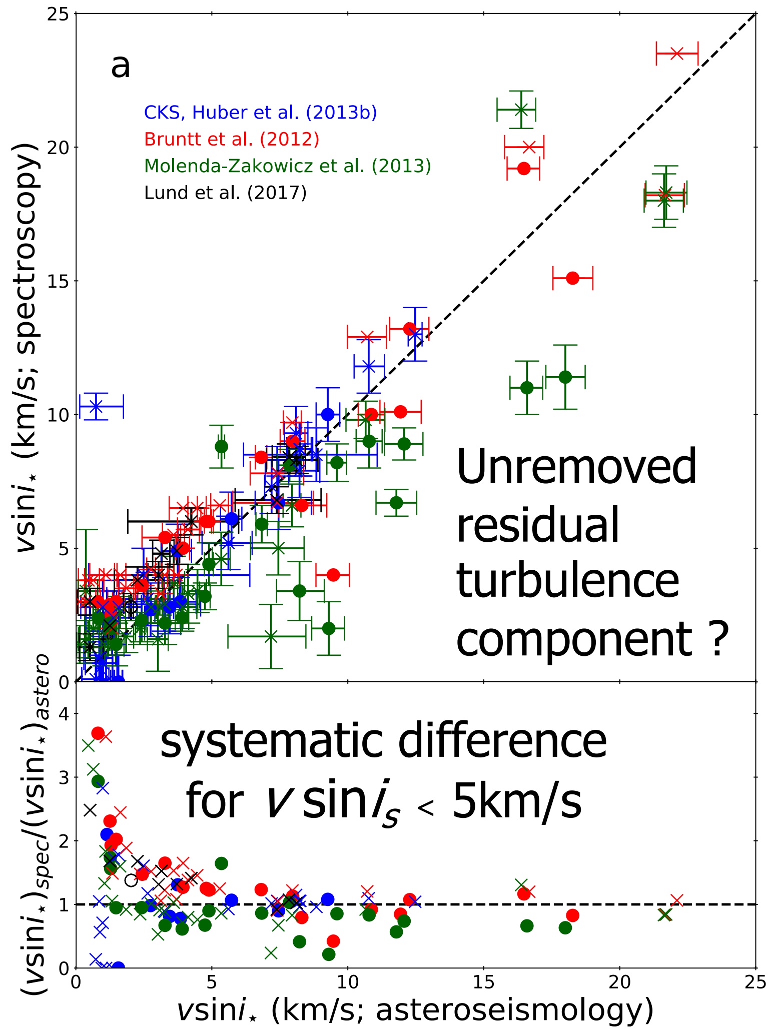
- 94 Kepler main-sequence stars
  - 33 with transiting planets
  - 61 with no known planets
- Transiting planet-host stars have systematically larger stellar obliquities (as expected)

Kamiaka, Benormar, and YS (2018)



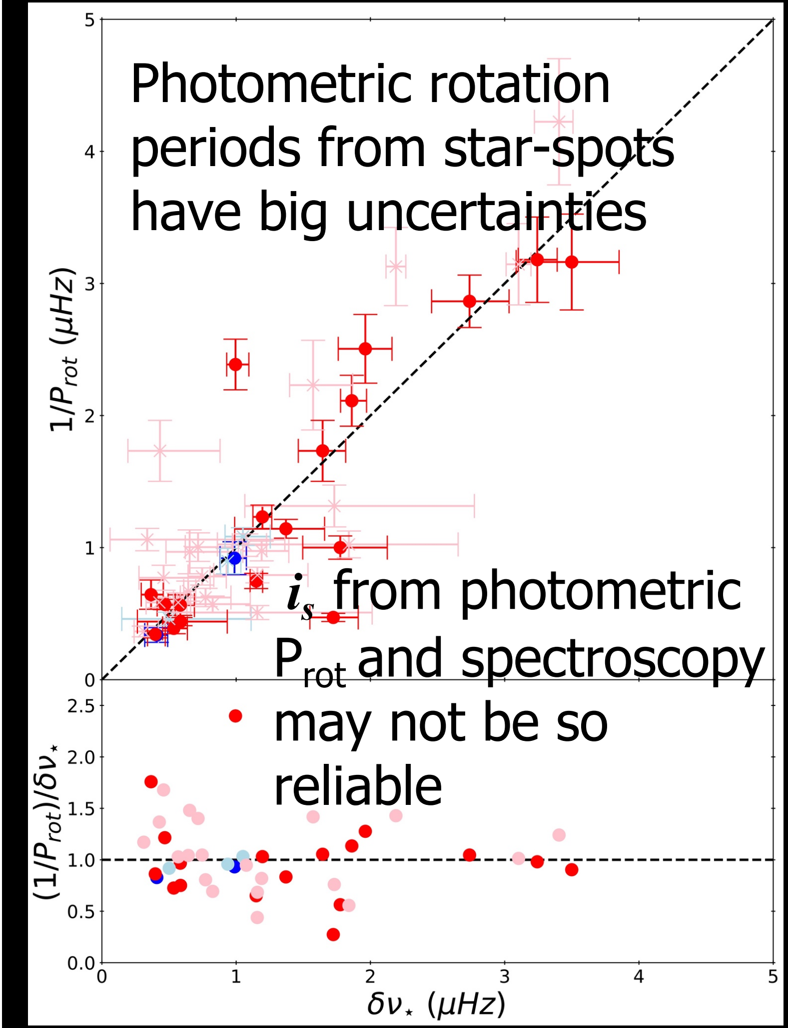
Comparison with independent observational estimates of $v\sin i_s$, P_{rot} and i_s

