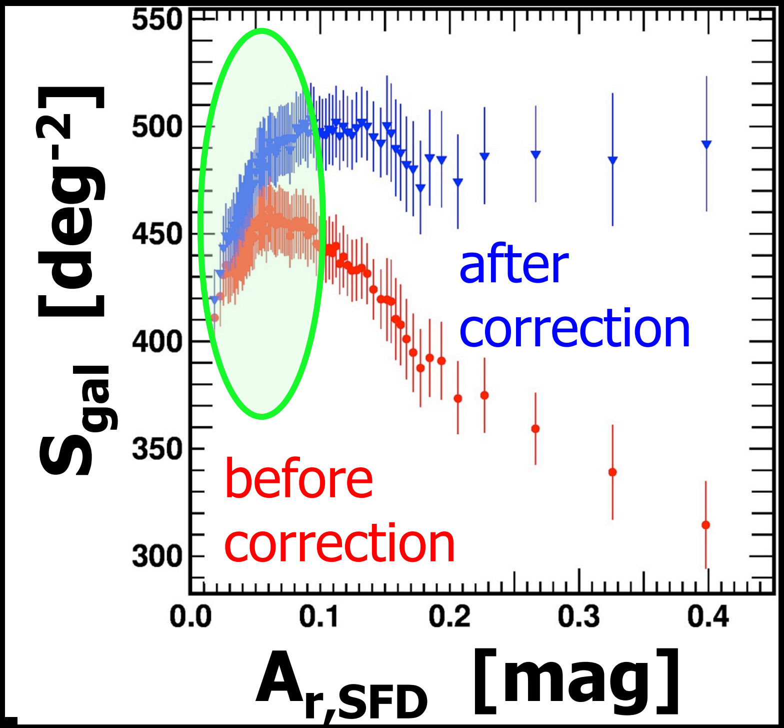
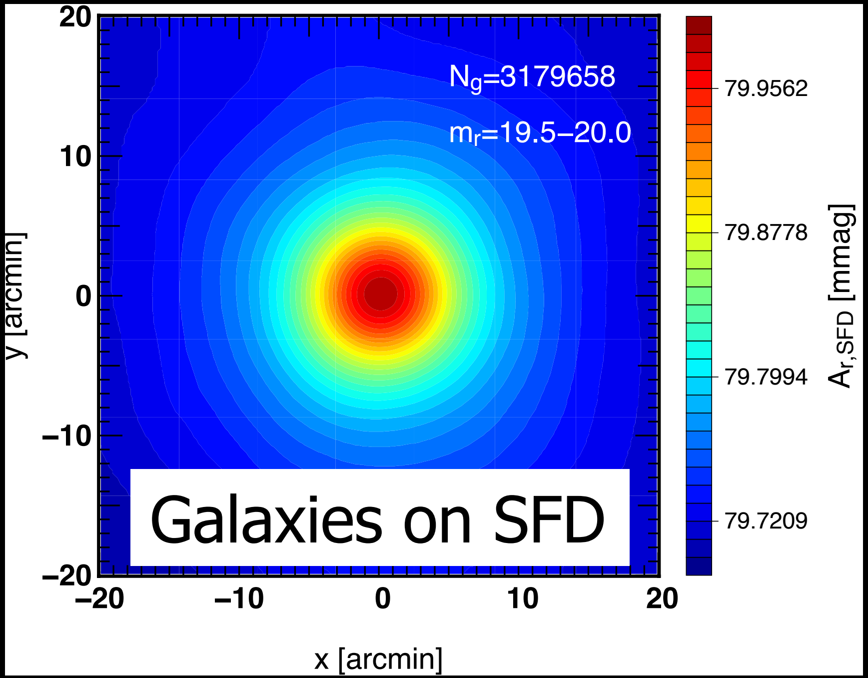
Anomaly in the SFD extinction map and discovery of FIR emission of galaxies by stacking analysis of the SDSS DR7 sample





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& Global scholar, Dept. of Astrophysical Sci., Princeton University

16:00-17:00 October 322, 2013

CAS seminar@#462, Bloomberg center, JHU

This talk is based on

- *Detection of Far Infrared Emission from Galaxies and Quasars in the Galactic Extinction Map by Stacking Analysis*
  - **T.Kashiwagi,** K.Yahata & YS
  - Publ.Astron.Soc.Japan 65(2013)43

- *The effect of FIR emission from SDSS galaxies on the SFD Galactic extinction map*
  - K.Yahata, A.Yonehara, YS, E.L.Turner T.Broadhurst, & D.P. Finkbeiner
  - Publ.Astron.Soc.Japan 59(2007)205

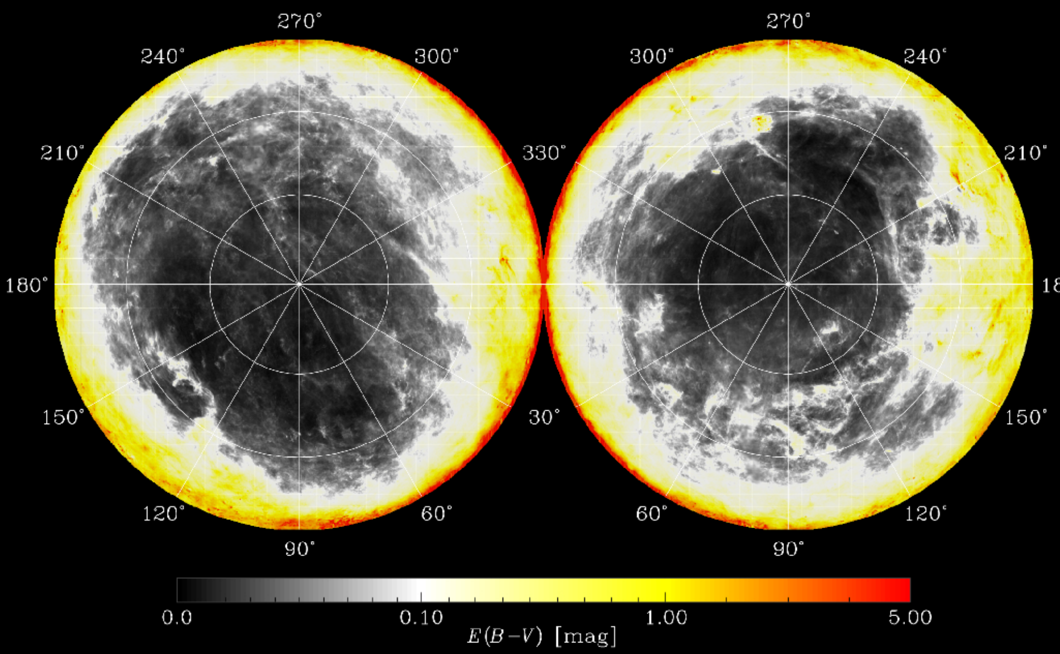
Testing a reliability of the SFD Galactic extinction map with SDSS galaxy number counts

- K.Yahata, A.Yonehara, YS, E.L.Turner, T.Broadhurst, & D.P. Finkbeiner

  Publ.Astron.Soc.Japan 59(2007)205
- T.Kashiwagi

  Master thesis (2011) submitted to U.Tokyo

SFD Galactic extinction map



Galactic extinction E(B-V) map (Schlegel, Finkbeiner & Davis 1998; SFD)

- The most fundamental dataset for all astronomical observations
- True large-scale structures revealed only after the extinction correction
- Its reliability is of vital importance in precision cosmology

SFD procedure to construct the Galactic extinction map

- COBE 100 μm+240 μm maps (0.7deg.pixel)
  - Remove zodiacal light and cosmic infrared background
  - Dust temperature map ⇒ temperature-dependent emissivity corrected 100 μm map
- Calibration of higher angular-resolution IRAS 100 μm map (5 arcmin. pixel)
- Assume

$$E(B-V)=pI_{100\mu m}X(T)$$

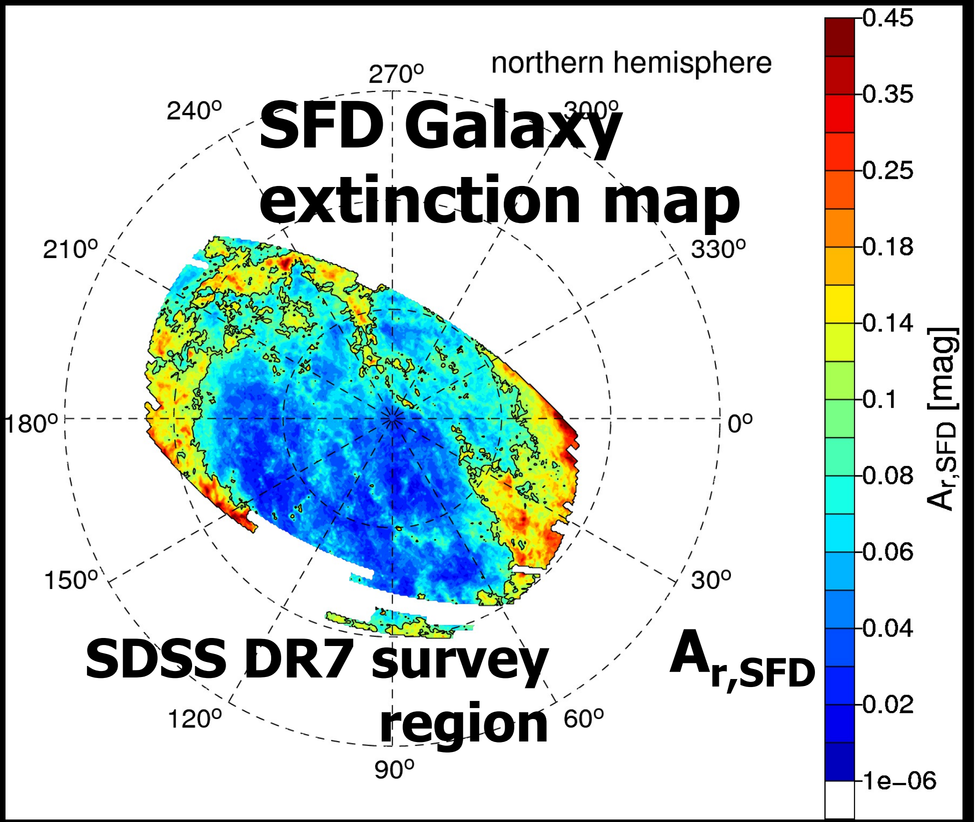
dust temperature correction factor

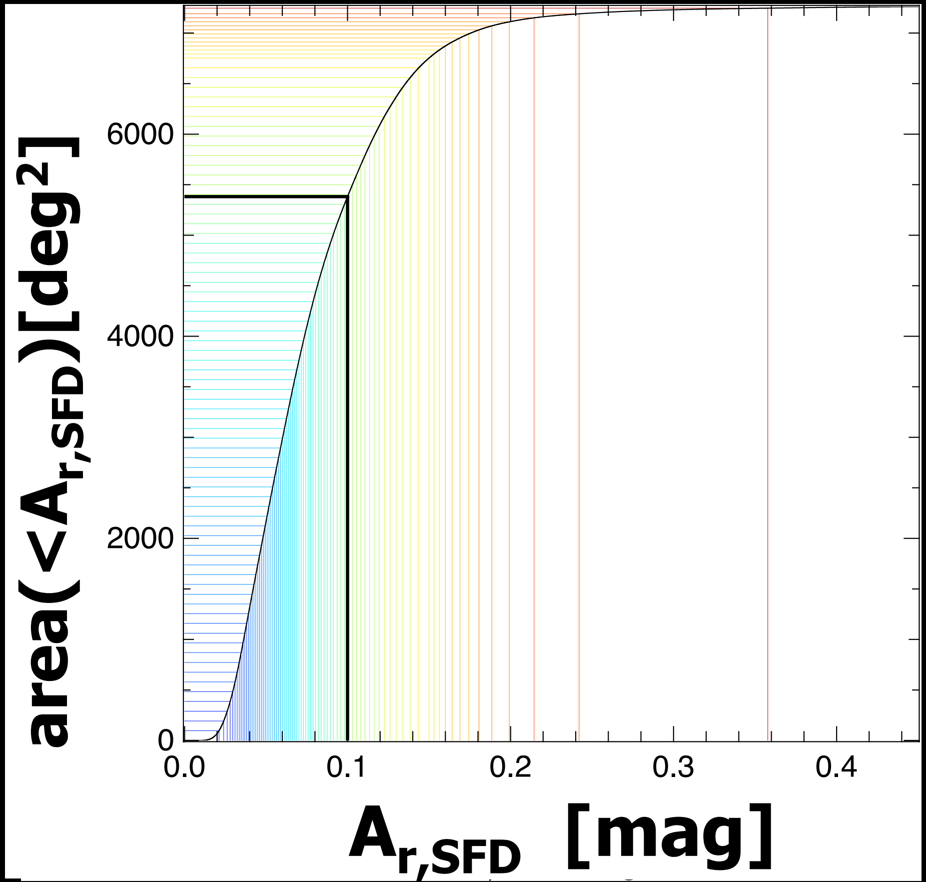
at each region and determine p～0.0184 from the data

