

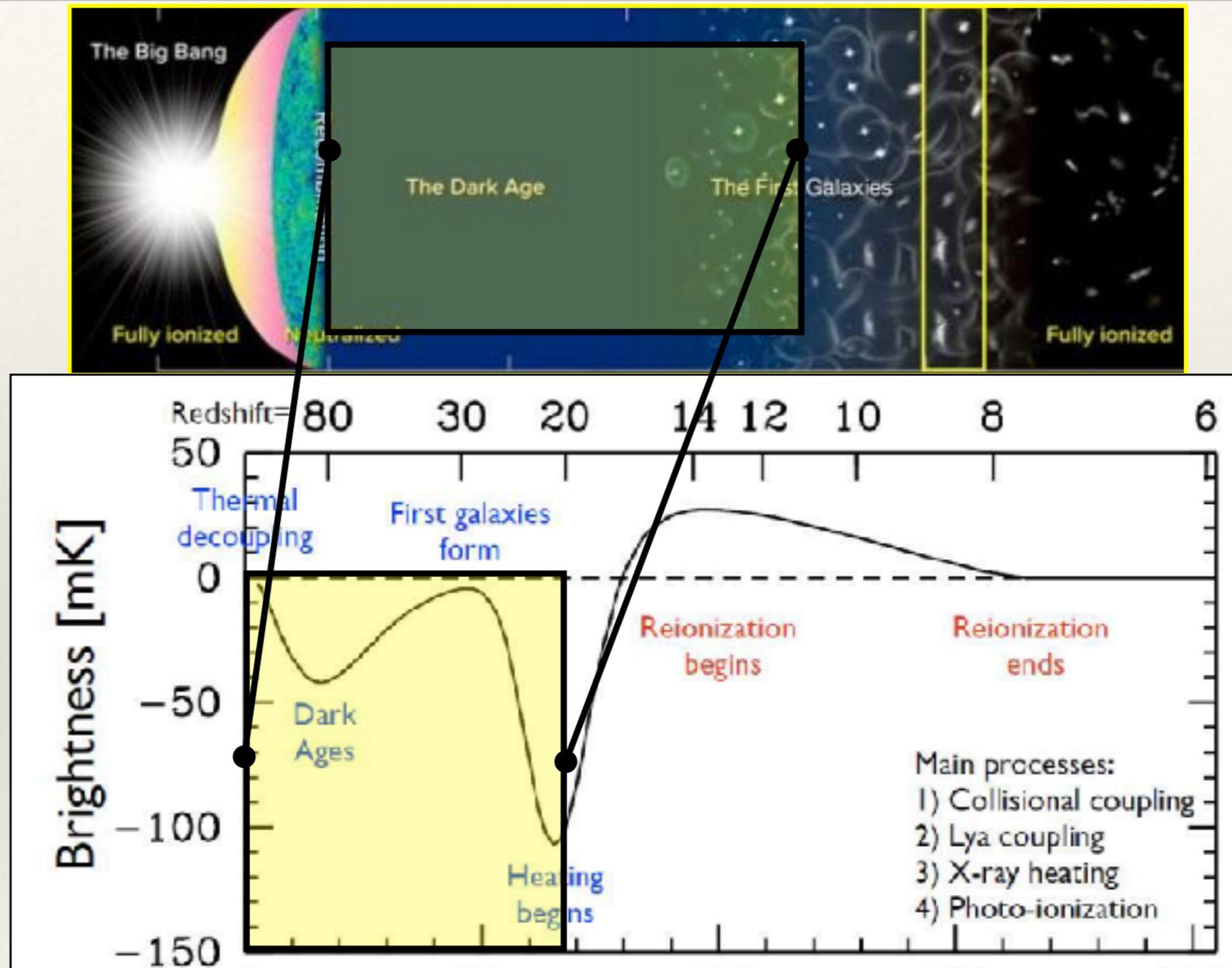
Hongo 21cm workshop (Oct 4)

**Modeling the 21 cm Signal from
the Cosmic Dark Ages:
Achieving Sub-Percent Accuracy**

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(Tsukuba University)

Collaborators: Rennan Barkana, Naoki Yoshida and so on

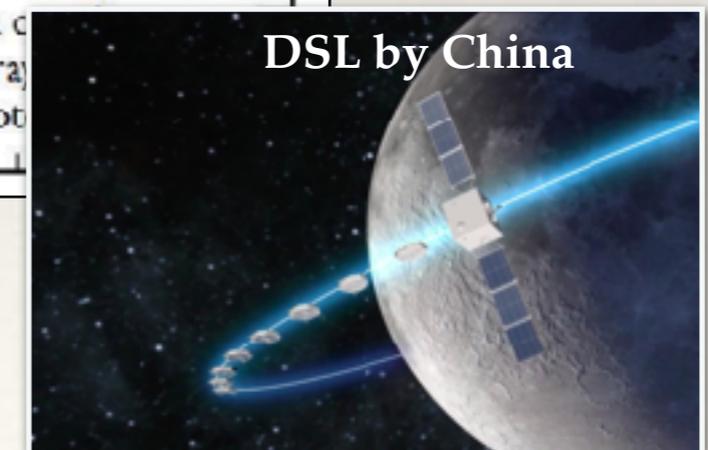
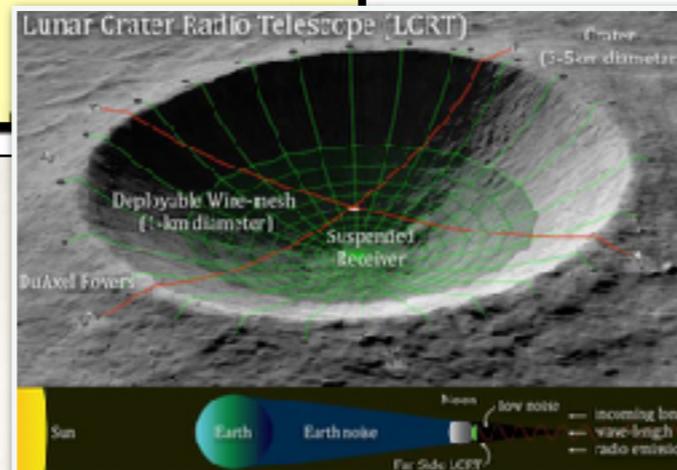
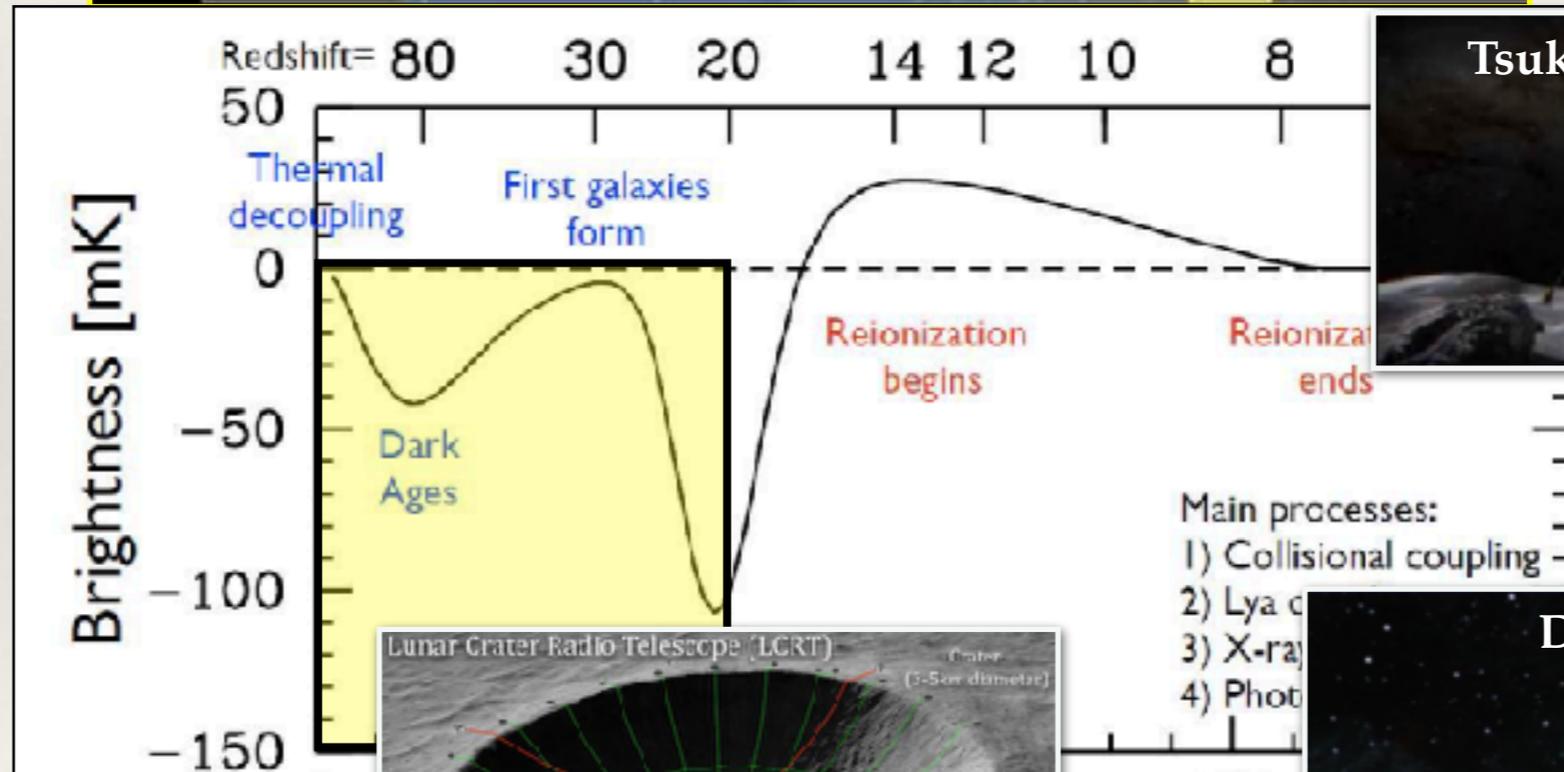
21cm Signal from $z > 20$