Kamiaka, Benormar, and YS (2018)



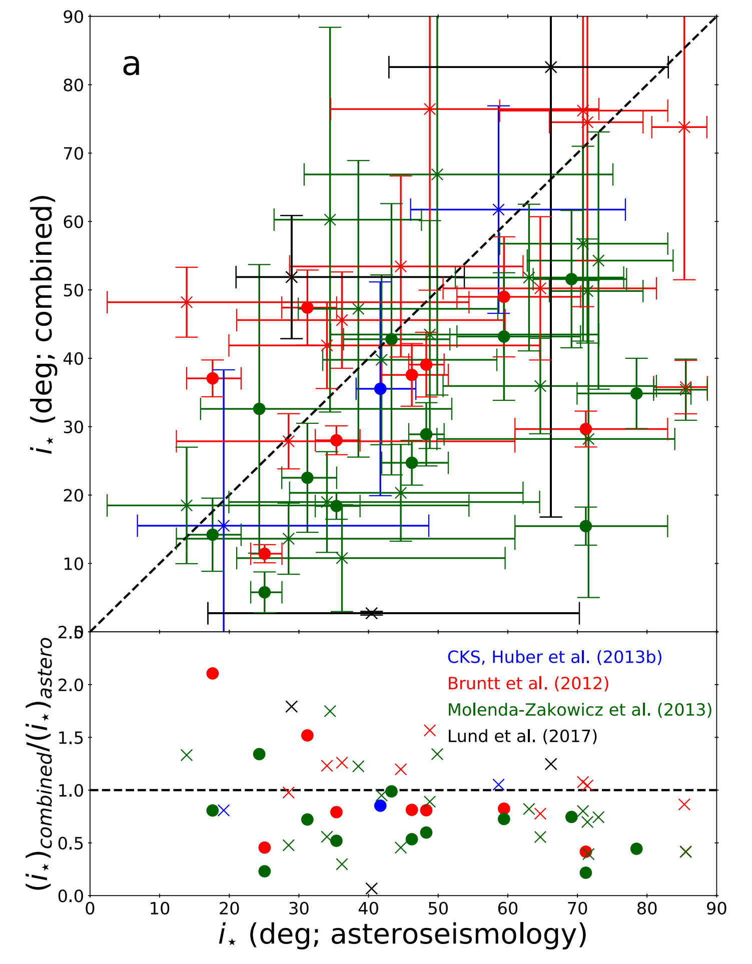
Unremoved residual turbulence component ?