- Convert E(B-V) to A_{band} adopting $R_V=A_V/E(B-V)=3.1$

A_{SFD} map in SDSS DR7 survey region

3.6×10^6 galaxies ($17.5<r<19.4$) in 7270 deg^2
from SDSS DR7 photometric catalog

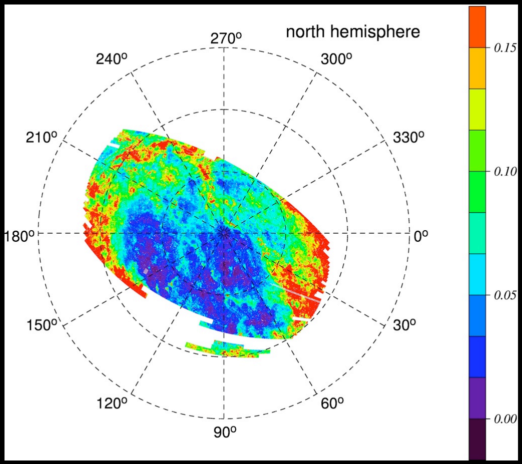


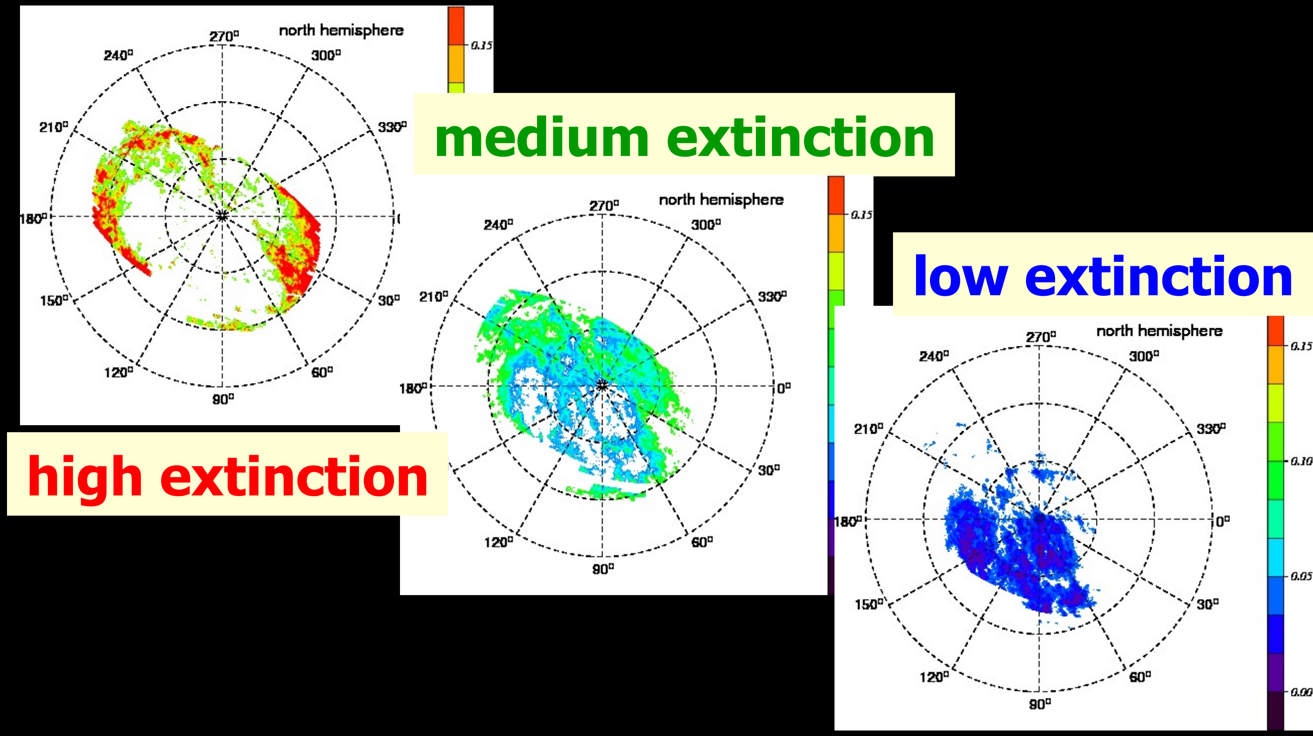


Kashiwagi (2011), Yahata et al.(2007)

Estimating Galactic extinction from SDSS galaxy surface density

SDSS DR7 survey area (color coded according to A_{SFD})

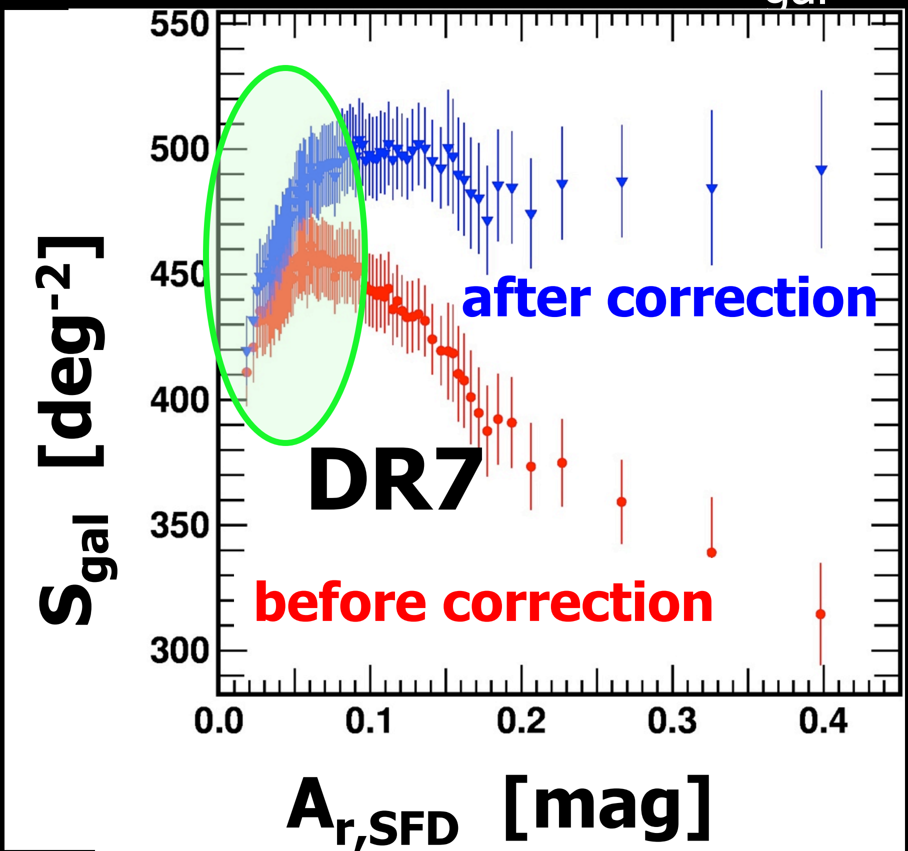




- divide the SDSS DR7 survey area into many small regions according to A_{SFD}
- combine those non-contiguous regions into 84 bins for A_{SFD} with ~100 deg^2 each
- compute the SDSS galaxy number density S_{gal} for those bins

Anomaly of SDSS galaxy surface density S_{gal} as a function of A_{SFD}

- ***If A_{SFD} is perfect, we expect that***
  - Before correction: S_{gal} should monotonically decrease as a function of A_{SFD}
  - After correction: S_{gal} should be constant



(cosmological principle!)

- OK for A_{SFD} >0.1, but *quite the opposite for A_{SFD} <0.1*
  - ~70% of the SDSS survey area has A_{SFD} <0.1 !
- First pointed out by Yahata et al. (2007) for DR4, and confirmed by Kashiwagi (2011) for DR7

Origin of the anomaly

- A_{SFD} is estimated assuming that the ***extinction*** is proportional to the FIR ***emission*** flux (100 μ m)
  - the anomaly indicates the positive correlation between galaxy surface density and the FIR flux at least where the real extinction is small
- 100 μ m flux = Galactic dust + galaxies
  - contamination by the FIR emission from galaxies proposed by Yahata et al. (2007); indeed originally suggested by Ed Turner

Numerical and analytic models to explain the anomaly of SFD map from the FIR emission of galaxies

- T.Kashiwagi

  Master thesis (2011) submitted to U.Tokyo
- T.Kashiwagi, YS, A.Taruya, I.Kayo, T.Nishimichi & K.Yahata

  to be submitted to ApJ (2013)

Mock simulations to test the FIR emission hypothesis

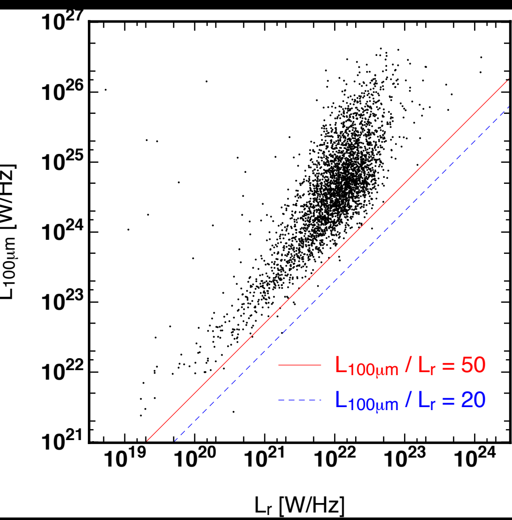
- Distribute random particles over the DR7 region
  - the same number and the same r-band magnitude distributions as SDSS galaxies
- Assign 100μm flux to each particle sampled from the log-normal distribution of $y=L_{100\mu m}/L_r$
- Add the 100μm flux to the original SFD map and compute the extinction

  $$A_{SFD} \Rightarrow A_{SFD} + \Delta A_{mock}$$

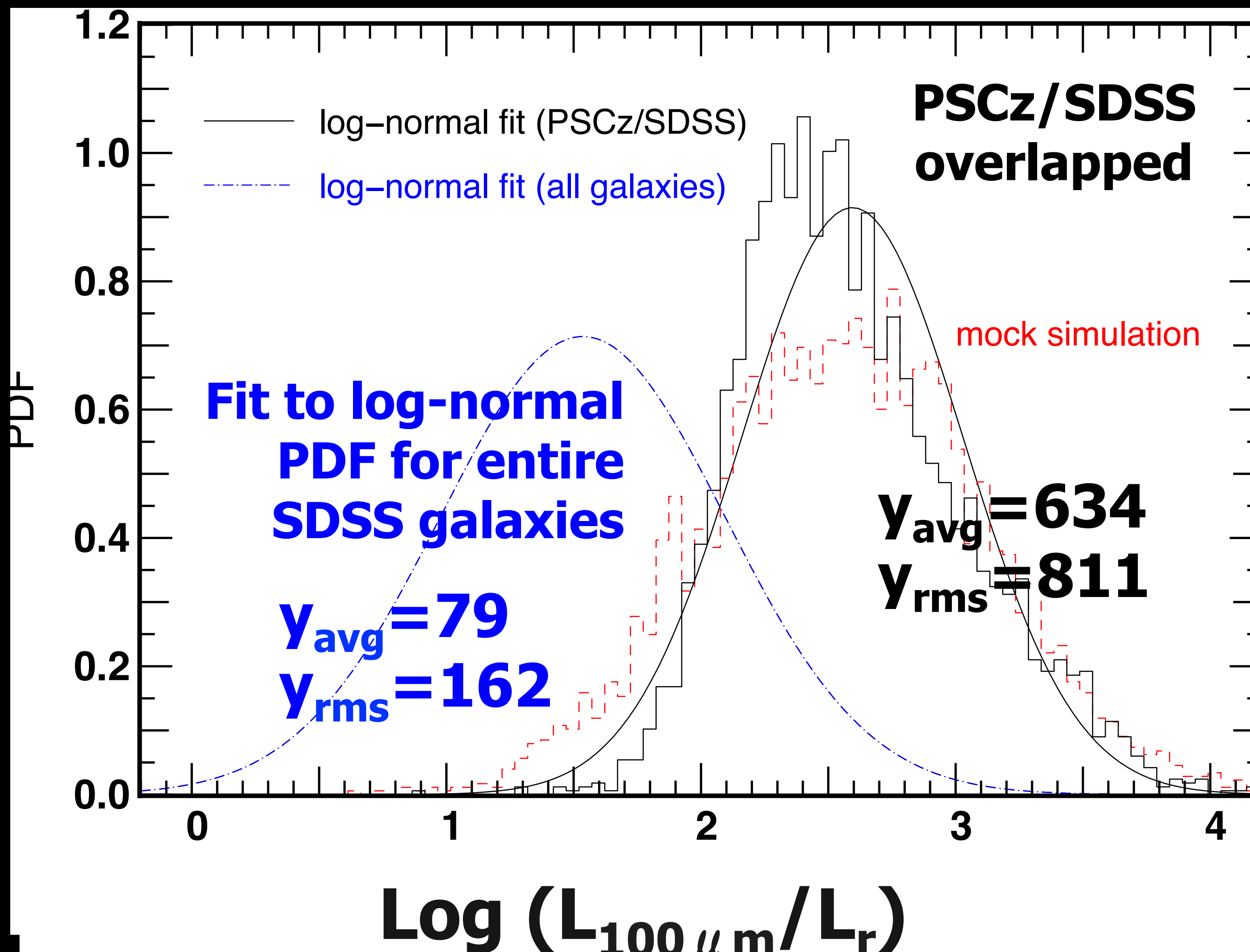
  at each pixel (5'x5')
- Compute particle surface density as a function of $A_{SFD} + \Delta A_{mock}$ at each pixel