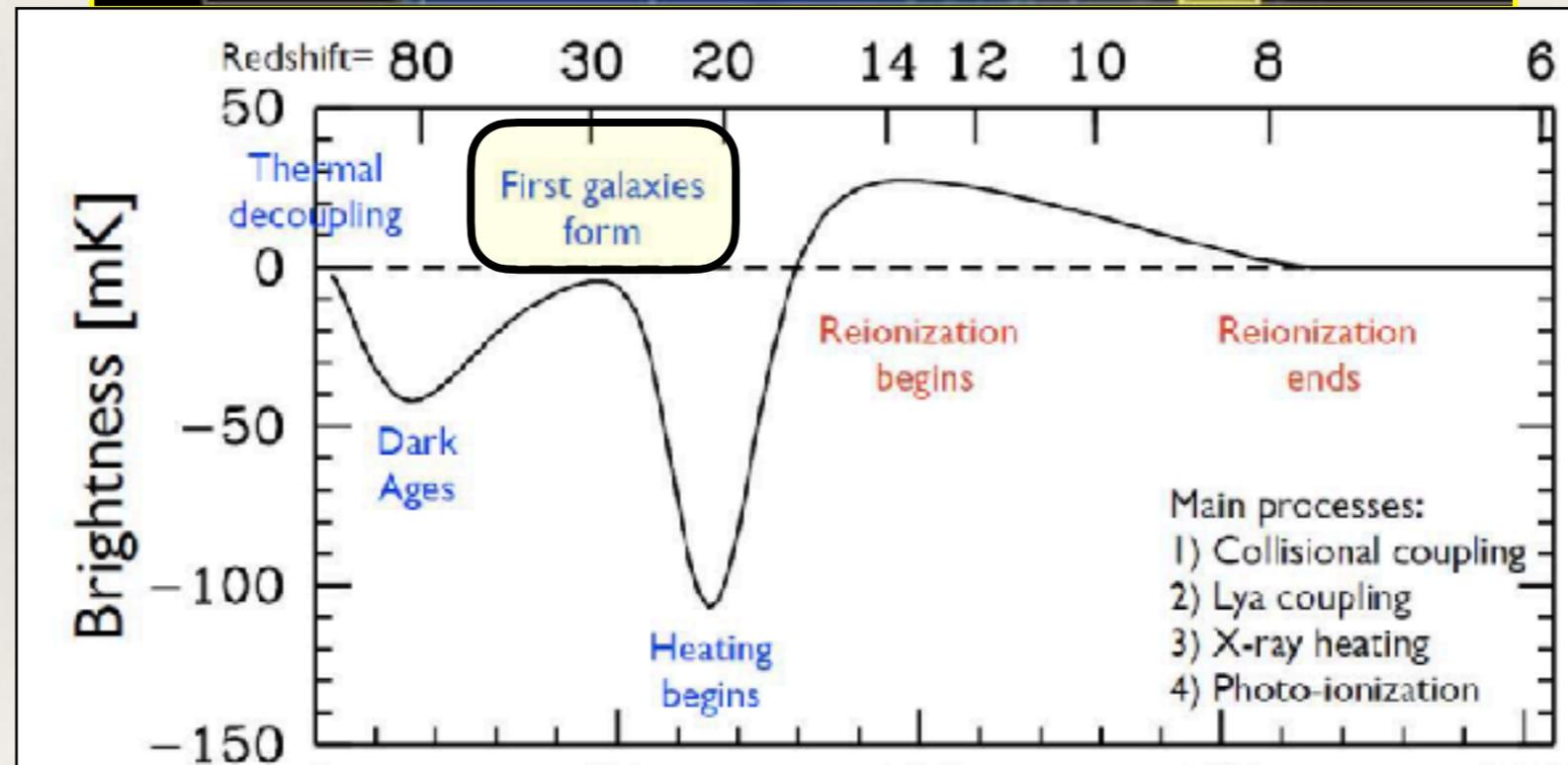
... comes from pristine gas before stars and galaxies ever formed and can constrain cosmology parameters without astrophysics uncertainties.

(Mondal & Barkana 2021)

21cm Signal from $z > 20$



21cm Signal from $z > 20$



The signal is often calculated for the homogeneous Universe, but structures are forming at sub-Mpc scales. We need to model the impact on the signal.

Impact of Structure Formation

$$T_b = (1 + z)^{-1} (T_S - T_{\text{CMB}}) (1 - e^{-\tau})$$

$$T_S^{-1} = \frac{T_{\text{CMB}}^{-1} + x T_{\text{gas}}^{-1}}{1 + x}$$

Coupling coefficient

$$x = x_\alpha + x_c$$

$$x_c \propto \kappa_{1-0}(T_{\text{gas}}) n_{\text{H}}$$

$$(1 - e^{-\tau}) \approx \tau \text{ for small } \tau$$

$$\approx 1 \text{ for large } \tau$$

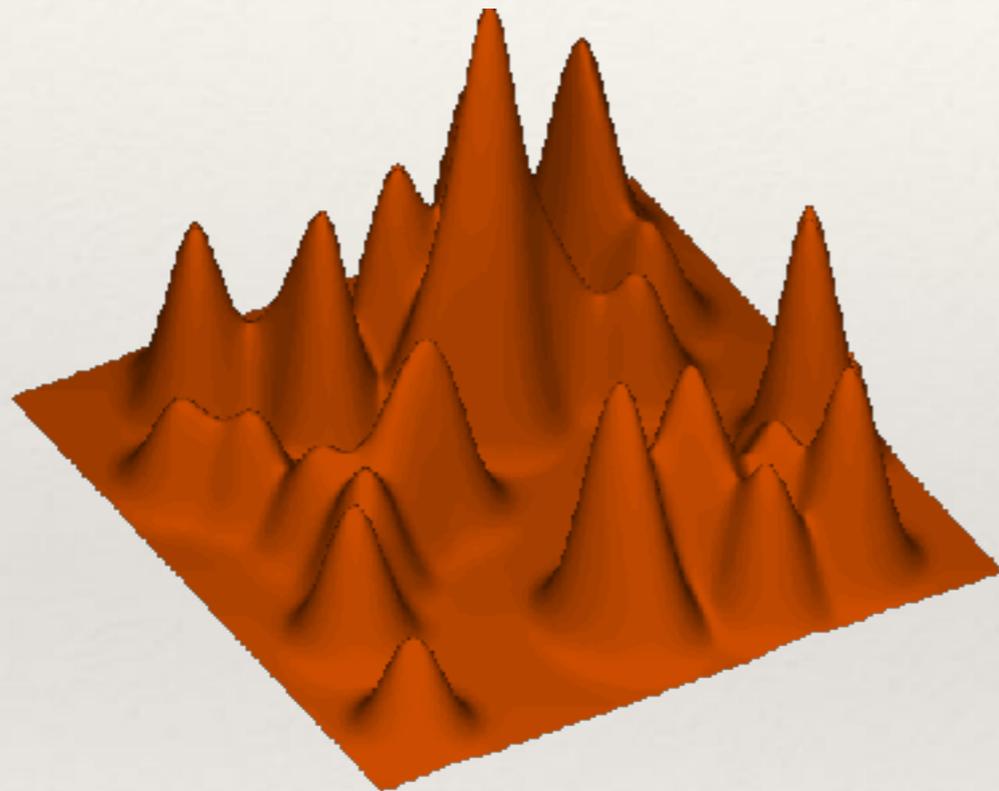
$$\tau(z) = \frac{3c\lambda_{21}h_p A_{10}n_{\text{HI}}}{32\pi k_{\text{B}}T_S(1+z)(dv_r/dr)} \propto \frac{n_{\text{HI}}}{dv_r/dr}$$

$H(z)$ in the homogeneous Universe

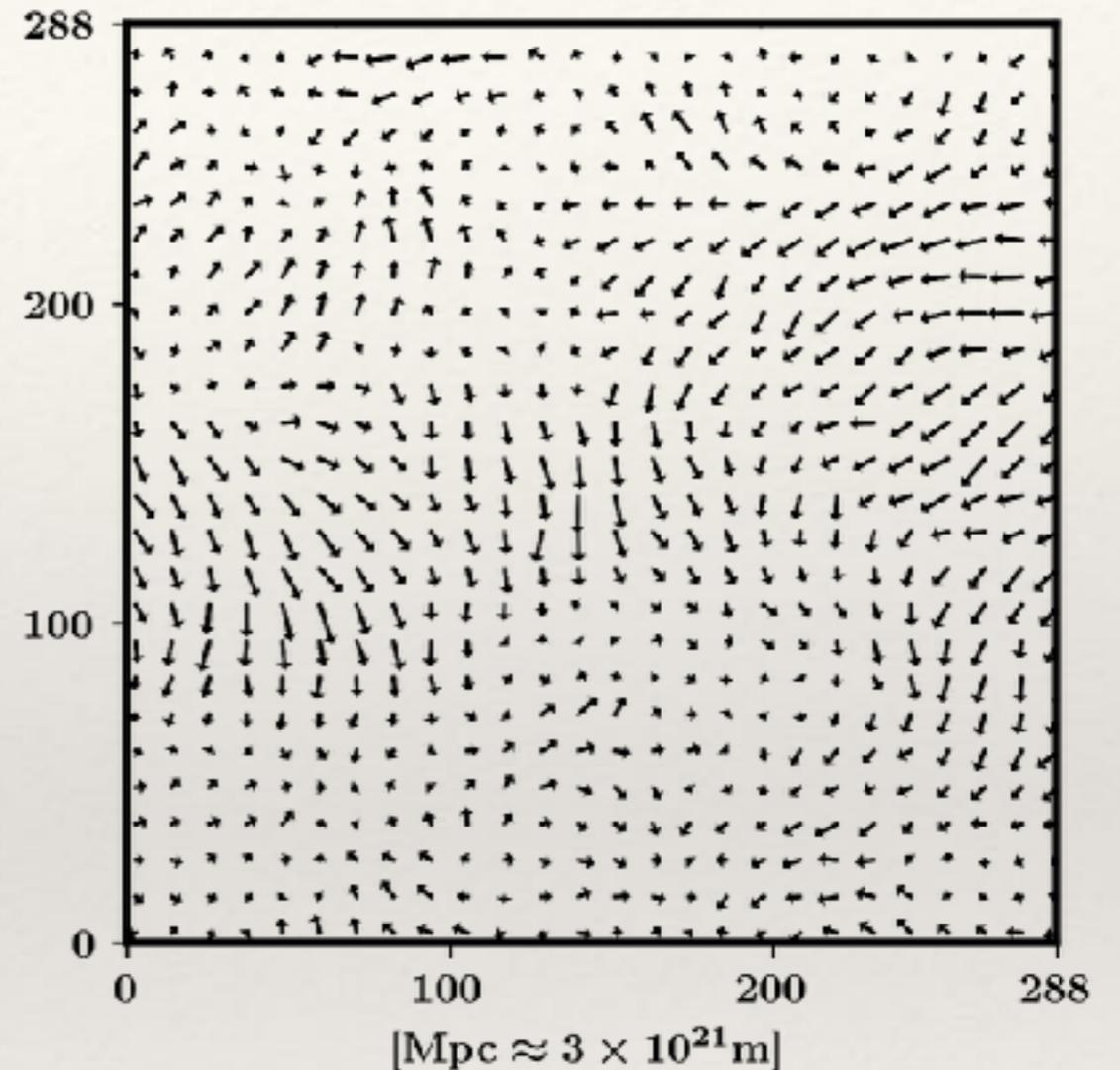
Structure formation introduces fluctuations in T_{gas} and n_{H} , and v_r
How will that affect T_b ?