systematic difference for $v\sin i_s$ < 5km/s

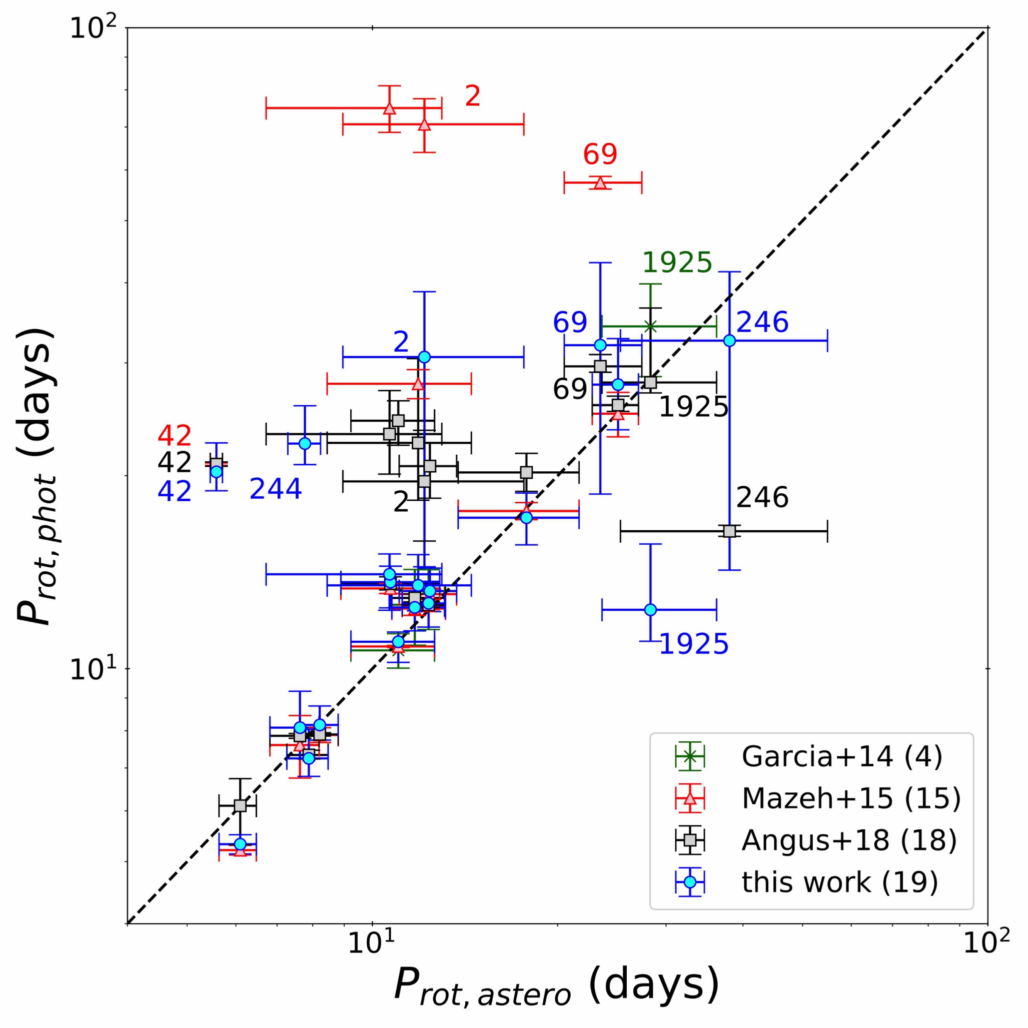


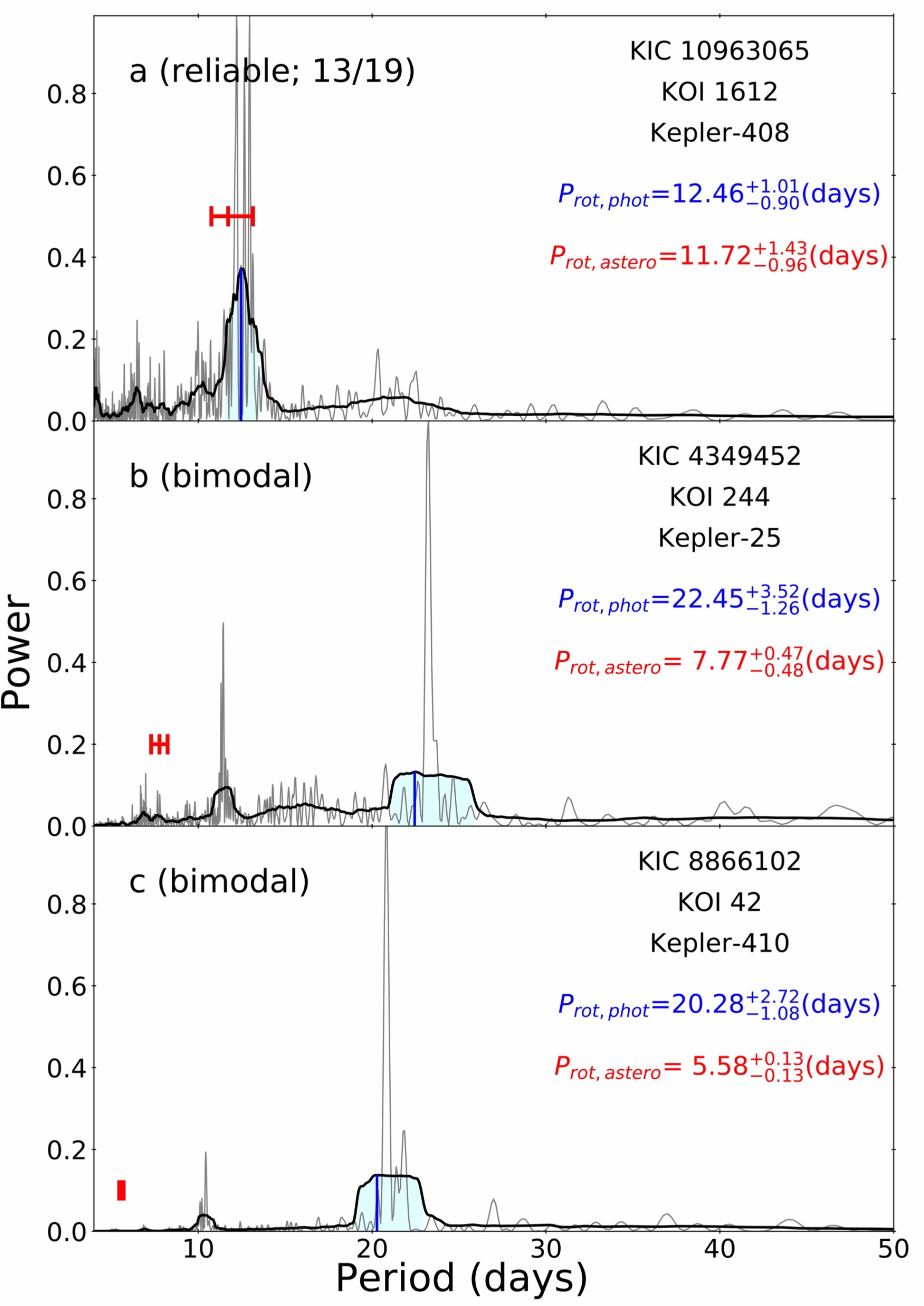
Photometric rotation periods from star-spots have big uncertainties