Correlation of $L_{100\mu m}$ and L_r for galaxies

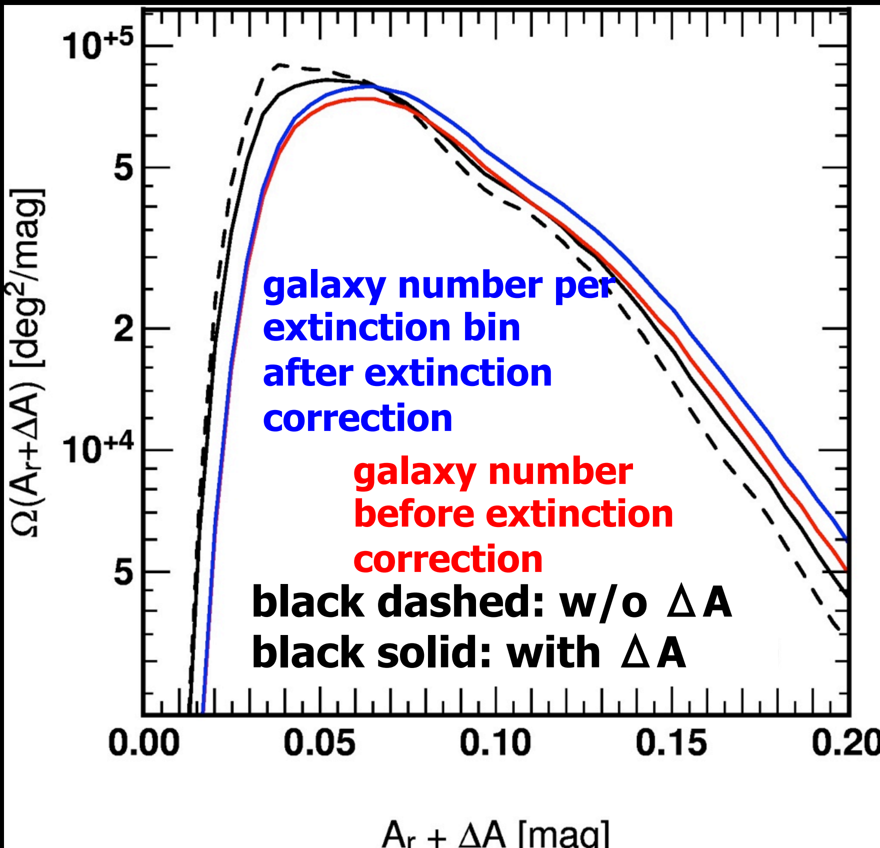
- PSCz & SDSS overlapped (FIR luminous) sample (~3300gals)

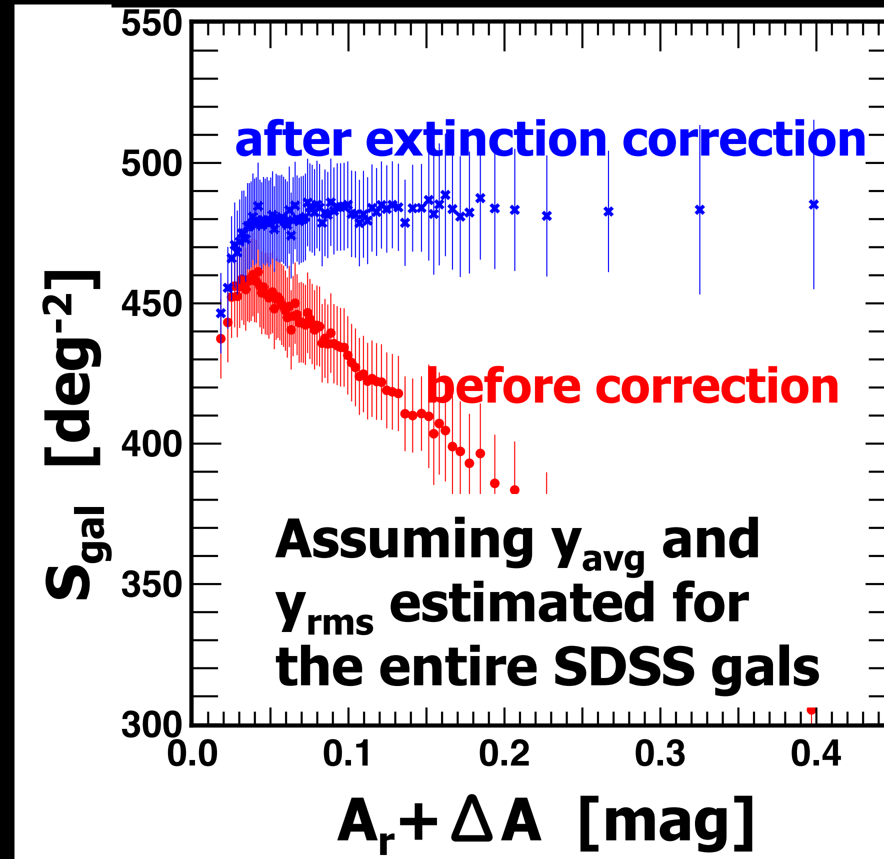


PSCz/SDSS overlapped sample (3300 galaxies)



Mock simulation result



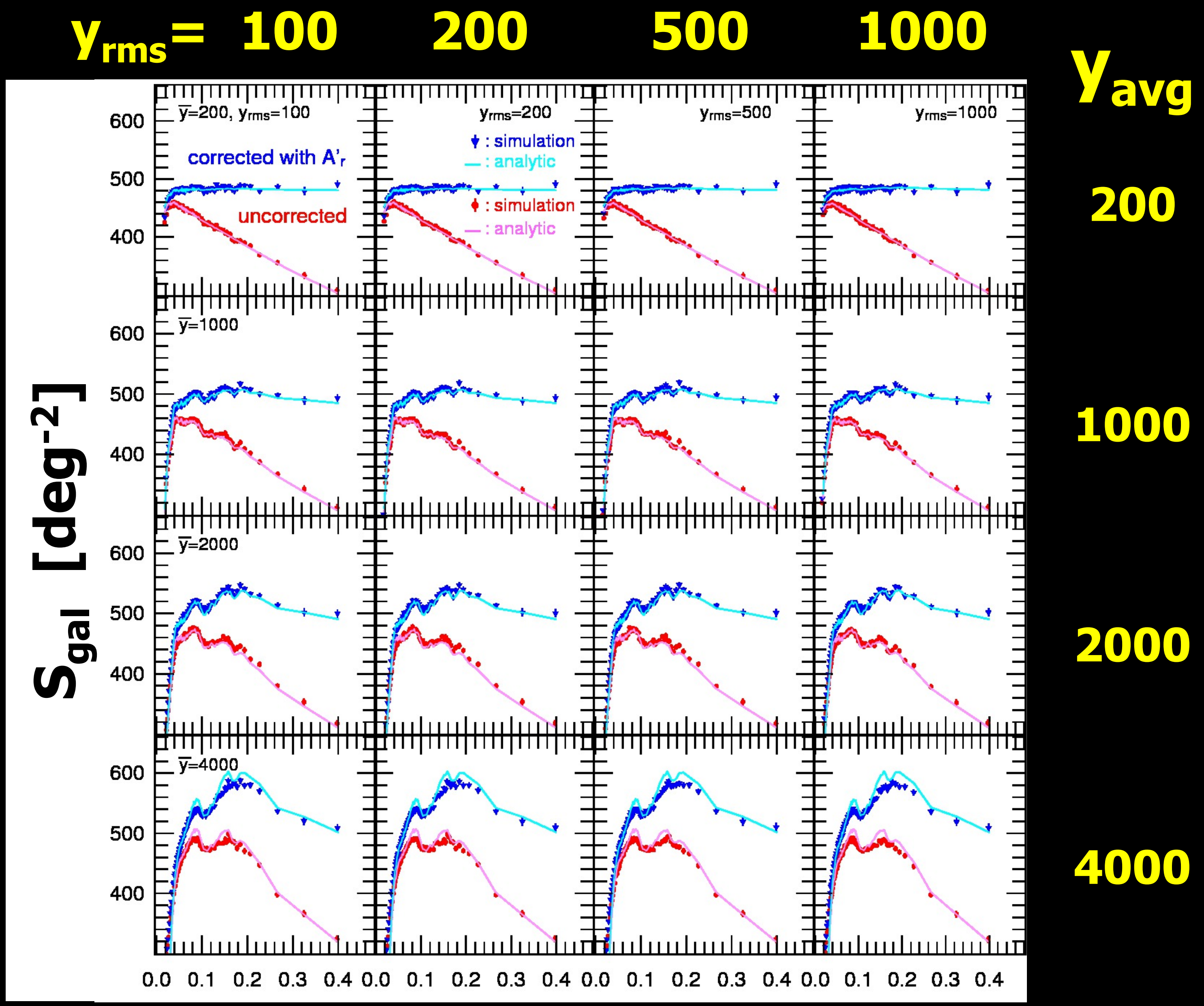


- the trend of the observed anomaly is reproduced, but weaker if the mean $L_{100\mu m}/L_r$ estimated for the entire SDSS galaxies is adopted

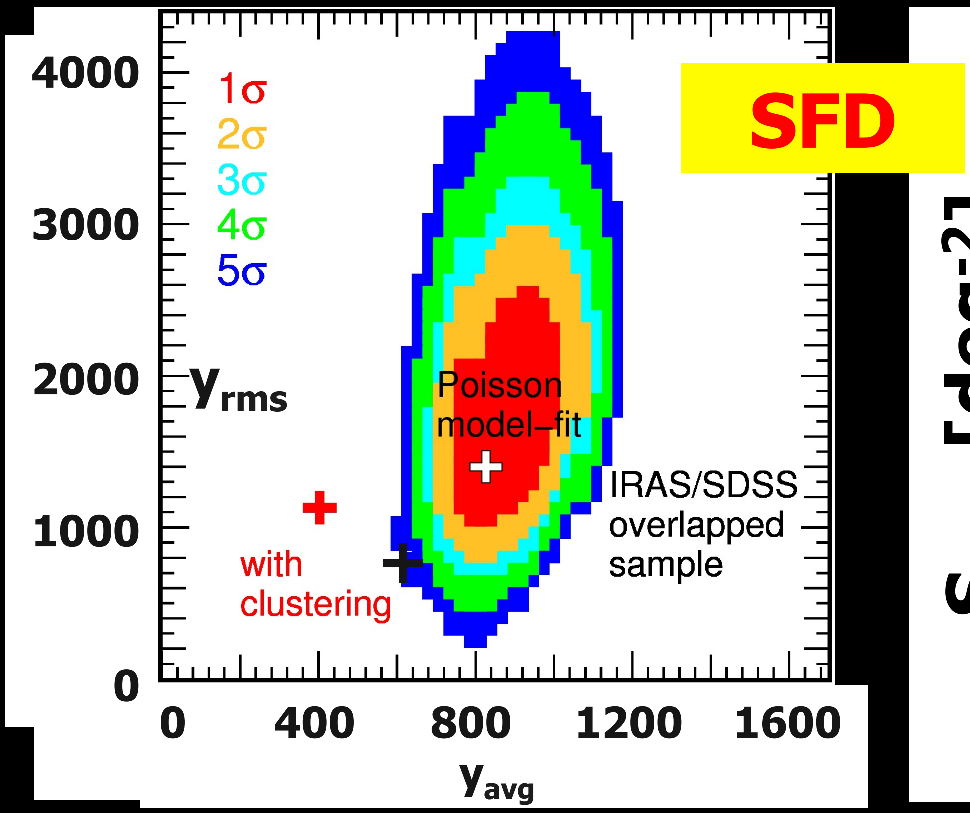
Analytic model for S_{gal} with contamination of FIR galaxy emission