Baryon-DM Streaming

At $z > 1090$,

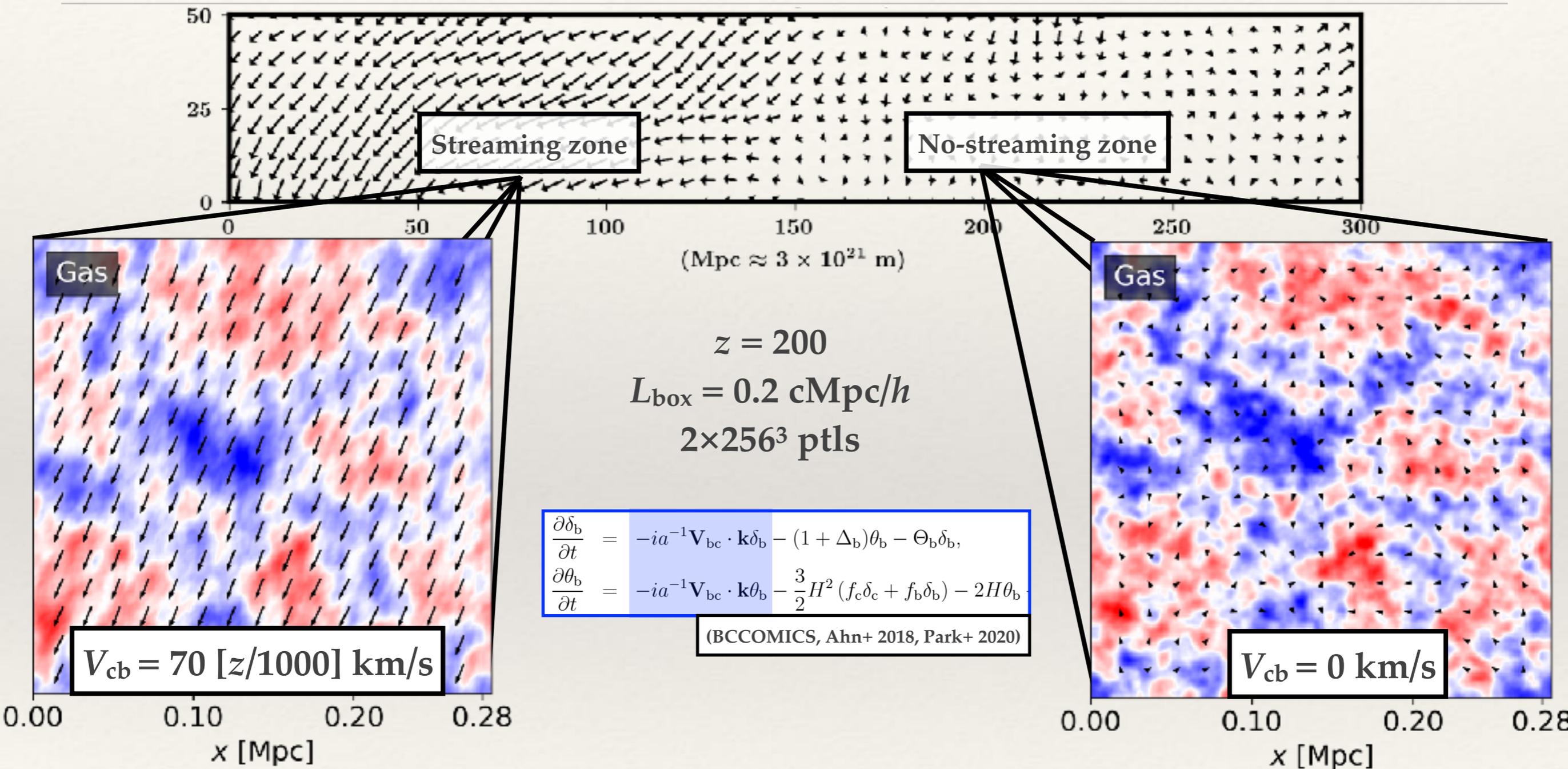


... pressure of photon-baryon fluid generates the Baryonic Acoustic Oscillations.
This sound wave became a *relative motion* between baryons and dark matter.

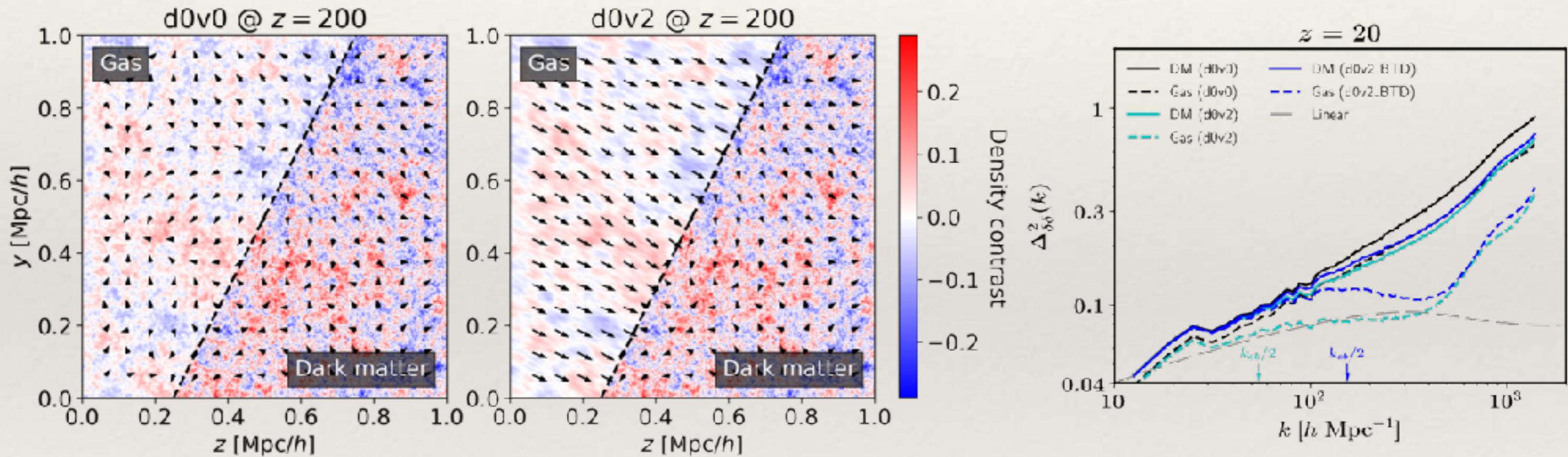


$V_{cb} = 0 - 70 [z/1000] \text{ (km/s)}$
Not strong enough to affect SF at $z \sim 6$
But, strong enough disrupt some of
small-scale gas structures

Initial Conditions with Streaming v

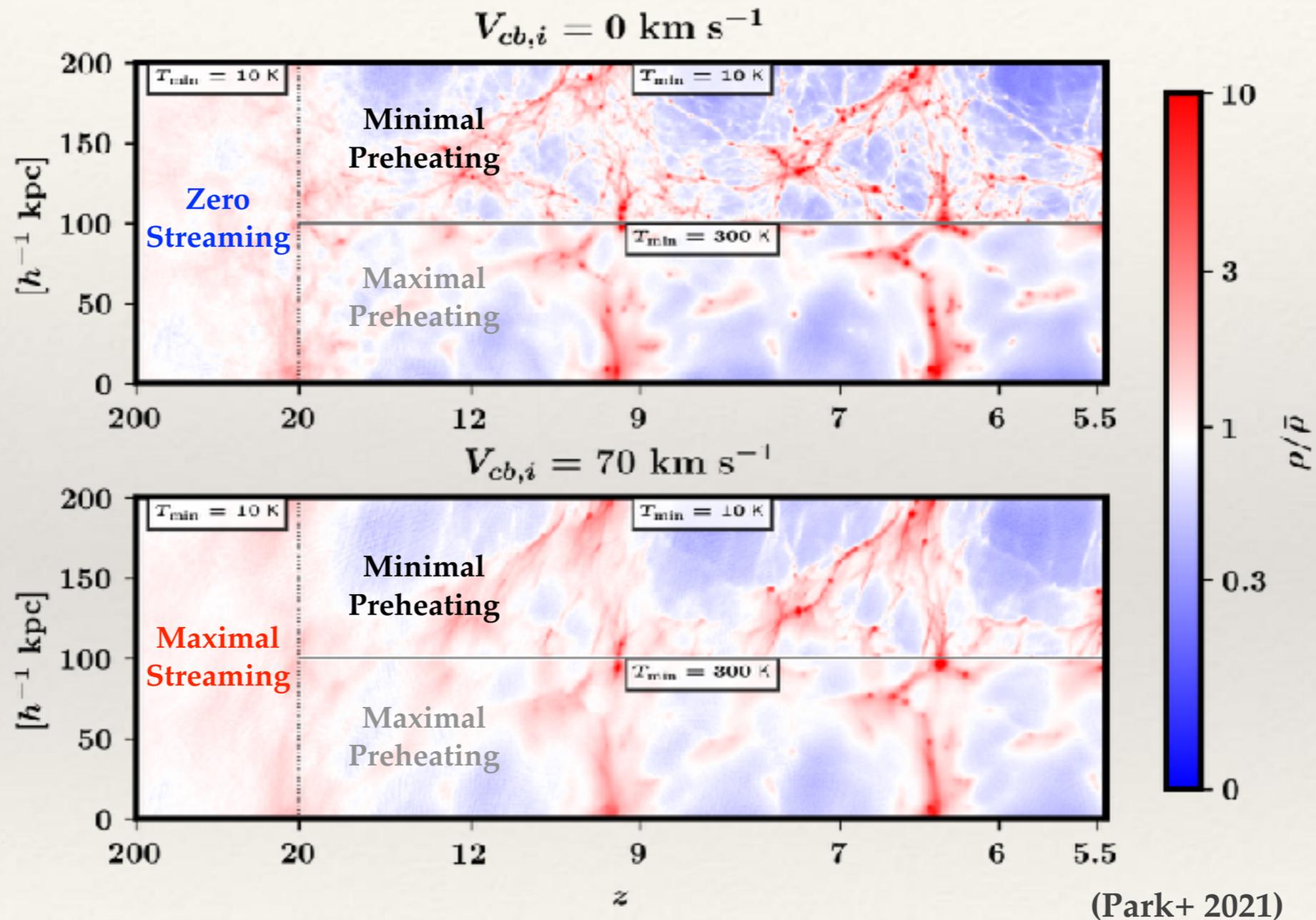


Baryon-dark Matter Streaming Velocity



(Park+ 2020)