i_s from photometric P_{rot} and spectroscopy may not be so reliable



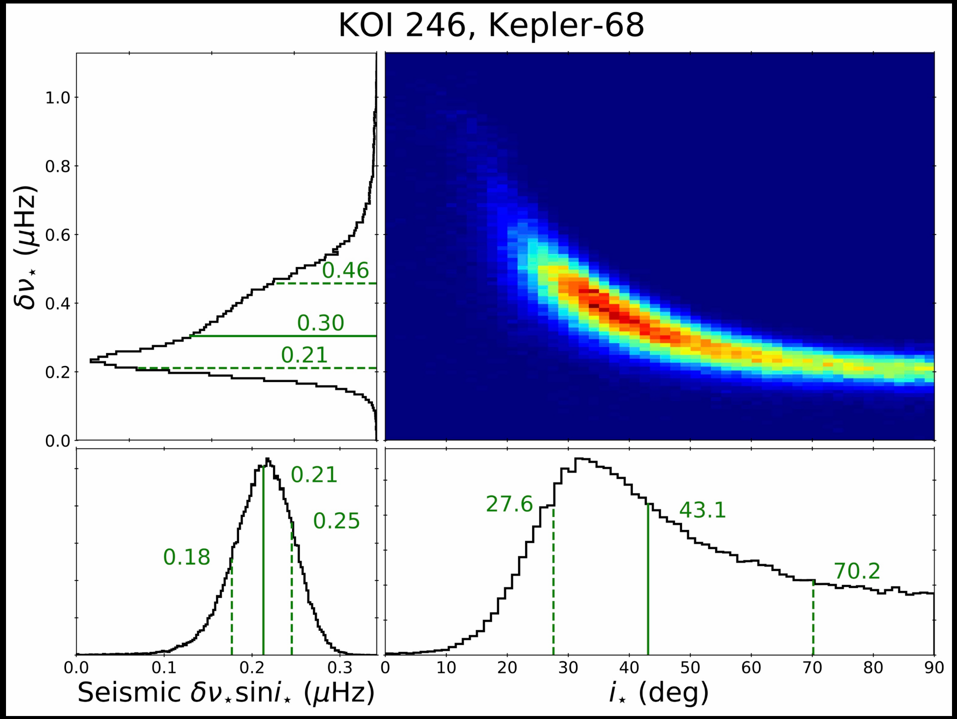
Photometric variation vs. asteroseismology