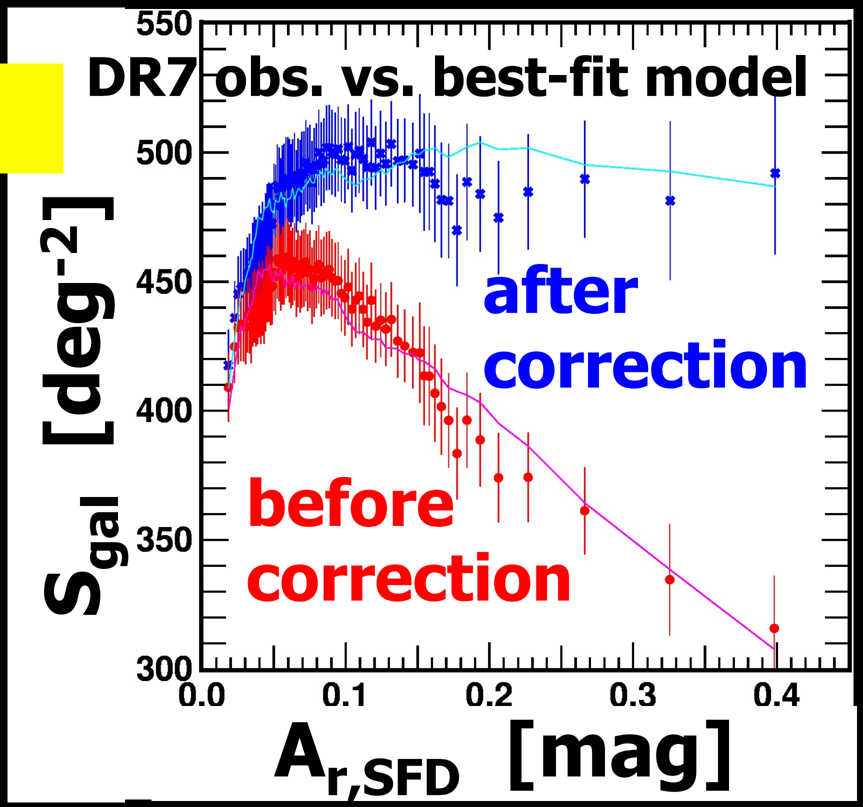
- Poisson distributed galaxies (spatial clustering is ignored)
- Log-normal PDF for $y = L_{100\mu m}/L_r$ of each galaxy (characterized by y_{avg} and y_{rms})
- Compute the PDF of the additional extinction (converted from the FIR emission of galaxies) for a pixel on the SFD map with N galaxies $P_N(\Delta A)$
- Compute $N(A')$ and $\Omega(A') \Rightarrow S_{gal}(A')$ where $A'=A+\Delta A$

Mock simulation vs. analytic prediction



Fit to the observed anomaly using the analytic model





- Observed trend is well reproduced by the model
- Best-fit y_{avg}=800 (clustering ignored)

  ⇒ y_{avg}=400 (corrected for clustering)

Can we directly detect the FIR emission of galaxies ? - Stacking image analysis of SDSS galaxies -

- T.Kashiwagi, K.Yahata & YS

  Publ.Astron.Soc.Japan 65(2013)43

Direct detection of FIR emission of galaxies

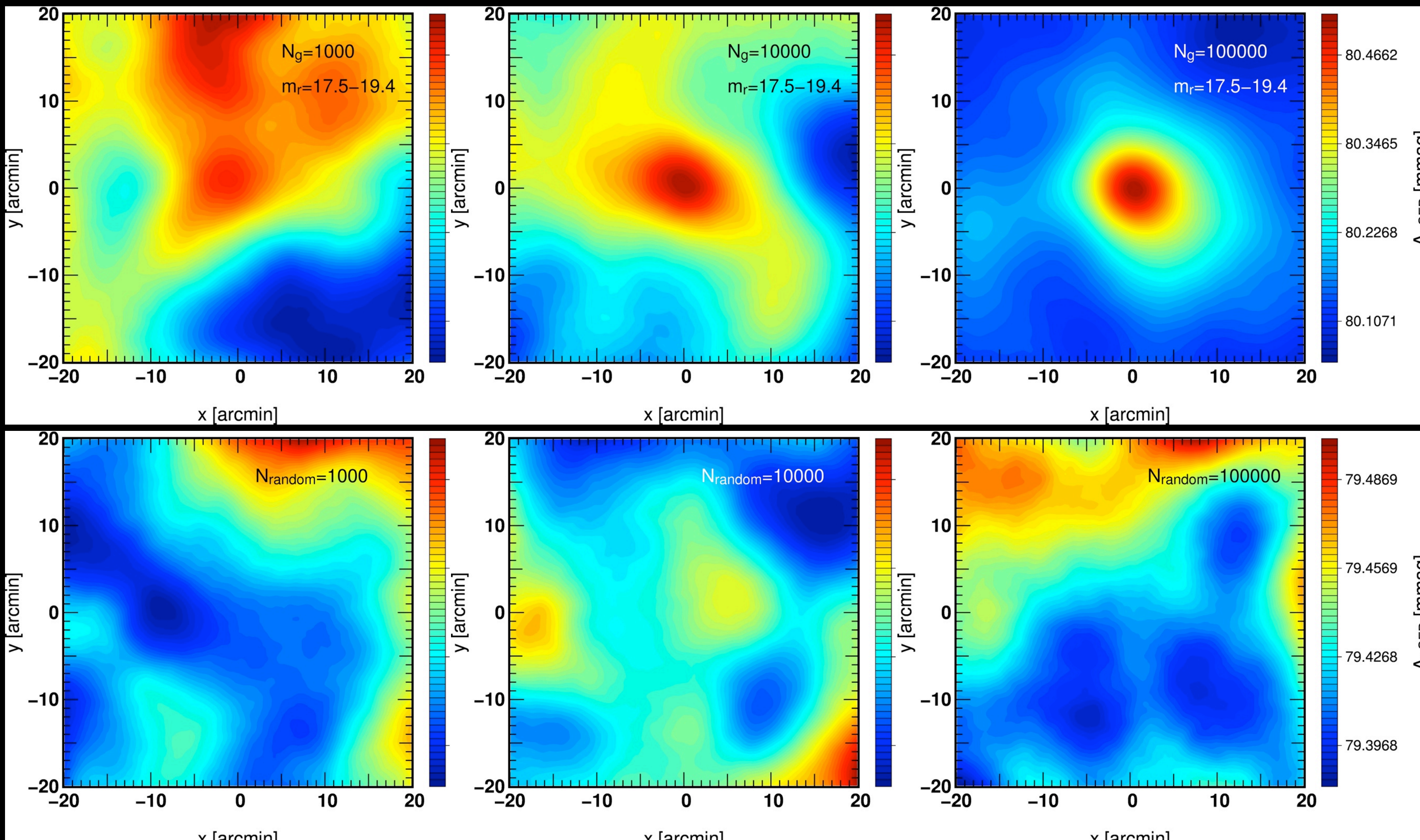
- FIR emission of a majority of SDSS photometric galaxies is weak and cannot be detected individually.
- Can we detect their FIR emission statistically through stacking SDSS galaxies over the SFD map ? ⇒ *Yes !*