Pre-reionization IGM with Streaming & Preheating

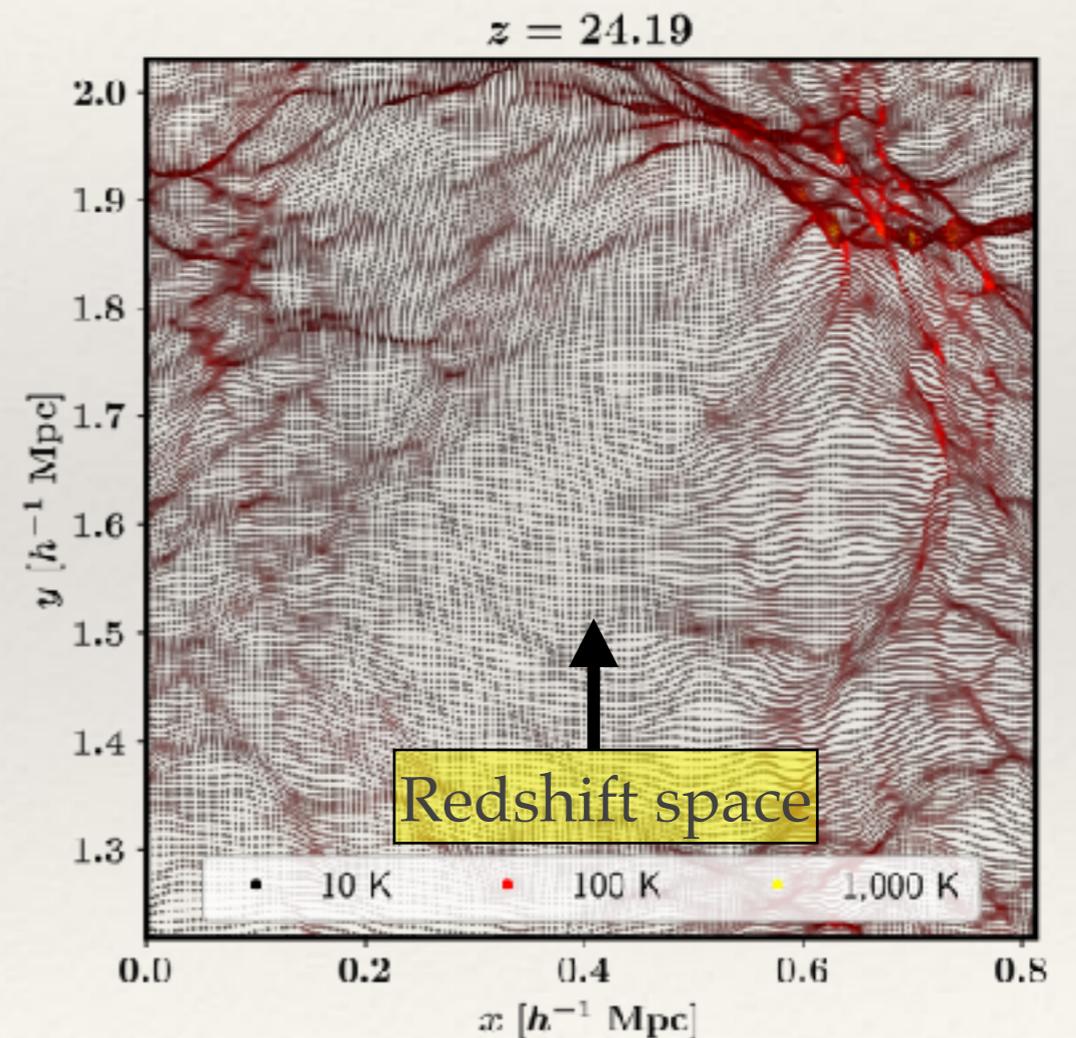
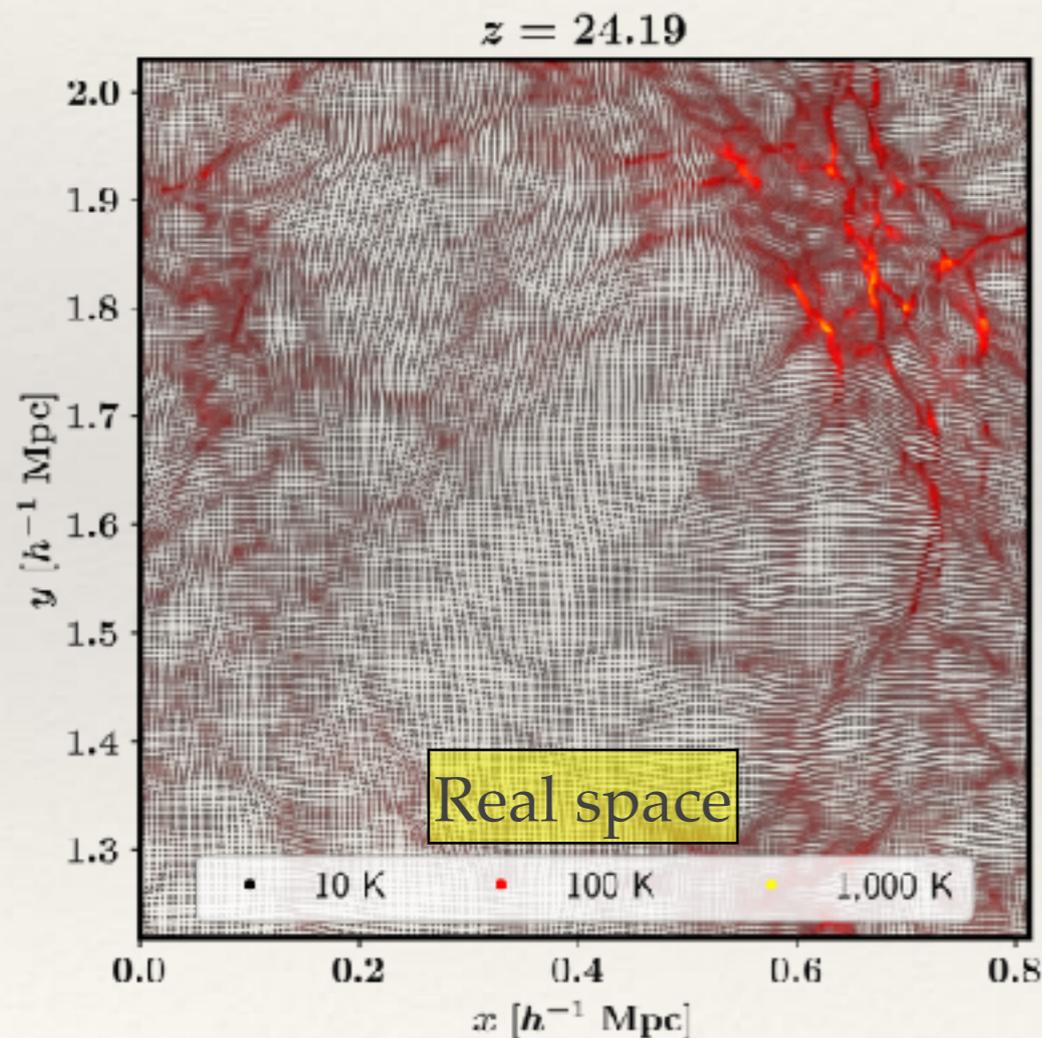


Streaming v & X-ray preheating erases small-scale gas structures.

Redshift Space Distortion

$$T_b = (1 + z)^{-1} (T_S - T_{\text{CMB}}) (1 - e^{-\tau})$$

$$\tau(z) = \frac{3c\lambda_{21}h_p A_{10}n_{\text{HI}}}{32\pi k_B T_S (1+z)(dv_r/dr)} \propto \frac{n_{\text{HI}}}{dv_r/dr} = n_{\text{HI,RS}} \quad \begin{array}{l} (1 - e^{-\tau}) \approx \tau \text{ for small } \tau \\ \approx 1 \text{ for large } \tau \end{array}$$