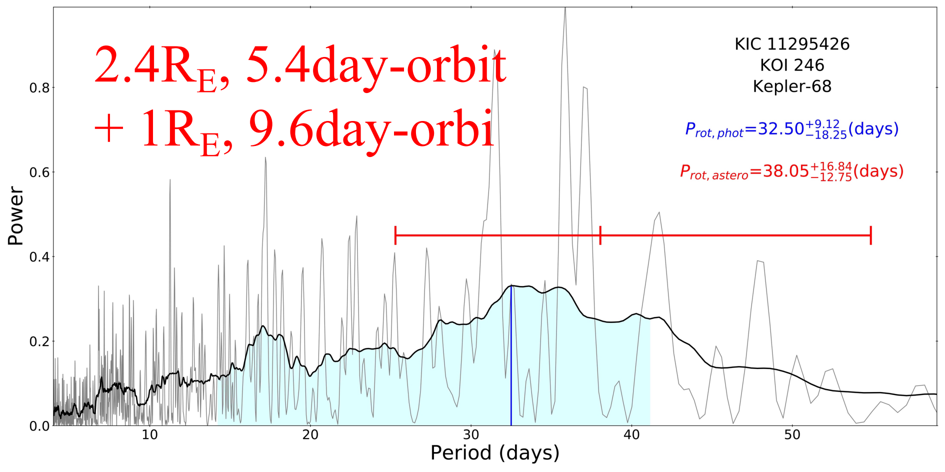


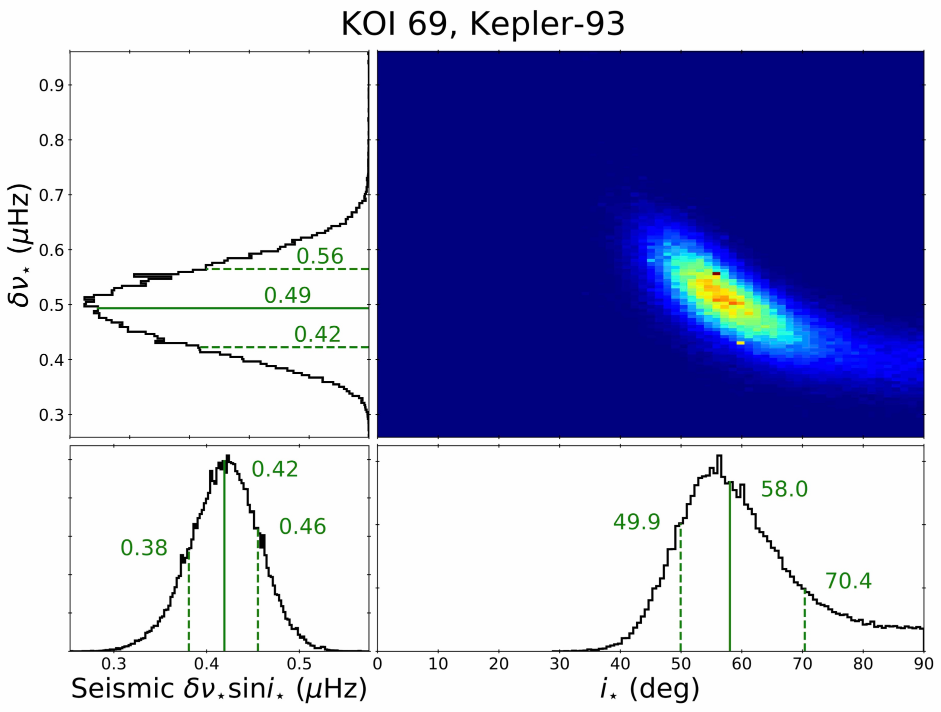


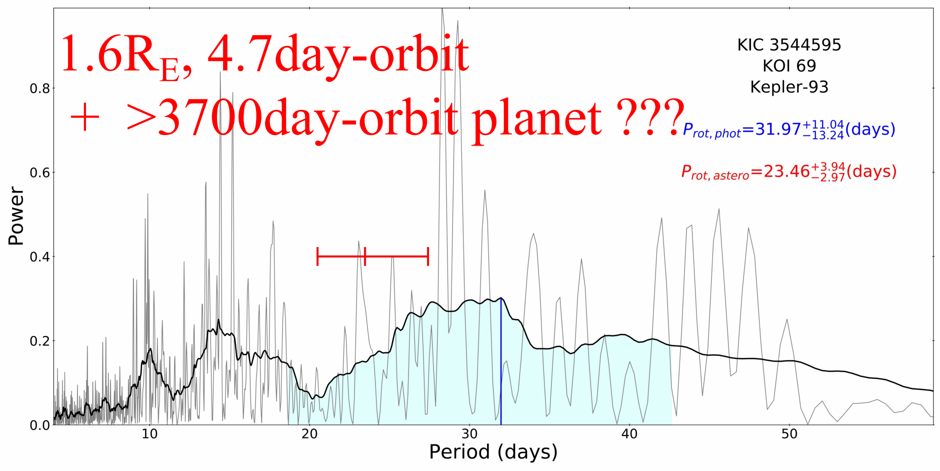
YS, Kamiaka & Benomar AJ 157(2019)172

multi-planetary systems of possible interest