Kashiwagi, Yahata & YS
Publ.Astron.Soc.Japan 65 (2013)43

Stacking analysis of SDSS galaxies on the SFD map

galaxy

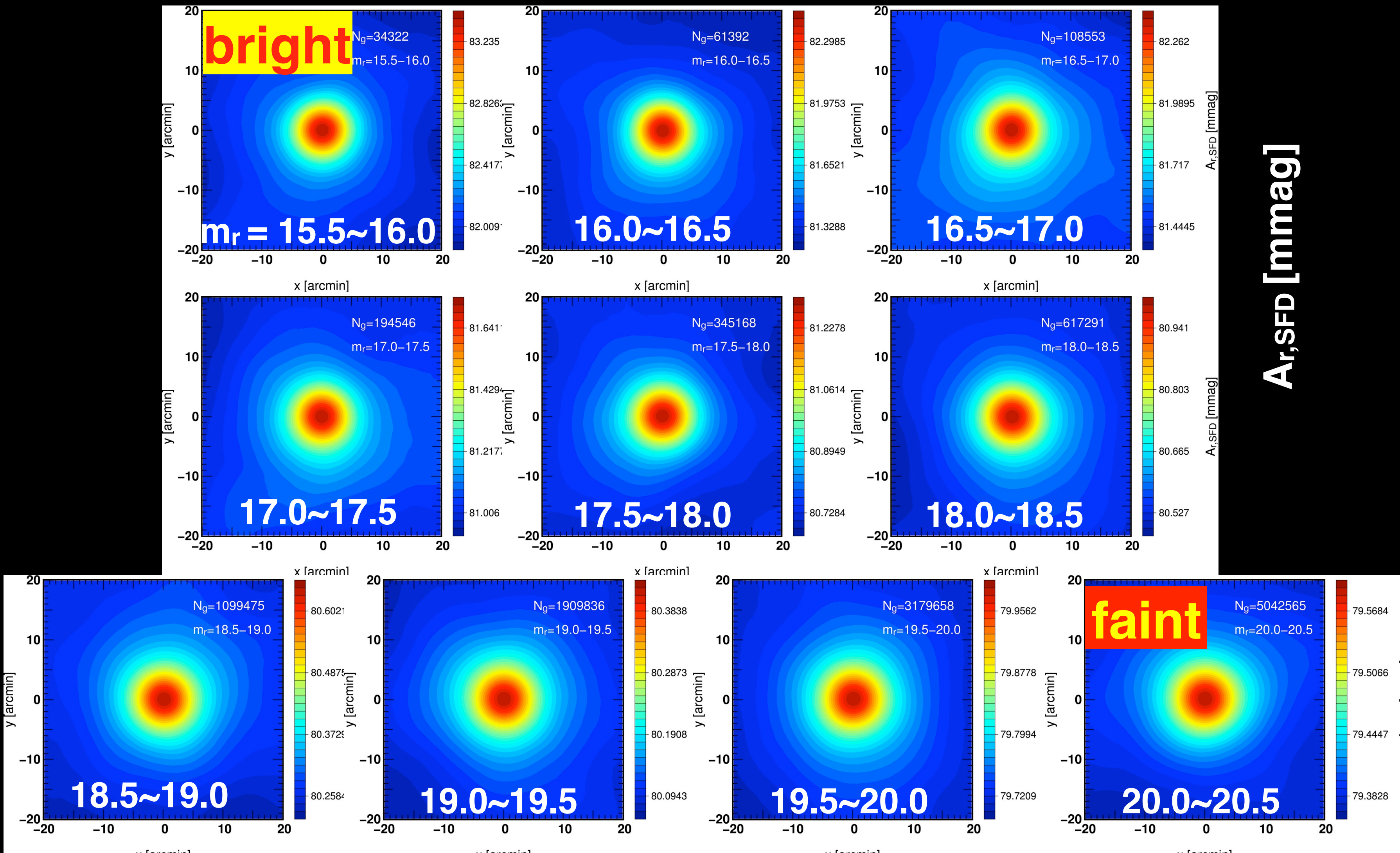
random



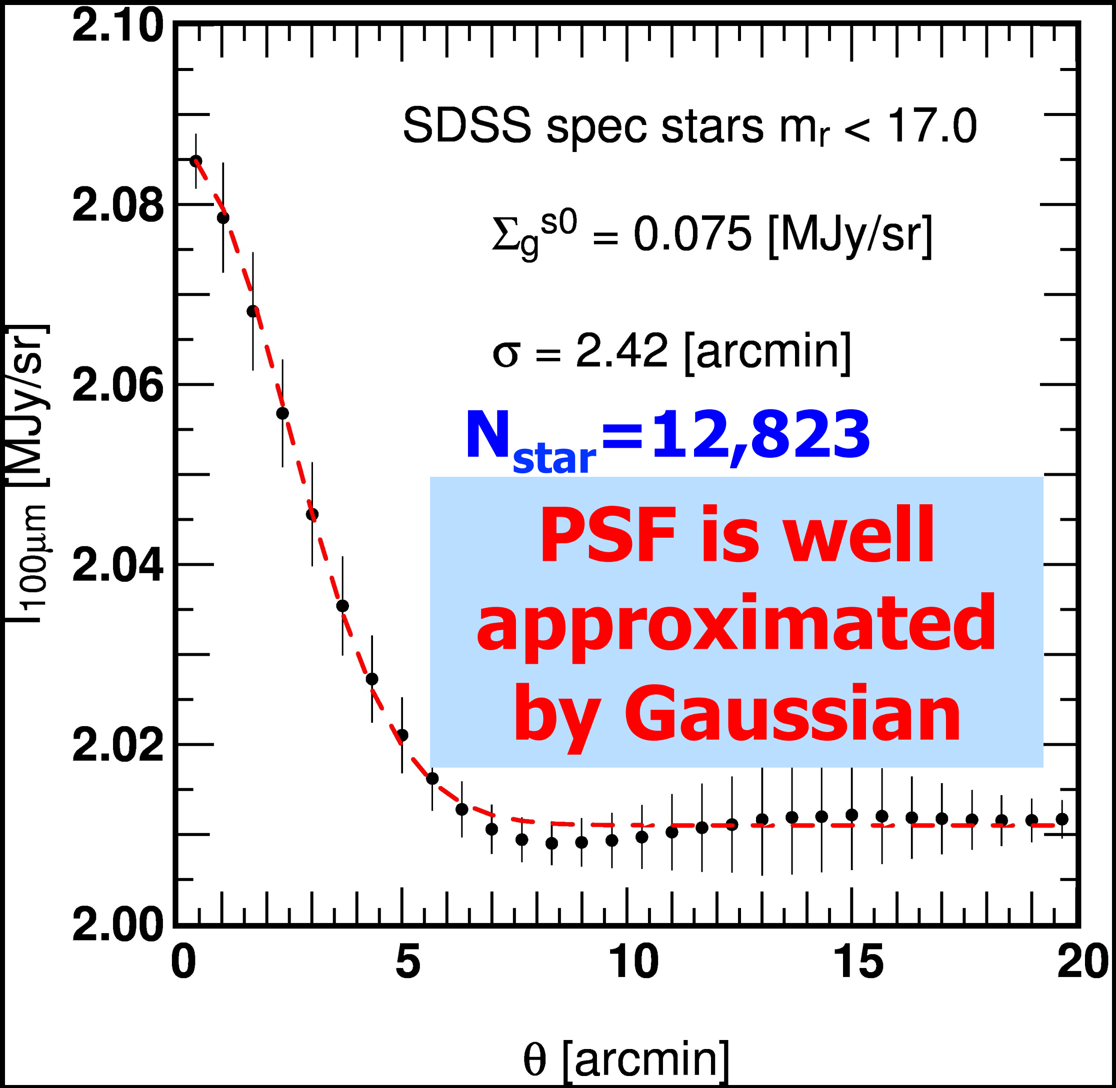
N = 10^3 N = 10^4 N = 10^5

Magnitude dependence

Stacking SDSS galaxies ($15.5 < m_r < 20.5$) over SFD map according to their r-band magnitude ($\Delta m_r = 0.5$)

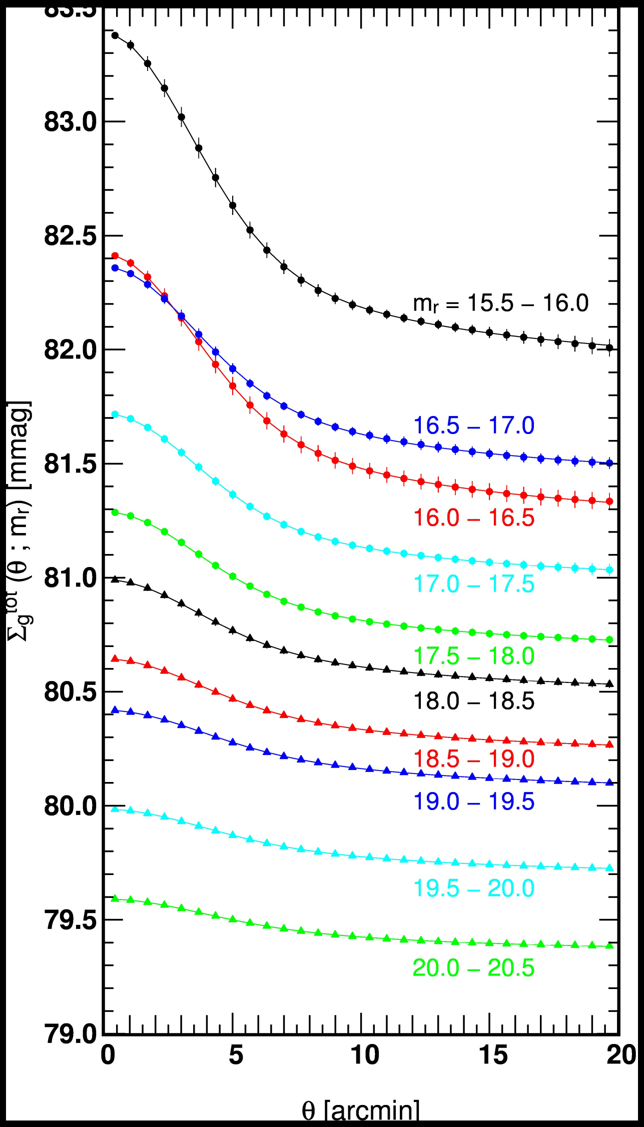


Point spread function of IRAS 100 μm map



- Angular resolution of SFD (IRAS) is low,
  ⇒ we need to its PSF to understand the stacked profile
- PSF is difficult to measure on SFD map since most of point sources are removed
- We use the original IRAS 100 μm diffuse map, and stack stars with r<17 mag.

Decompositions into single galaxy and clustering terms



$$\Sigma_{\rm g}^{\rm tot}(\theta;m_r) = \Sigma_{\rm g}^{\rm s}(\theta;m_r) + \Sigma_{\rm g}^{\rm c}(\theta;m_r) + C$$

$$\Sigma_{\rm g}^{\rm s}(\theta;m_r) = \Sigma_{\rm g}^{\rm s0}(m_r)\exp\left(-\frac{\theta^2}{2\sigma^2}\right)$$

$$\Sigma_{\rm g}^{\rm c}(\theta;m_r) = \iint dm' d\varphi\ \Sigma_{\rm g}^{\rm s}(\theta-\varphi;m') \times w_{\rm g}(\varphi;m',m_r)\frac{dN_{\rm g}(m')}{dm'}$$

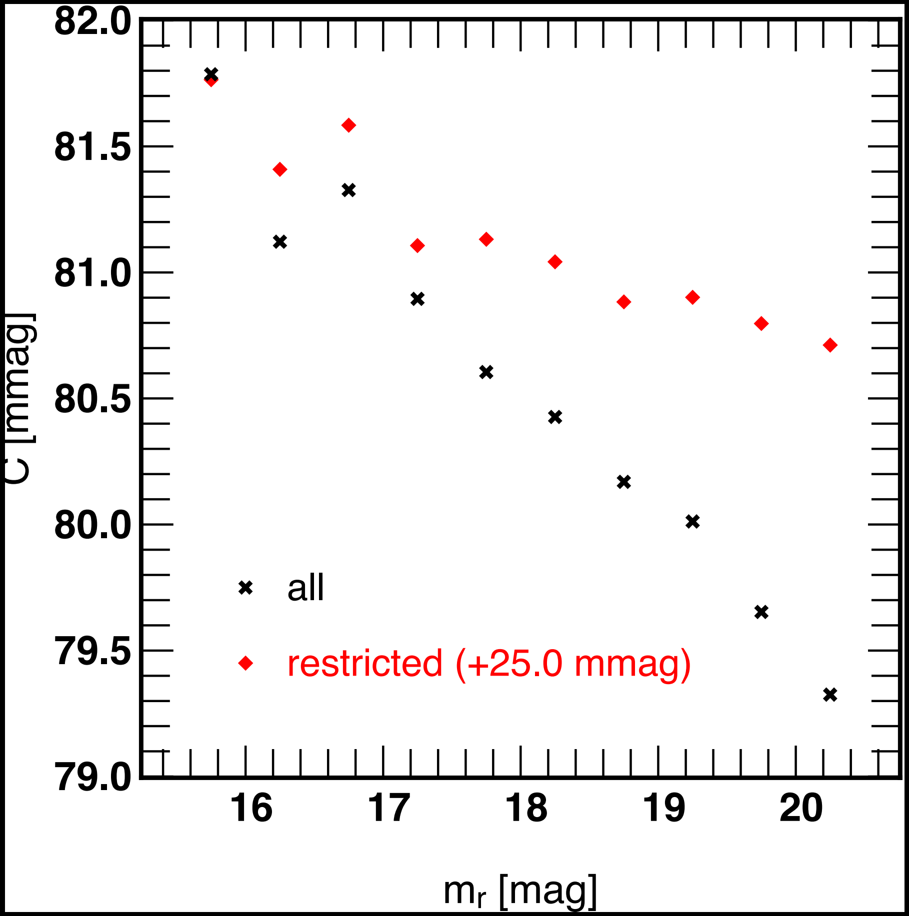
$$w_{\rm g}(\varphi;m',m_r) = K(m',m_r)(\varphi/\varphi_0)^{-\gamma},$$

$$\Sigma_{\rm g}^{\rm c}(\theta;m_r) = \Sigma_{\rm g}^{\rm c0}(m_r)\exp\left(-\frac{\theta^2}{2\sigma^2}\right) \times {}_1F_1\left(1-\frac{\gamma}{2};1;\frac{\theta^2}{2\sigma^2}\right)$$

simultaneous fit ⇒ $\sigma = 3.1'$

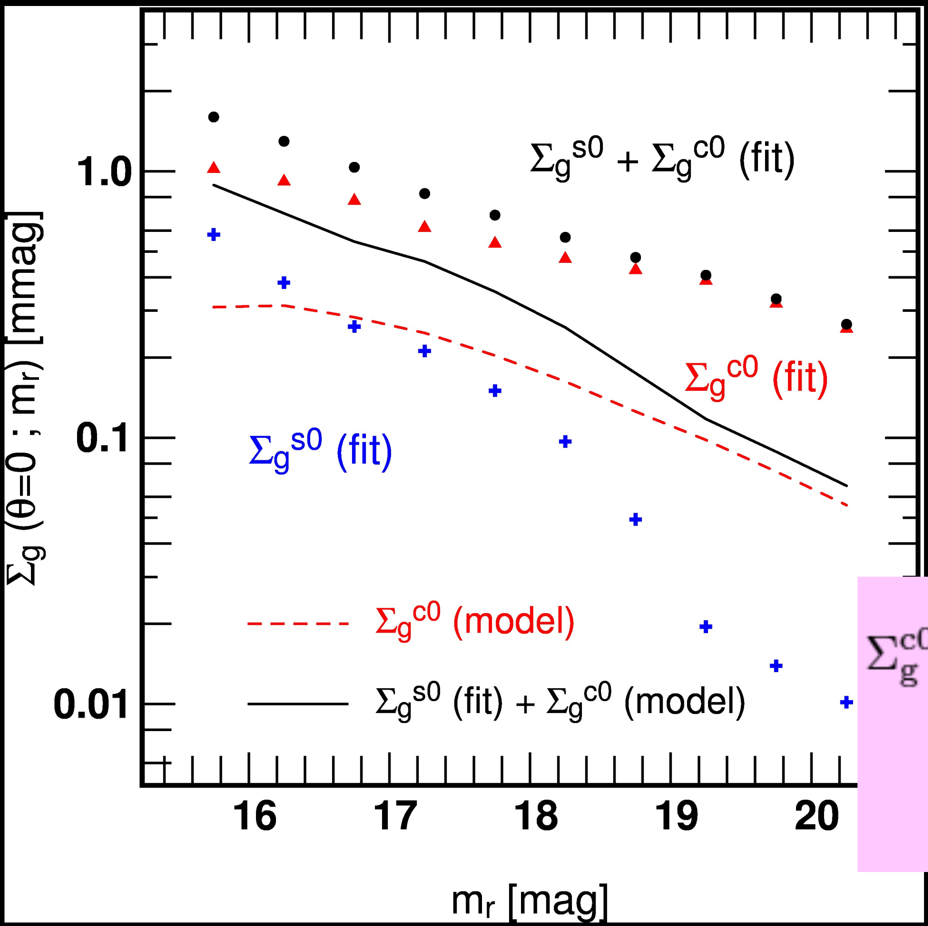
The reason why the fitted C is not constant

$$\Sigma_{\rm g}^{\rm tot}(\theta; m_r) = \Sigma_{\rm g}^{\rm s}(\theta; m_r) + \Sigma_{\rm g}^{\rm c}(\theta; m_r) + C$$



- This is largely due to the cfa Great Wall that is preferentially located in relatively high-extinction region of Galaxy.
  - The trend disappears when we remove the cfa Great Wall region.