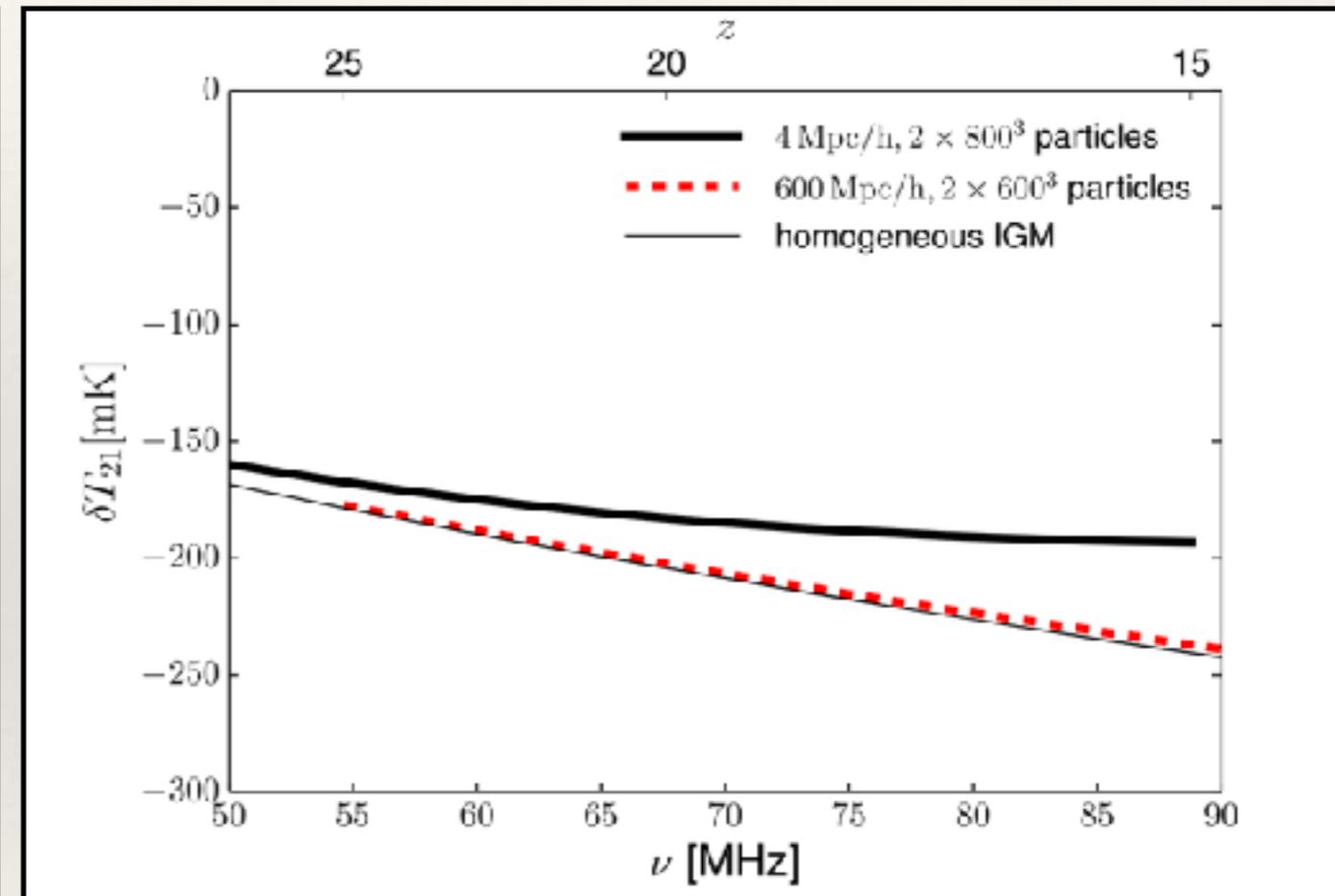
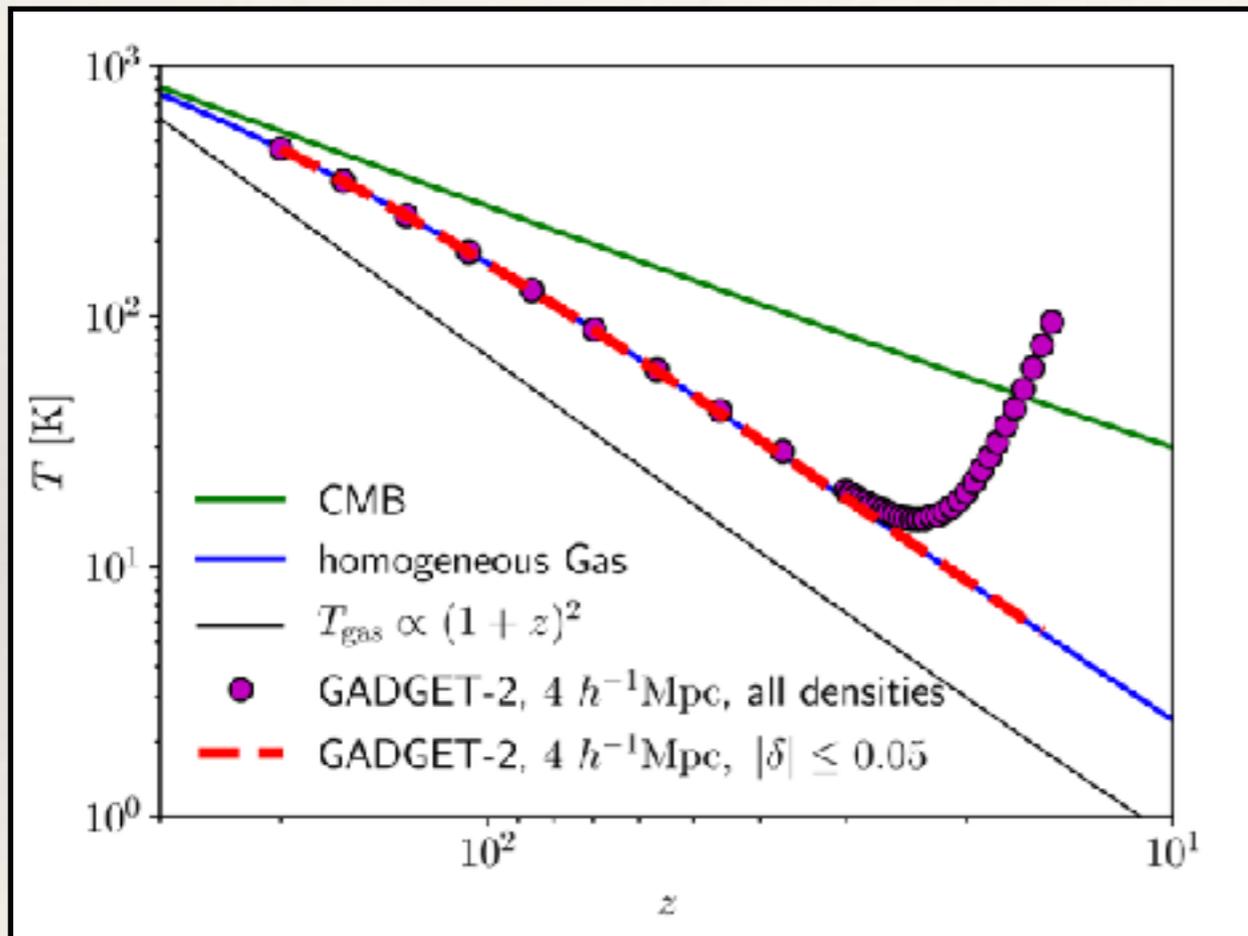


The RSD effect can reduce the signal from dense gas.

Impact on the Maximal 21 cm Signal at $z=17$

(Xu+ 2021)

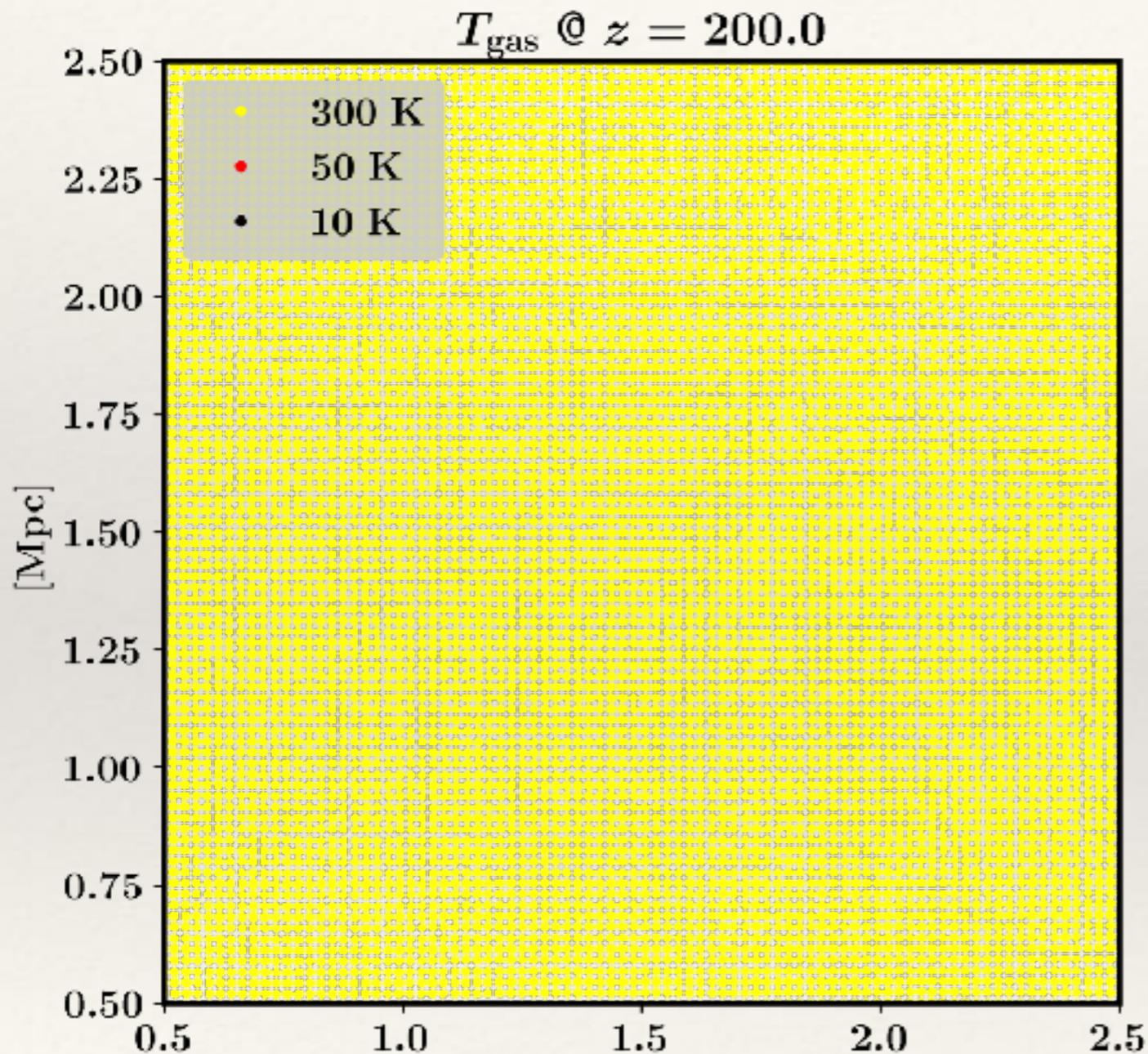
Maximal signal: $x_a = \infty \rightarrow T_S = T_{\text{gas}}$



Gas heated by compression and shocks
from structure formation

Weaker maximum absorption due to
the heating

Simulation Setup & Gas Temperature



Code: GADGET-2

Box size: 3 Mpc

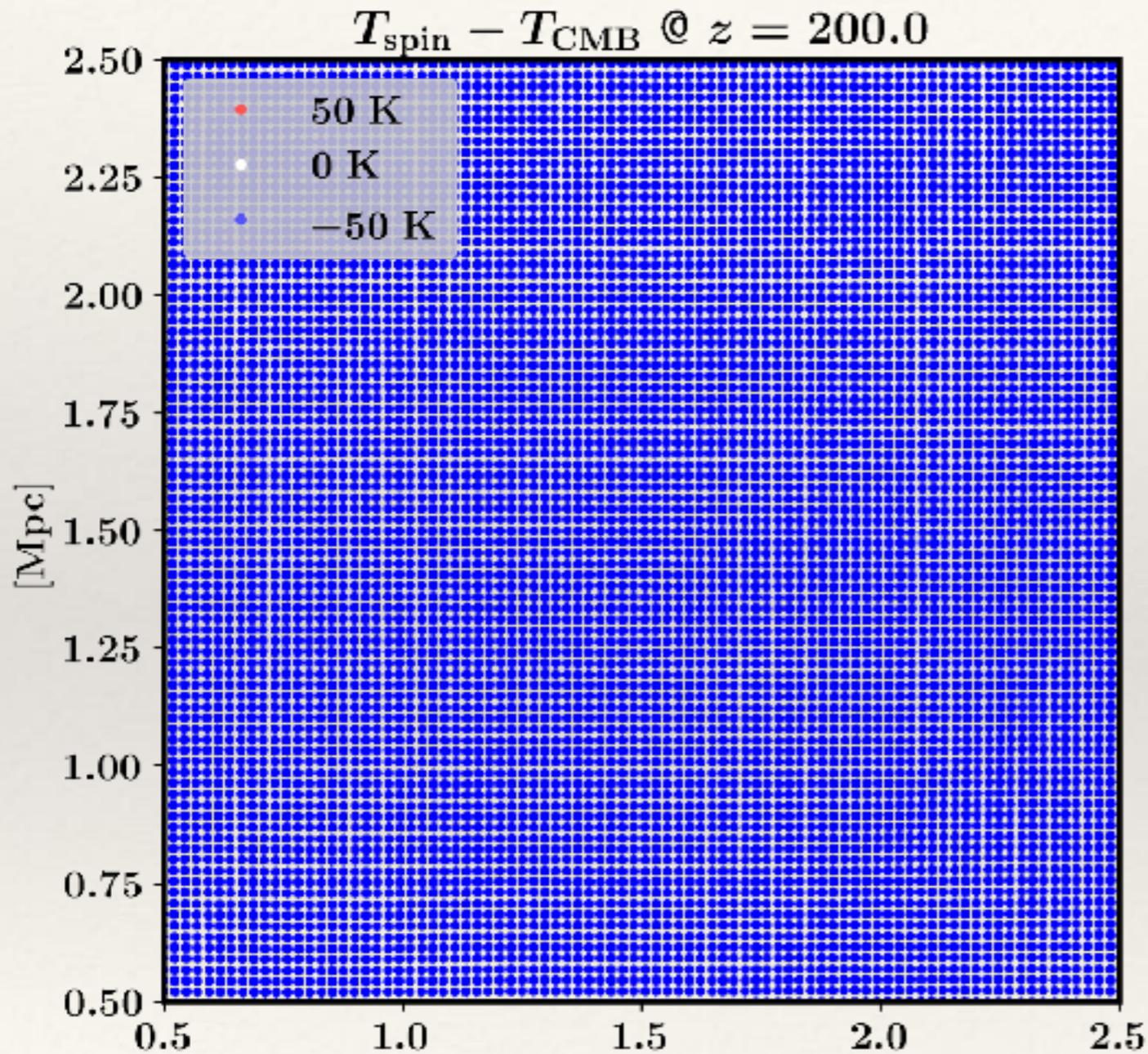
Resolution: 512^3 ptls

Chemical processes

- Compton heating
- $e + \text{H}^+ \rightarrow \text{H}$
- $e + \text{He}^+ \rightarrow \text{He}$
- Cooling by H and H_2 disabled
- No star formation

Our goal is to model the dark age signal from $z > 20$, assuming $\kappa_a = 0$.