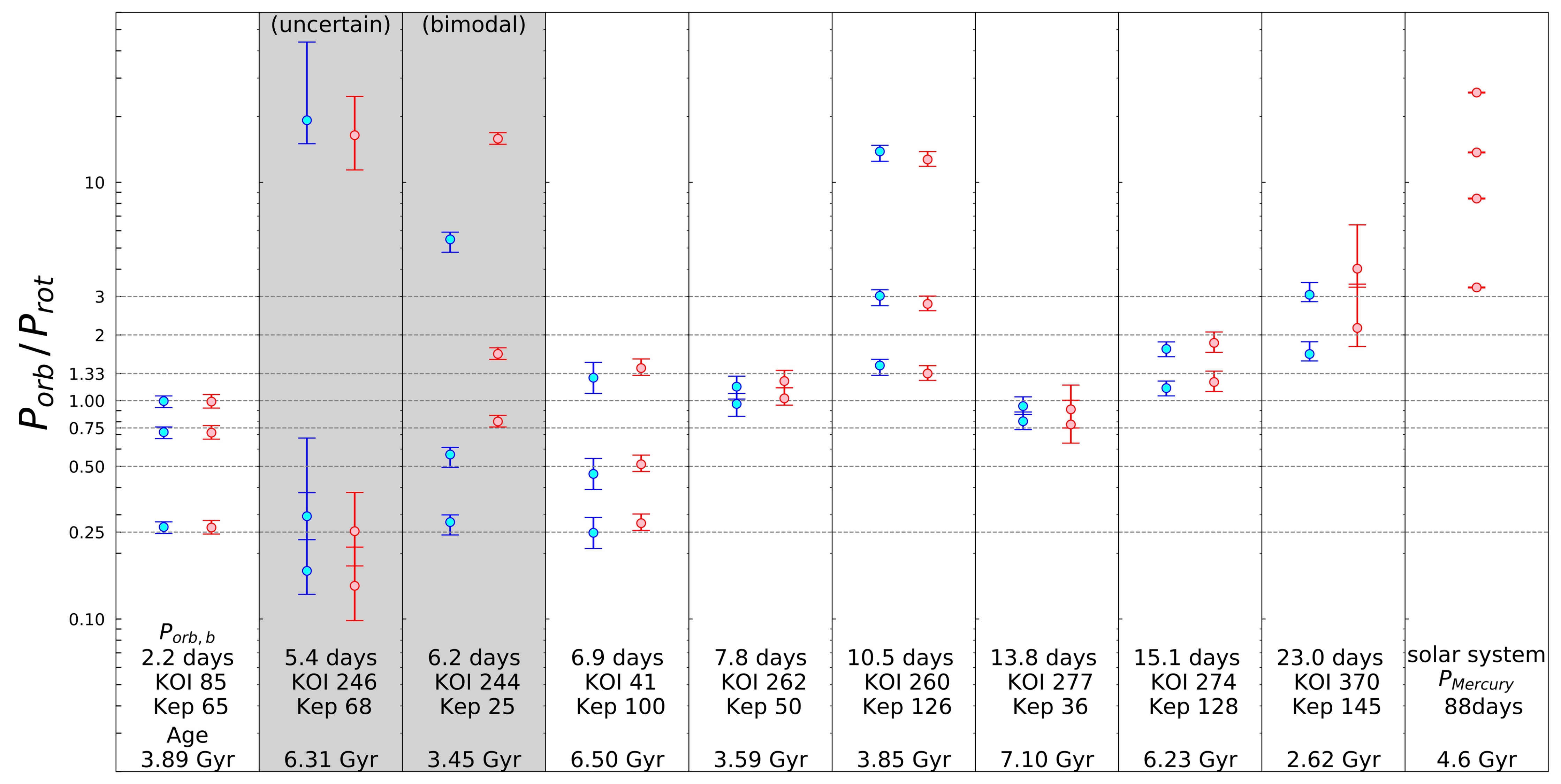



2.4R_E, 5.4day-orbit
+ 1R_E, 9.6day-orbi

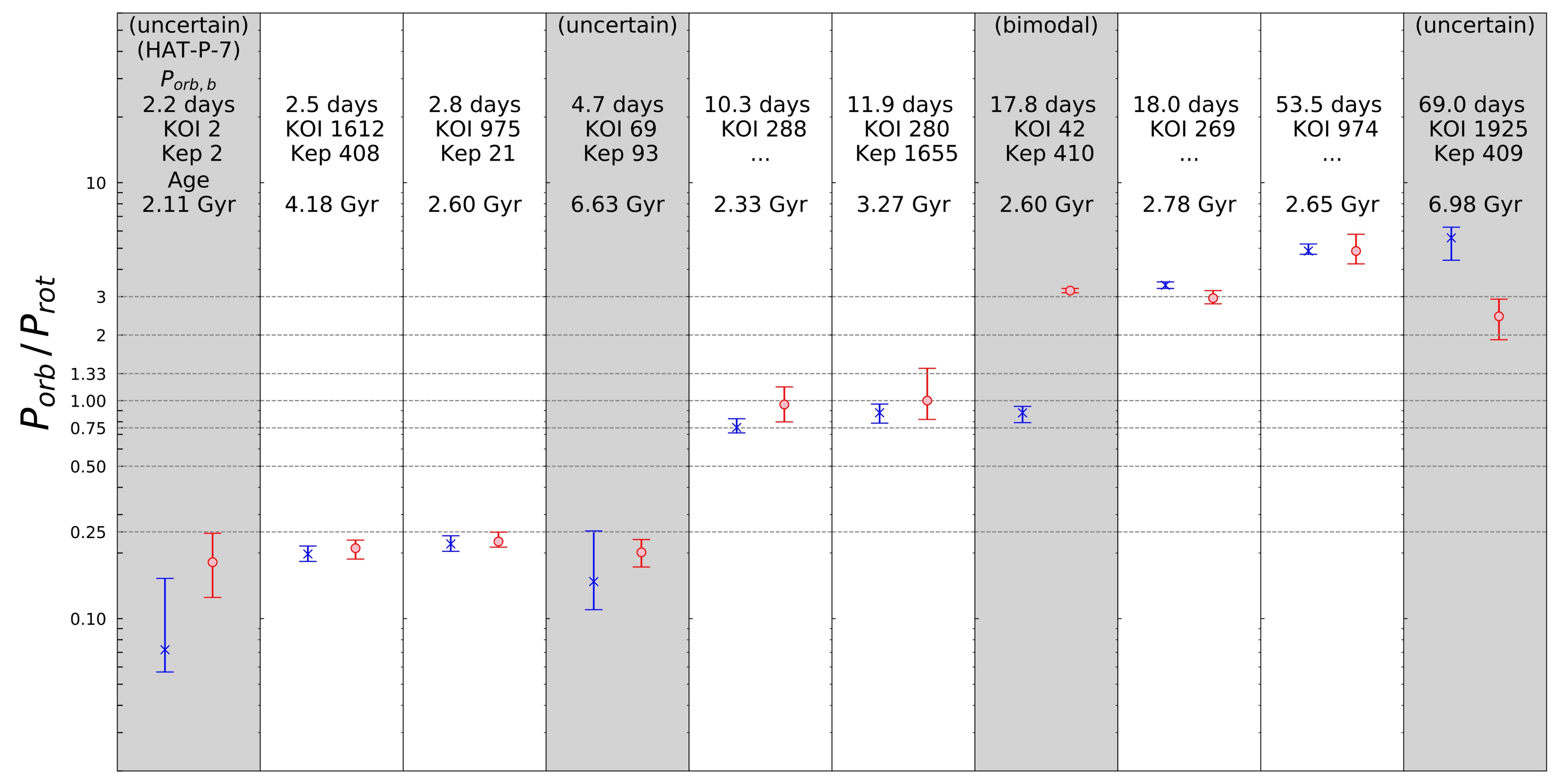



1.6R_E, 4.7day-orbit
+ >3700day-orbit planet ???