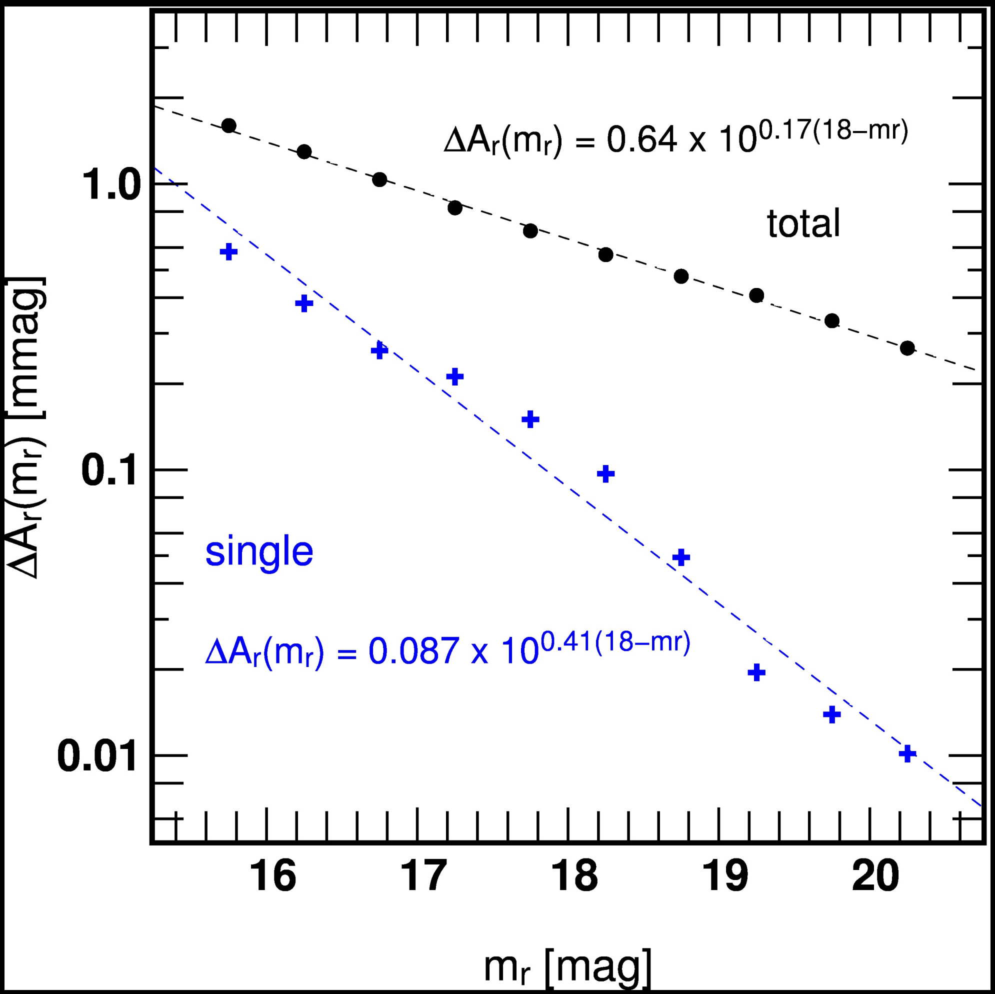
Extended dust emission around the halo hosting the central galaxy and/or contribution from unresolved galaxies ?



- The fitted clustering term is a factor of 2-3 larger than that expected from the measured angular correlation functions of resolved SDSS galaxies

$$\Sigma_{\rm g}^{\rm c0}(m_r) = 2\pi\sigma^2 \left(\frac{\varphi_0}{\sqrt{2}\sigma}\right)^{\gamma} \Gamma\left(1-\frac{\gamma}{2}\right) \times \int dm' \Sigma_{\rm g}^{\rm s0}(m') K(m', m_r) \frac{dN_{\rm g}(m')}{dm'}$$

Average contribution to A_r against m_r of the central galaxy (SDSS)



- Could be used as an empirical correction for the SFD extinction of a galaxy with m_r
  - but tiny and just statistical
  - not clear if this correction works in reality

c.f. Peek & Graves (2010)

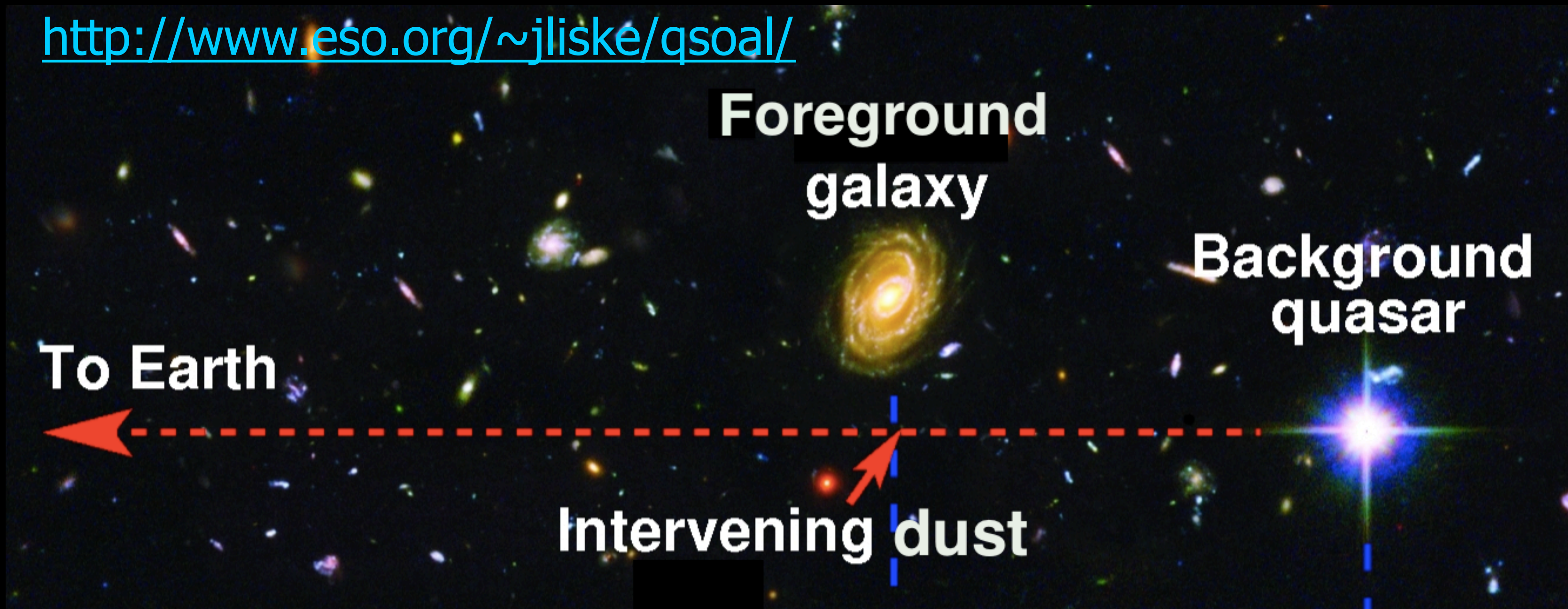
The spatial extent of the dust:

- associated with individual galaxies or extended over their common halos ? -

- T.Kashiwagi & YS: in preparation

Very preliminary !

Intergalactic dust is universal?



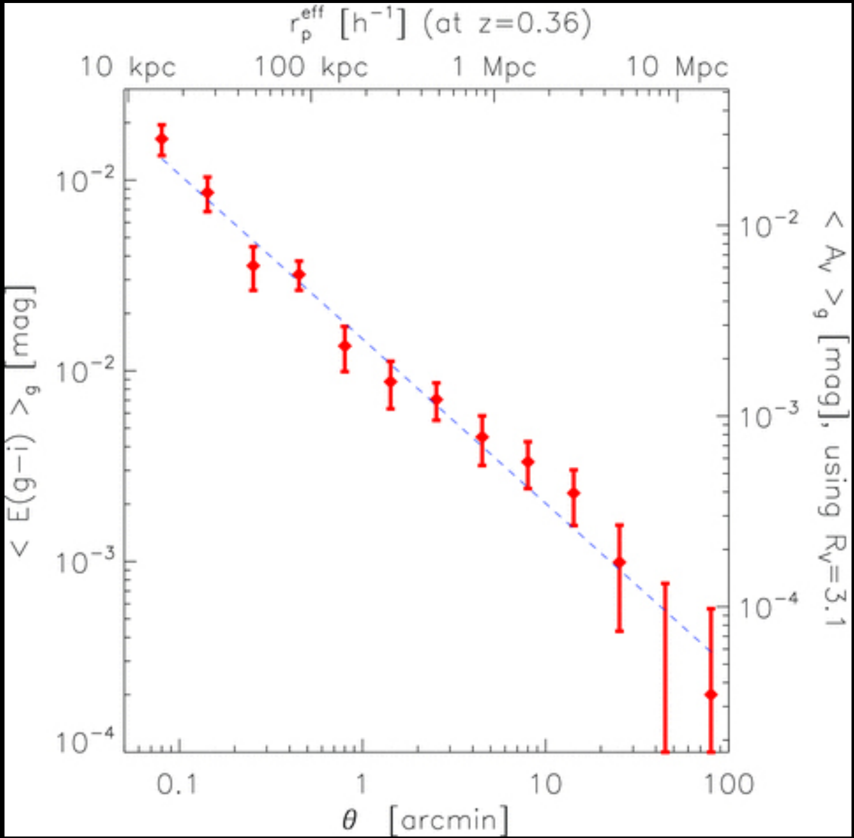
- Ménard, Scranton, Fukugita & Richards: MNRAS 405 (2010) 1025
  - Measure the reddening of background quasars due to the dust of SDSS galaxies from $<\delta m_Q(\Phi)\,\delta_g(\Phi+\theta)>$
  - Detected the presence of dust from 20kpc to several Mpc

Spatial distribution of intergalactic dust ?

$$\langle E(g-i)\rangle(\theta) = (1.5 \pm 0.4) \times 10^{-3} \left(\frac{\theta}{1\,\mathrm{arcmin}}\right)^{-0.86\pm0.19}$$

Ménard et al. (2010)

- Extended much beyond each galaxy ?
- Sum of dust associated with galaxies ?
  - Very similar to the galaxy angular correlation function power-law…



Stacking IRAS map to detect 100 μm emission of SDSS galaxies

- MSFR measure the ***absorption*** of dust
- Combining with the measurement of the ***emission*** of dust of galaxies, we constrain the dust temperature, which would distinguish intragalaxy and intracluster dust.
- Repeat the same stacking procedure
  - SDSS galaxies with $17<m_i<21$ that MSFR use
  - decomposition into three terms

$$I_{\rm total}(\theta, m_i) = I_{\rm single}(\theta, m_i) + I_{\rm clustering}(\theta, m_i) + C$$

Emission and absorption of dust

- Optical depth (extinction=scattering+absorption)

$$\tau(\theta,\lambda) = \kappa_{\rm ext}(\lambda)\Sigma_{\rm d}(\theta)$$

- Color excess

$$E_{g-i}(\theta,z) = \frac{2.5}{\ln 10}\left[\tau\Big(\theta,\lambda_g^{\rm rest}(z)\Big) - \tau\Big(\theta,\lambda_i^{\rm rest}(z)\Big)\right]$$

- Emission (optically thin approximation)

$$I(\lambda_{100\mu m},\theta,z) \quad = \quad \frac{1}{(1+z)^4}B\Big(\lambda_{100\mu m}^{\rm rest},T_{\rm d}(z)\Big)\tau(\lambda_{100\mu m}^{\rm rest}(z))$$