Spin Temperature



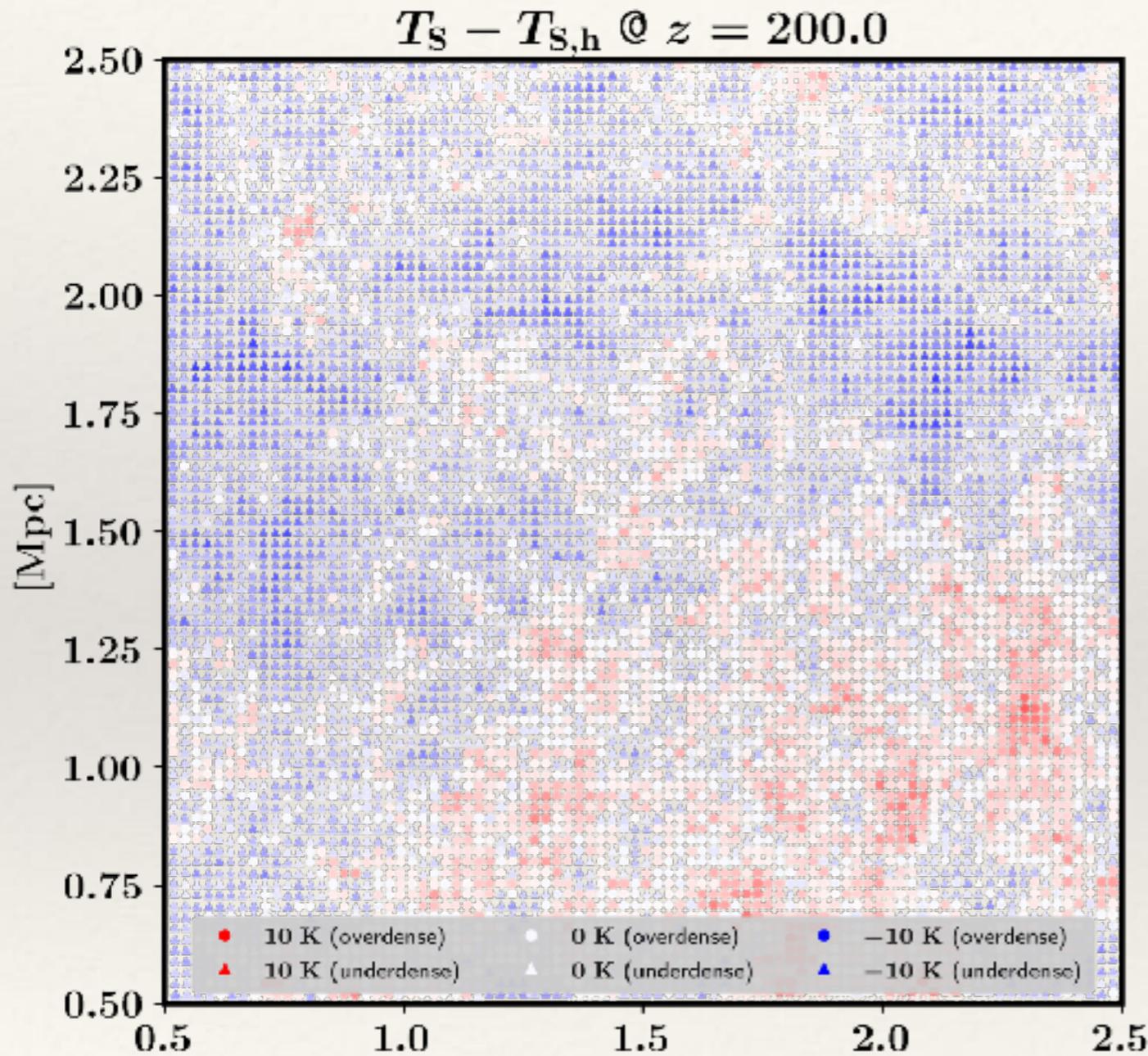
$$T_S^{-1} = \frac{T_{\text{CMB}}^{-1} + x_c T_{\text{gas}}^{-1}}{1 + x_c}$$

where

$$x_c = \frac{\kappa_{1-0}(T_{\text{gas}}) n_{\text{HI}} T_*}{A_{10} T_{\text{CMB}}}$$

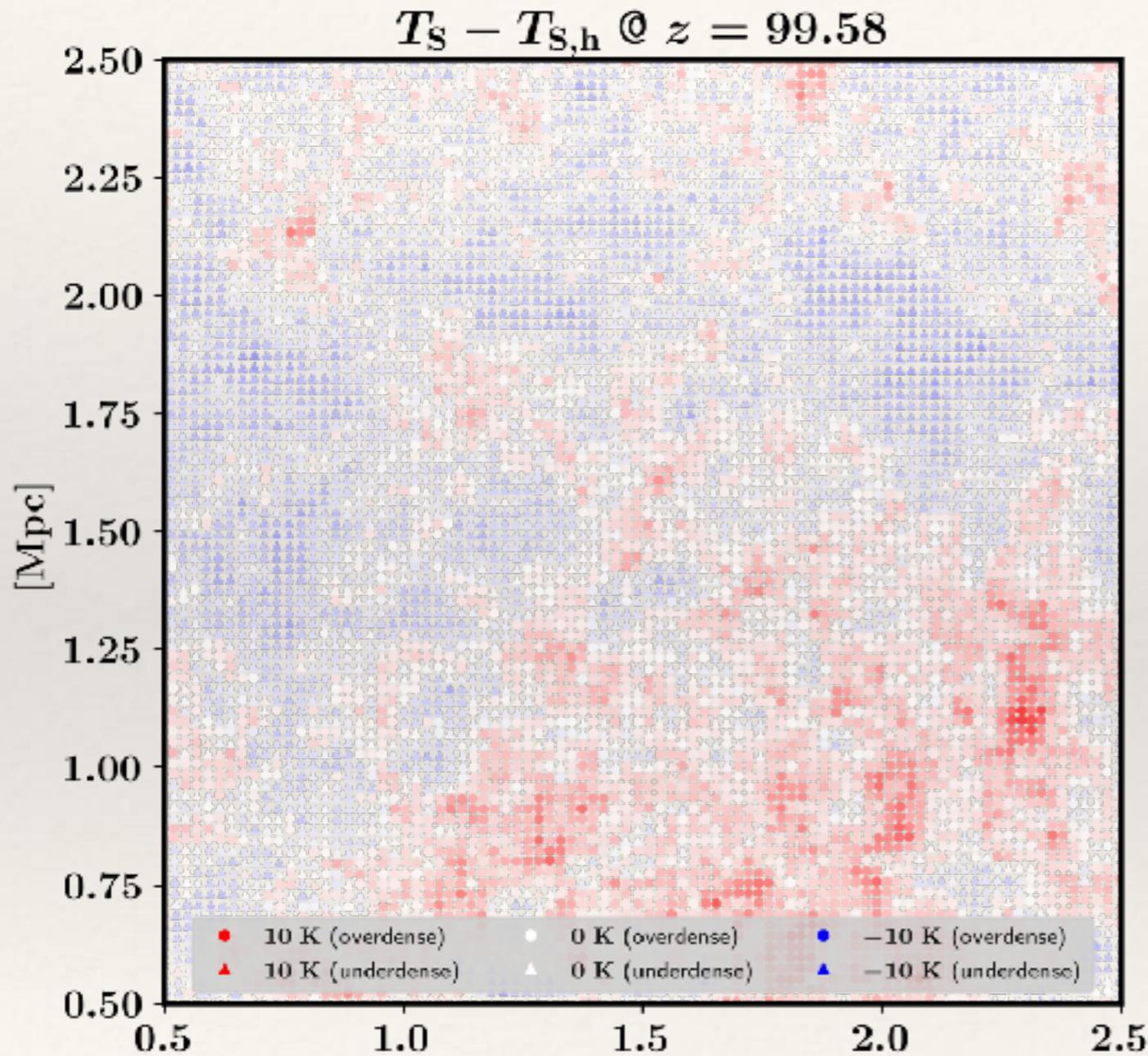
Strong coupling when $n_{\text{HI}} > 0.1 \text{ cm}^{-3}$
Weak coupling when $n_{\text{HI}} < 0.1 \text{ cm}^{-3}$