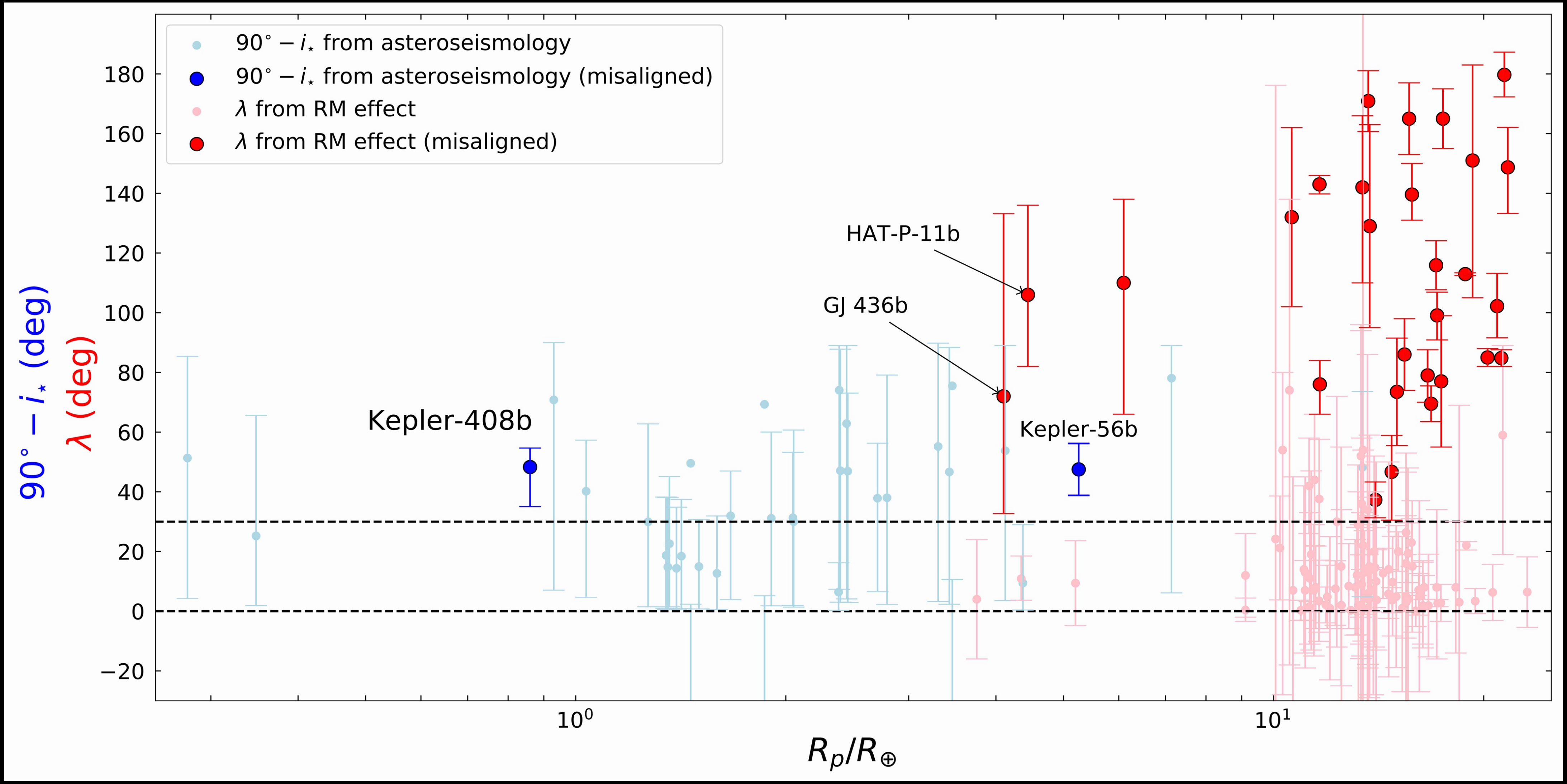
Possible spin-orbit synchronization ?



Possible spin-orbit synchronization ?