- Emission/absorption

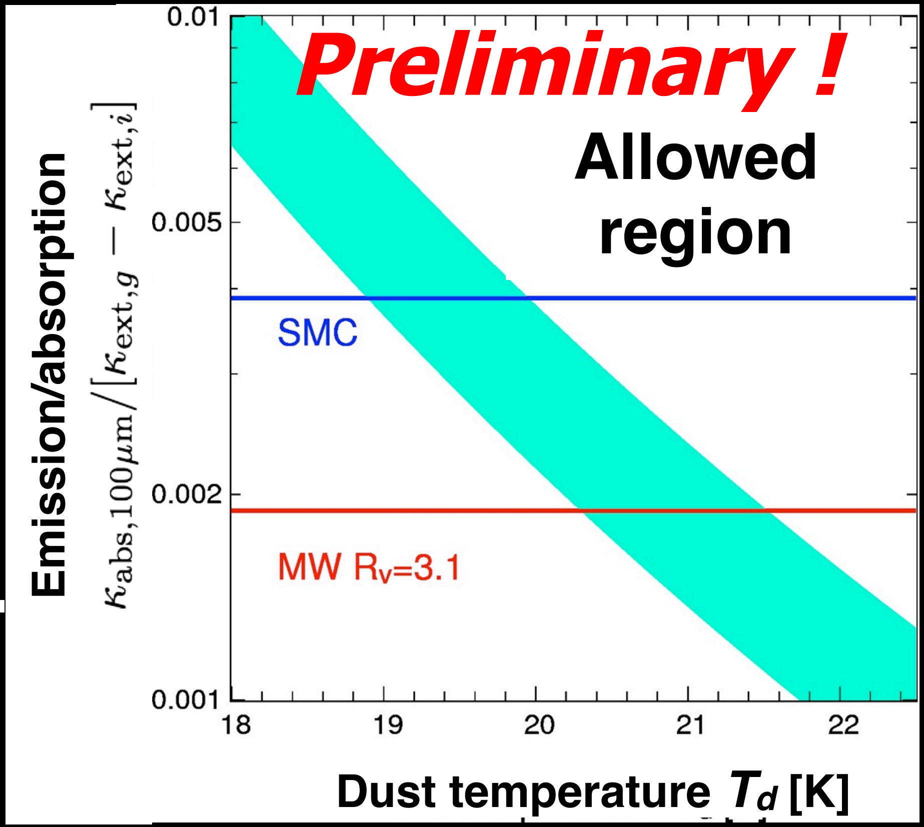
$$\frac{I(\lambda_{100\mu m},\theta,\bar{z})}{E_{g-i}(\theta,\bar{z})} = \frac{\ln 10}{2.5}\frac{1}{(1+\bar{z})^4}B(\lambda_{100\mu m}^{\rm rest}(\bar{z}),T_{\rm d}(\bar{z}))\frac{\kappa_{\rm abs}(\lambda_{100\mu m}^{\rm rest}(\bar{z}))}{\kappa_{\rm ext}\Big(\lambda_g^{\rm rest}(\bar{z})\Big) - \kappa_{\rm ext}\Big(\lambda_i^{\rm rest}(\bar{z})\Big)}.$$

constraining the dust temperature

- The ratio of emission and absorption compared with MW and SMC models (Weingartner & Draine 2001)

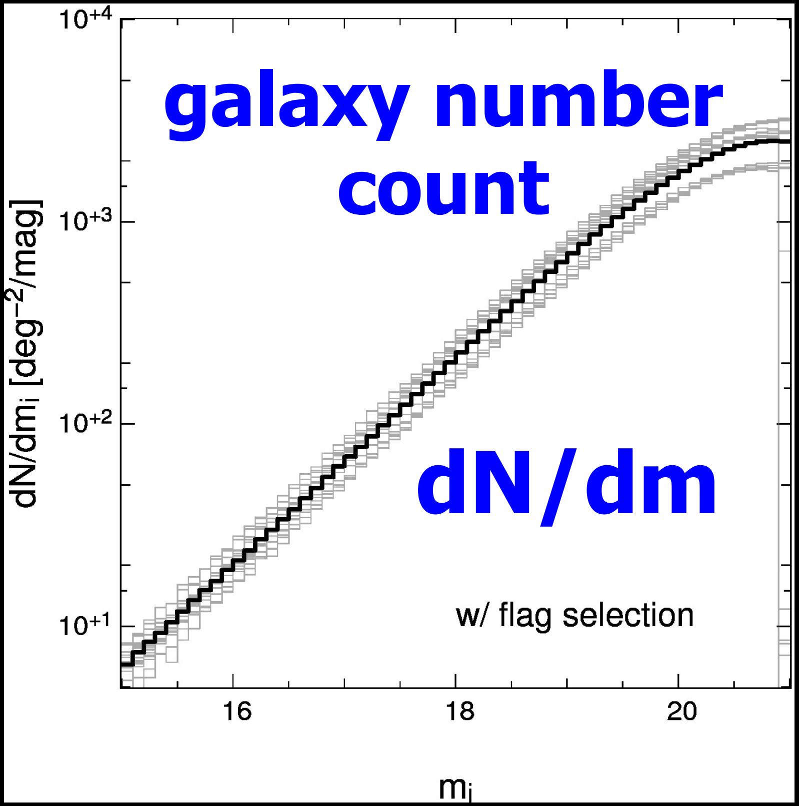
  http://www.astro.princeton.edu/~draine/dust/dustmix.html

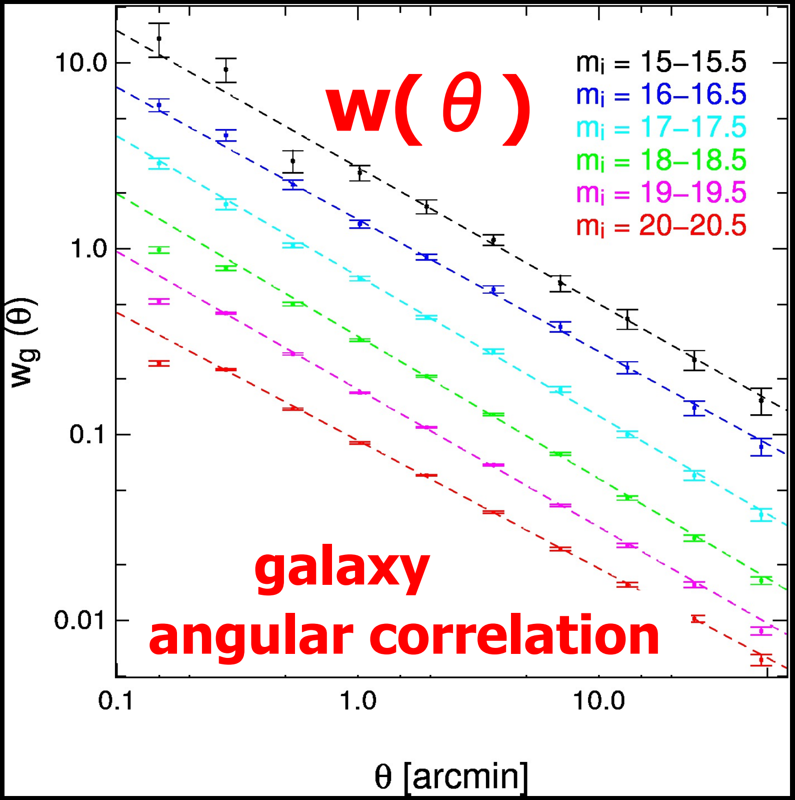
- **consistent with typical dust temperature of galaxies (~20K)**
  - **MSFR and we observed the absorption and emission of the same component, respectively.**
  - **Temperature of intergalactic dust ???**



Contribution of SDSS galaxies to the 100 μ m emission $I_{\mathrm{clustering}}(\theta, m_i)$

$$I_{\mathrm{clustering}}(\boldsymbol{\theta}, m_i) = \int dm' \frac{dN(m')}{dm'} \int d\boldsymbol{\varphi}\, I_{\mathrm{single}}(\boldsymbol{\theta} - \boldsymbol{\varphi}; m') w(\boldsymbol{\varphi}; m', m_i)$$