Impact of Structure Formation on Spin Temperature

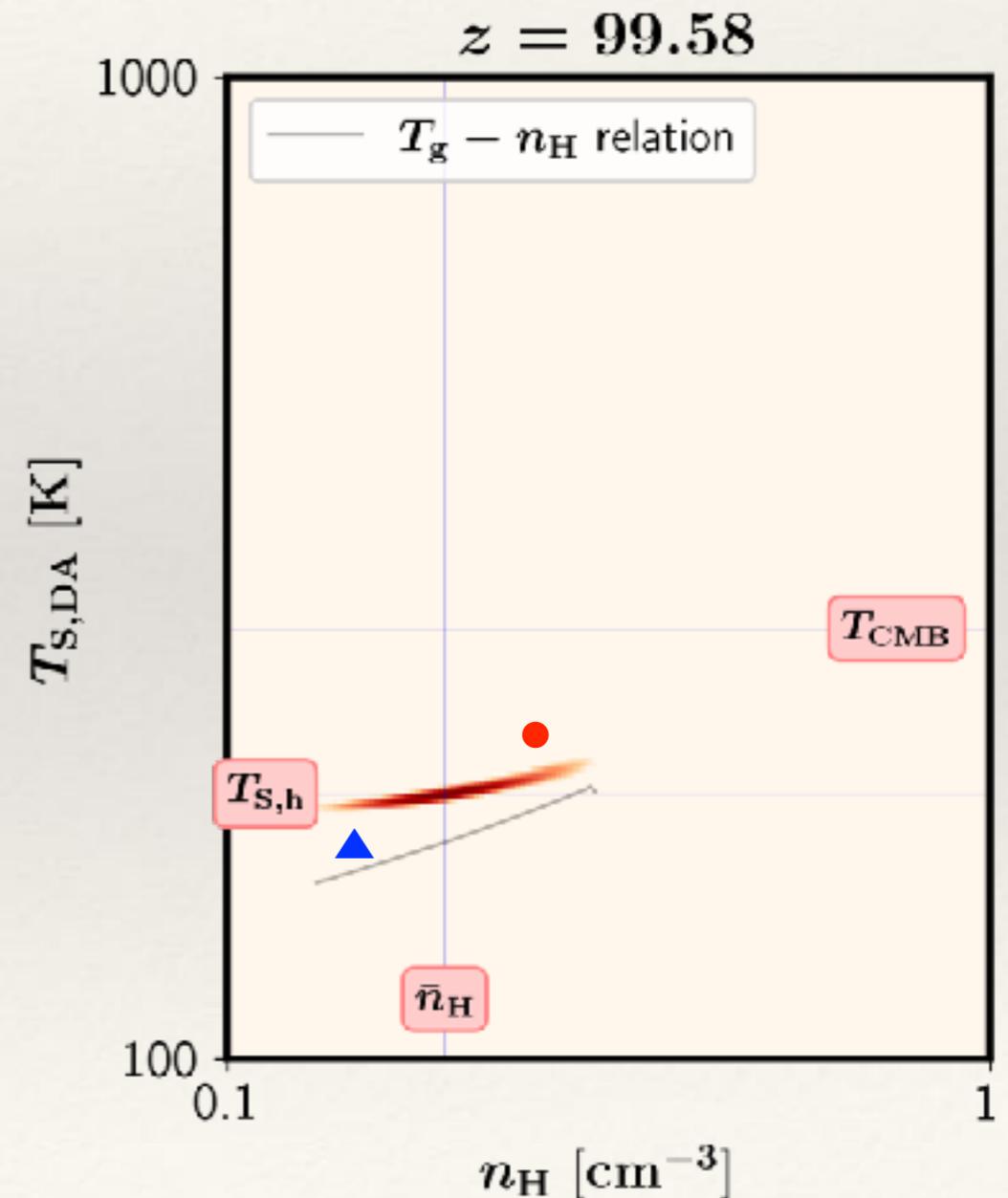


- Positive relation btw ΔT_S and δ at $z > 70$ (▲ & ●)
- Negative relation btw ΔT_S and δ at $z < 60$ (▲ & ●)
- $\Delta T_S = 0$ in underdense regions (◻) &
 $\Delta T_S < 0$ or $\Delta T_S > 0$ in overdense regions at $z < 30$ (● & ●)

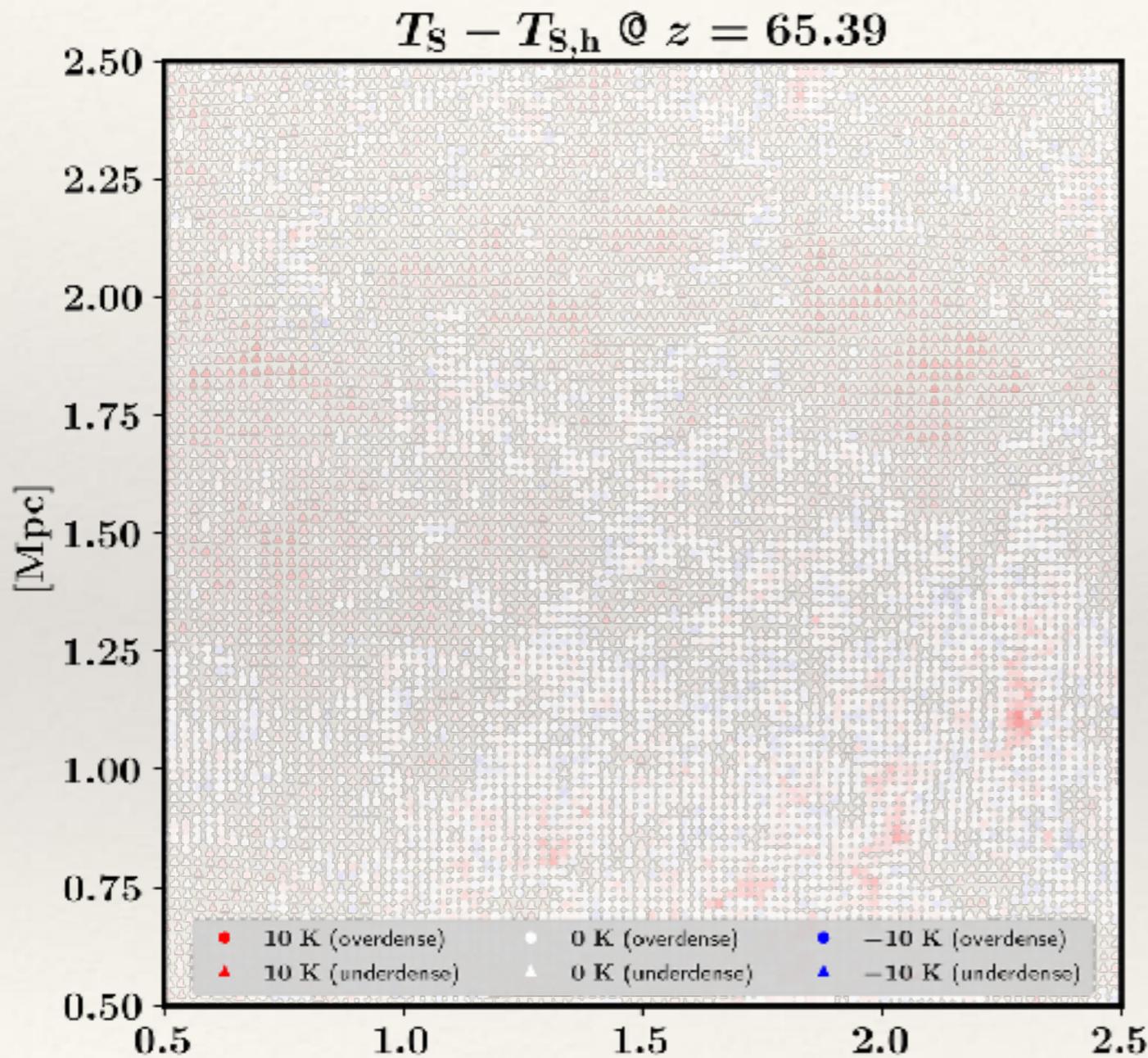
ΔT_S due to Structure Formation ($z > 70$)



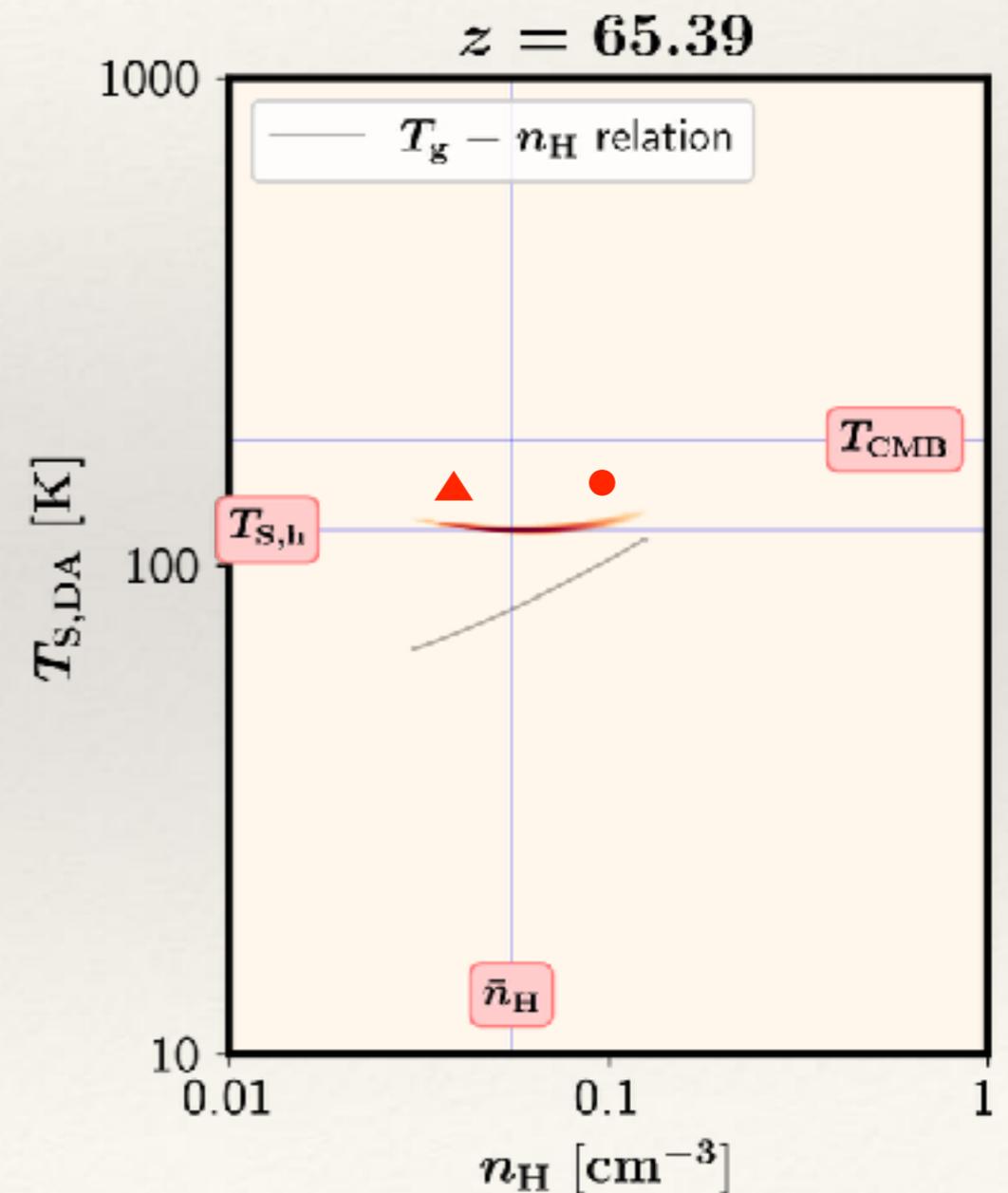
Positive relation btw ΔT_S and δ at $z > 70$
as T_S is tightly coupled to T_{gas} (▲ & ●)