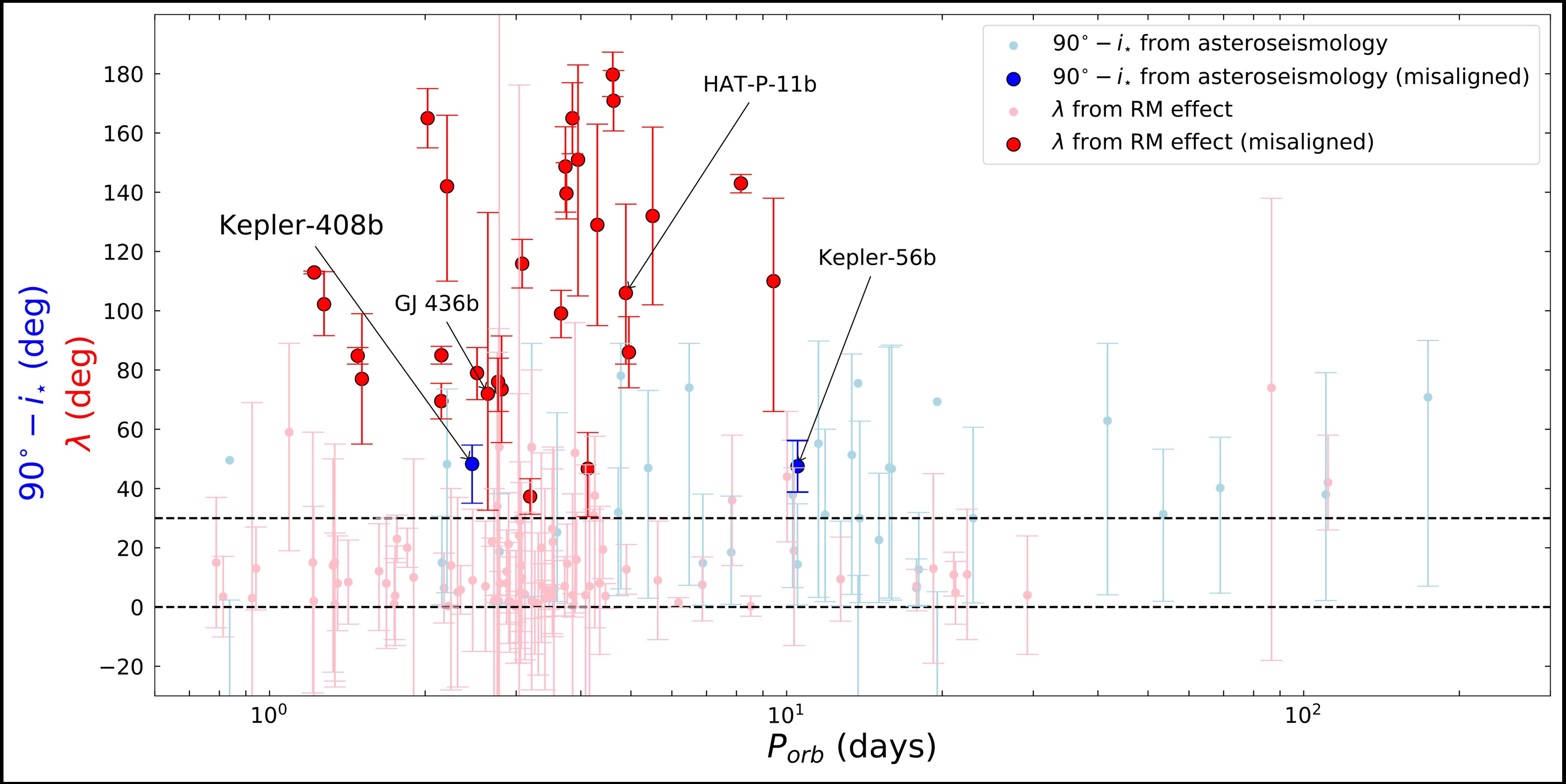


Spin-orbit angles against R_p



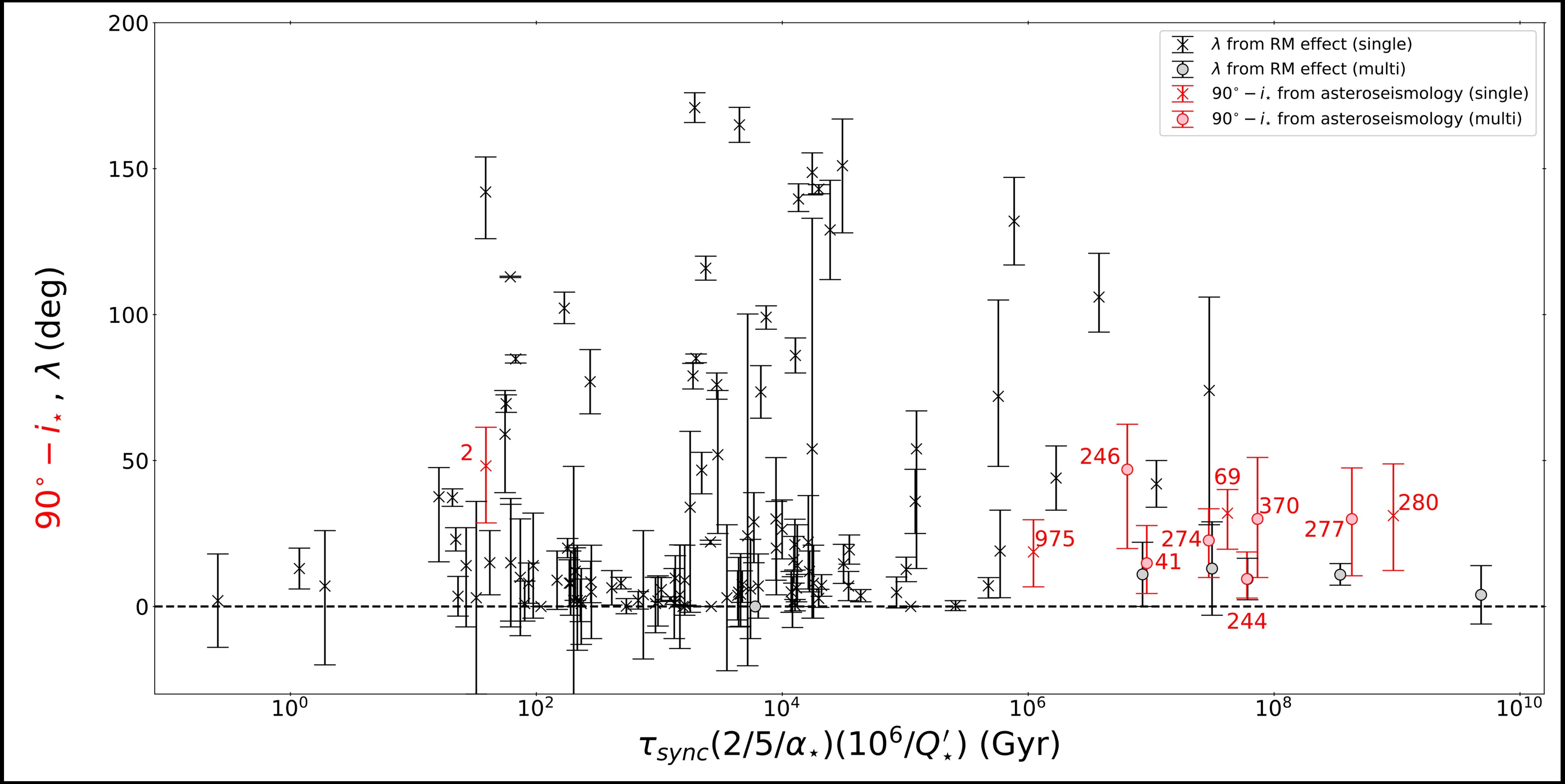
Kamiaka, Benomar, YS, Dai, Masuda, & Winn, AJ 157(2019)137

Spin-orbit angles against P_{orb}



Kamiaka, Benomar, YS, Dai, Masuda, & Winn, AJ 157(2019)137

Spin-orbit angles against the tidal synchronization time-scale



YS, Kamiaka & Benomar AJ 157(2019)172

Summary

- The Rossiter-McLaughlin effect and asteroseismology revealed quite unexpectedly large diversities in the spin-orbit architecture of planetary systems (～30 percent misaligned)
- The origin is not well understood
  - Nature vs. Nurture ?
  - Initial condition imprinted in protoplanetary disks ?
  - Chaotic dynamics in planet-planet interaction ?
  - Tidal interaction between the host star and planets ?
- Numerical simulations with realistic initial conditions !