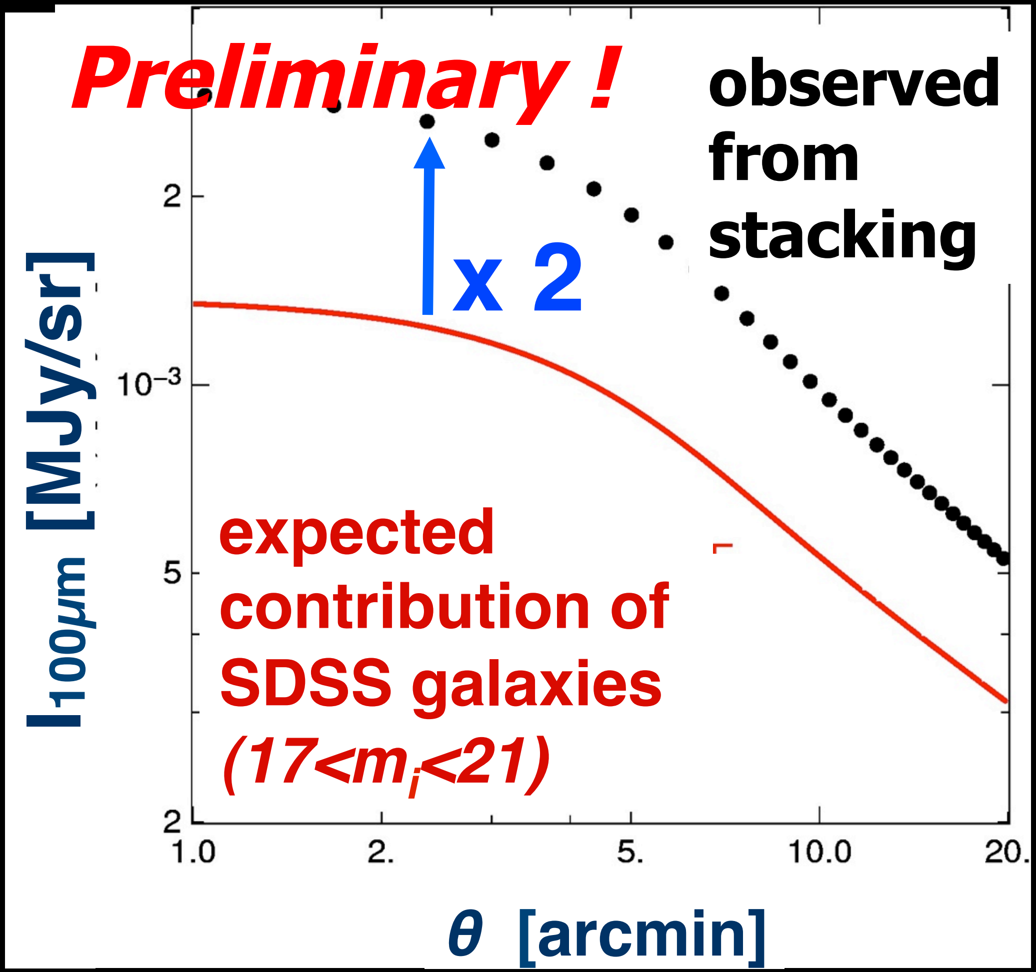




~50% is from SDSS galaxies

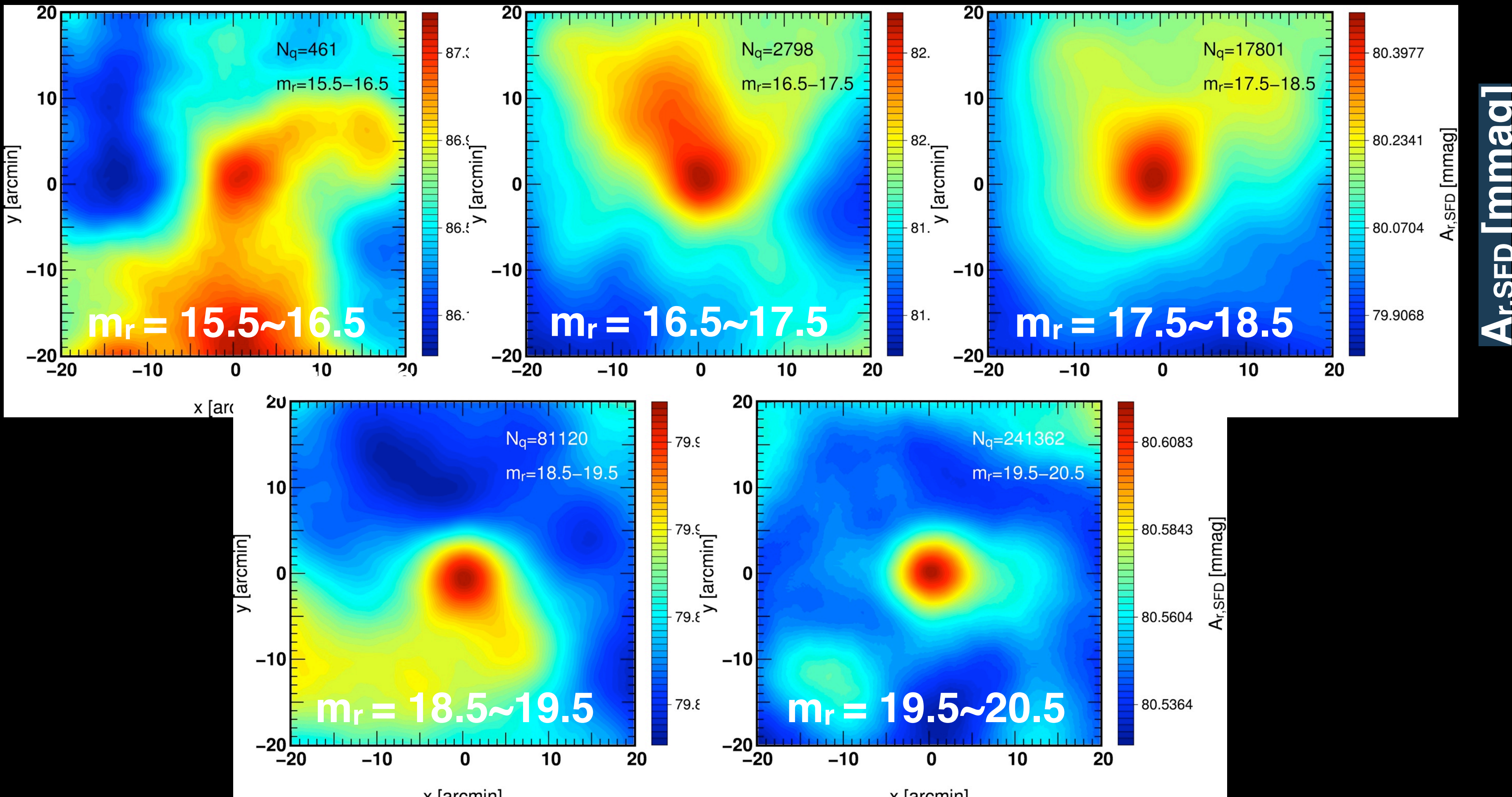
$$I_{\mathrm{clustering,MSFR}}(\theta) = \int_{17.0}^{21.0} I_{\mathrm{clustering}}(\theta; m'_i)\frac{dN}{dm'_i}dm'_i \times \left(\int_{17.0}^{21.0} \frac{dN}{dm'_i}dm'_i\right)^{-1}$$

- SDSS galaxies explain 50% of the FIR flux
- The rest may be
  - extended dust component over cluster scales ?
  - sums of optically faint infrared galaxies ?

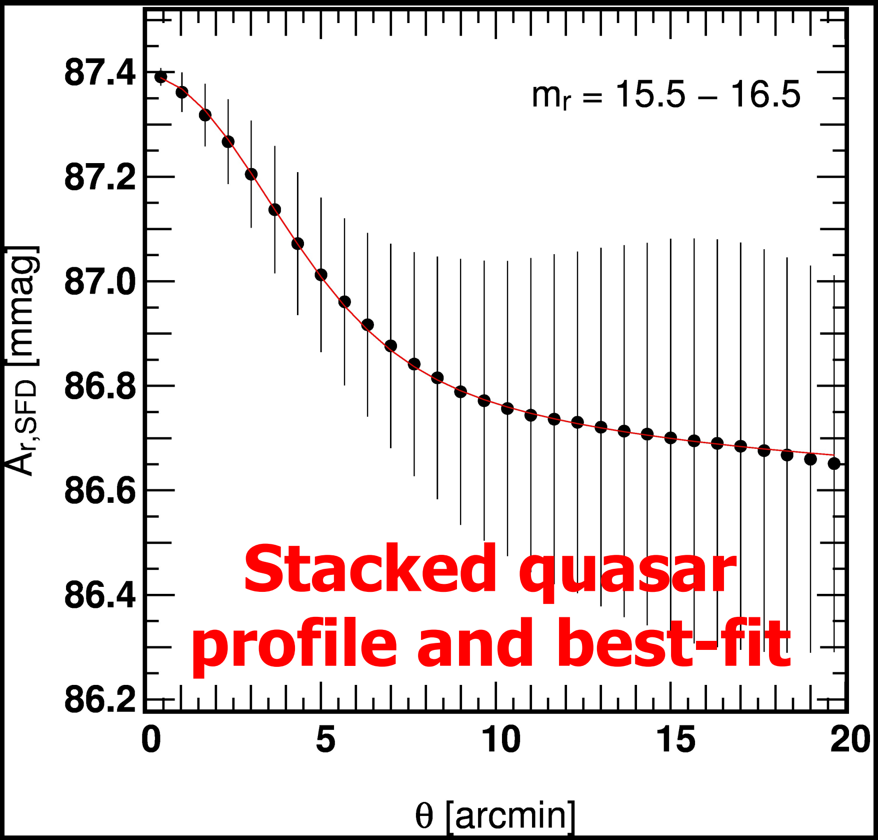


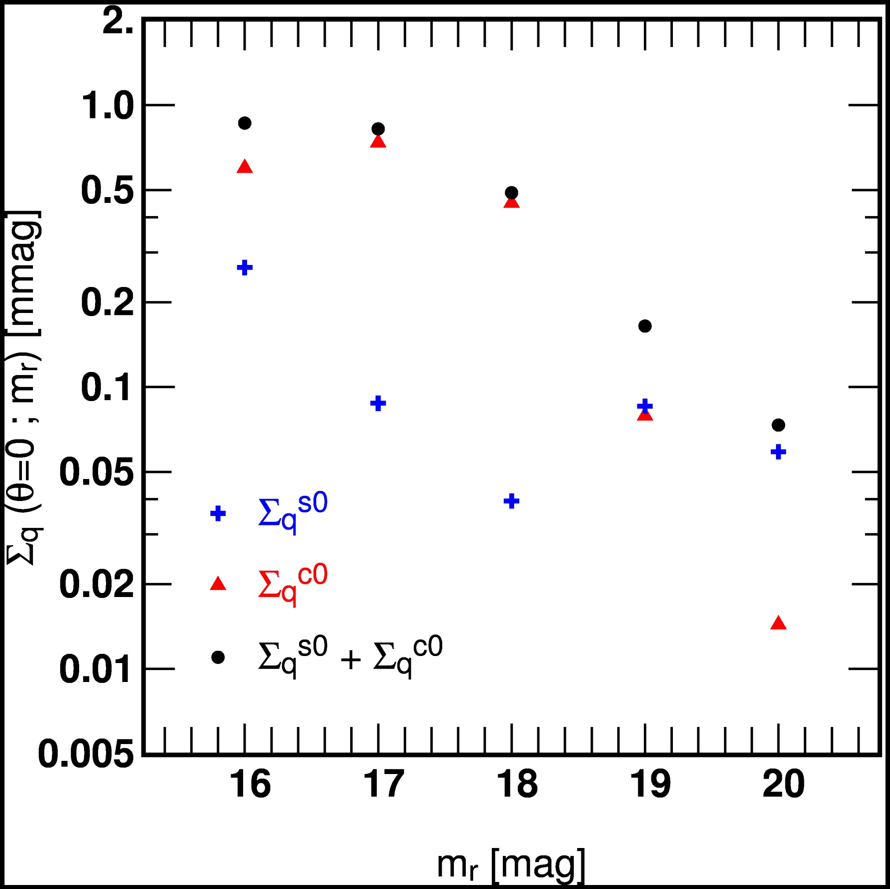
Stacking SDSS quasars

Similar stacking analysis of SDSS photometric quasar catalogue (Richards et al. 2009) indicates the FIR signals as well.



Profiles of stacked quasar images





Best-fit decomposition into the single and clustering terms

Quasars stacking should deserve further study…

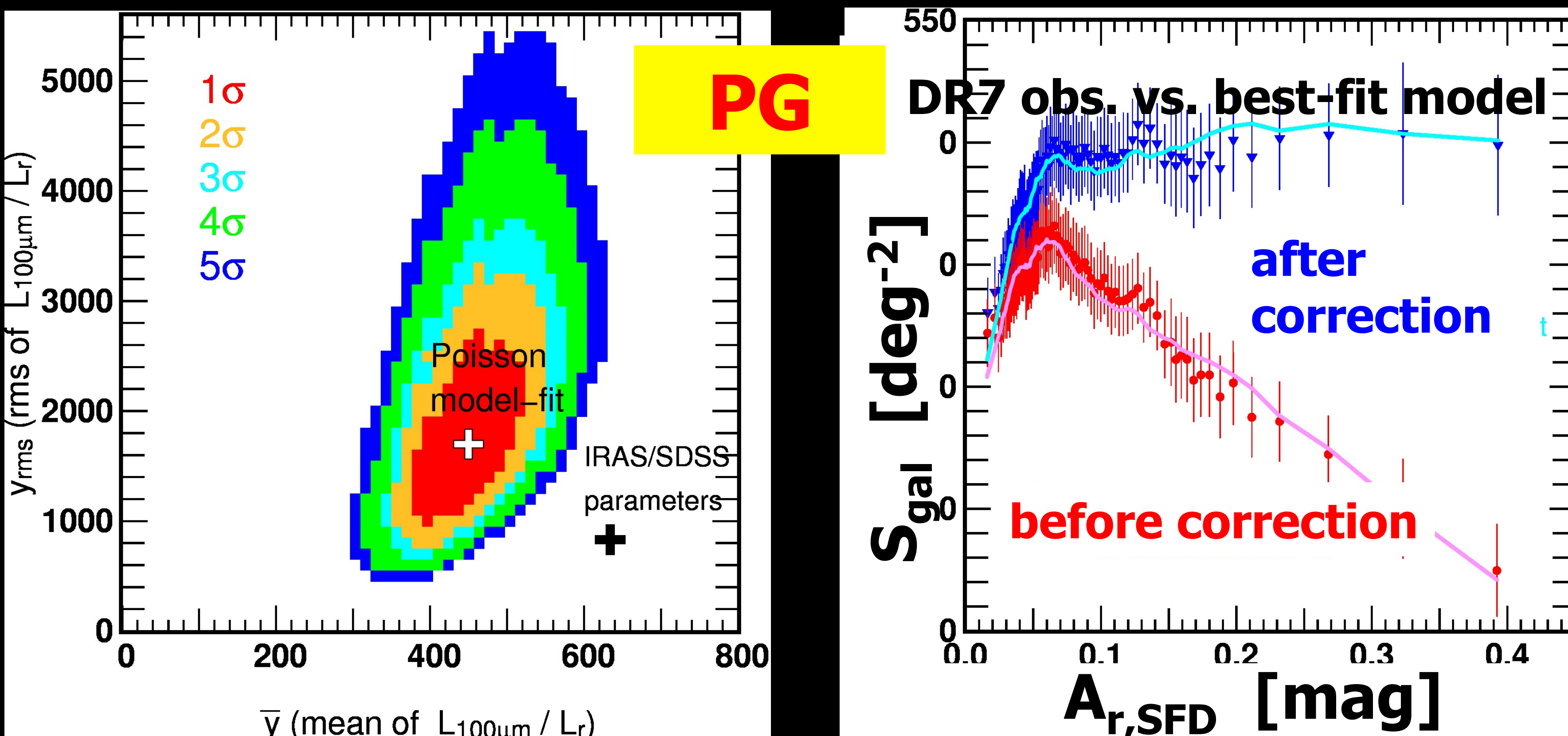
Summary of the first part

- Detection of FIR emission from SDSS galaxies by stacking image analysis over the SFD map (～ IRAS 100 μm map)
  - Largely explains the anomaly of SDSS galaxy number counts as a function of A_{SFD} discovered by Yahata et al.(2007)
  - Possible correction to the SFD map and a future Galactic extinction map with Planck
  - A new probe of unresolved (dusty) galaxy correlations and/or dust profile of the hosting halo

Preliminary Summary of the second part

- Sum of dust associated with SDSS galaxies explains ~50% of the amount discovered via absorption by Ménard et al.(2010)
- Where is the remaining 50% ?
  - Associated with non-SDSS galaxies ?
  - Not directly associated with individual galaxies, but extended over cluster scales ?
- FIR emission from SDSS quasars detected, should be explored in future.

Fit to the observed anomaly using an analytic approximation model



- Corrected map by Peek & Graves (2010)
  - the anomaly still exists but agrees better with our model prediction than the original SFD map