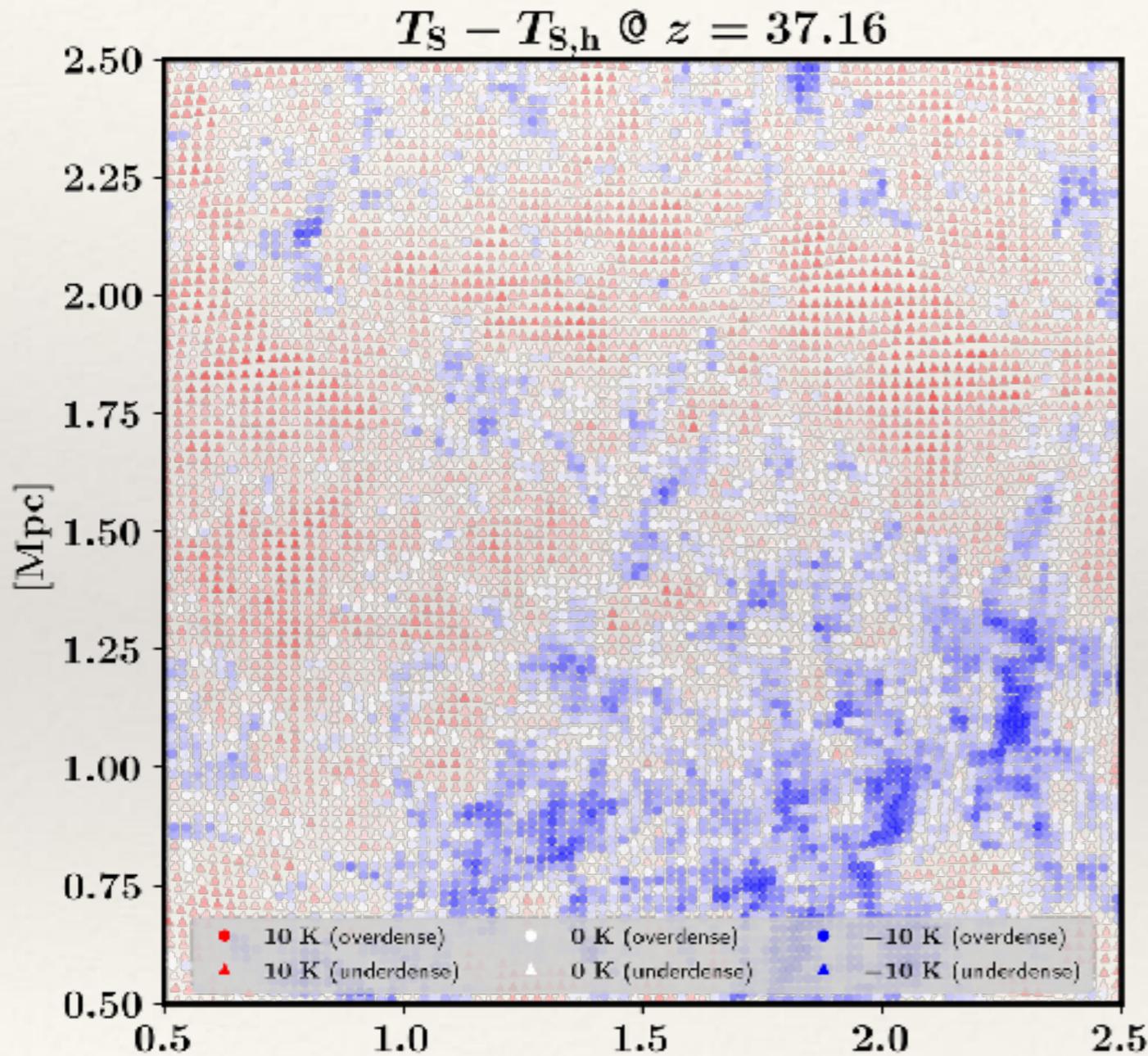
ΔT_S due to Structure Formation ($z = 65$)



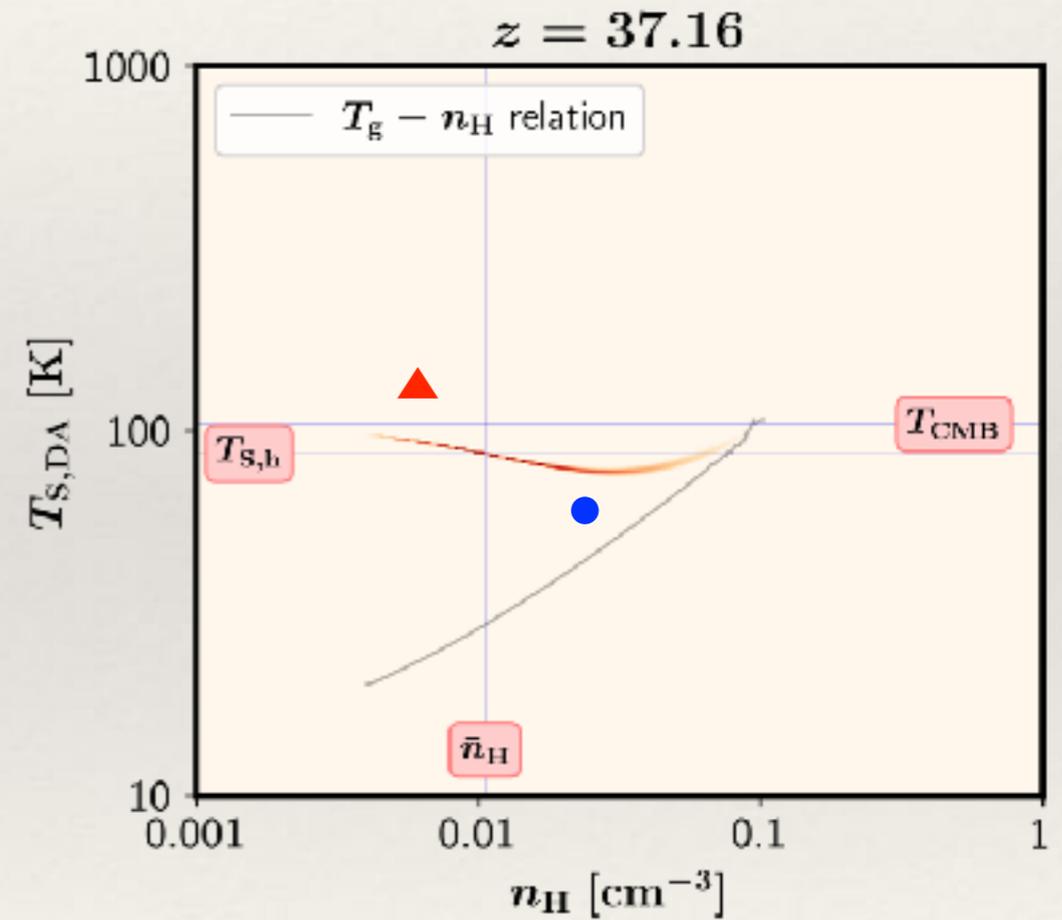
T_S starts to decouple from T_{gas} from $\delta < 0$
→ Positive ΔT_S from both $\delta < 0$ and $\delta > 0$ (▲ & ●)



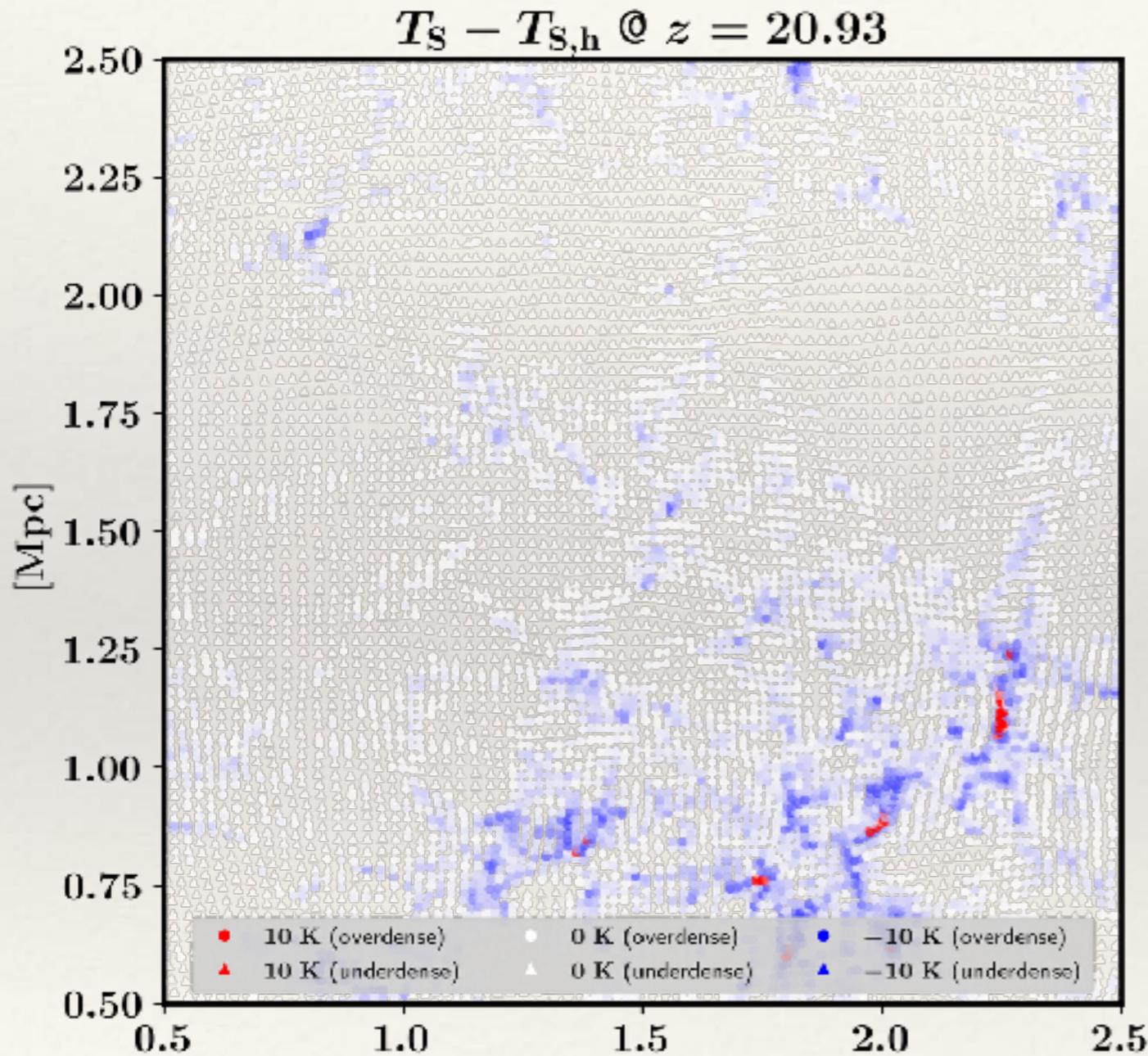
ΔT_S due to Structure Formation ($z < 60$)



Negative relation btw ΔT_S and δ at $z < 60$
as T_S driven by decoupling from T_{CMB} (▲ & ●)

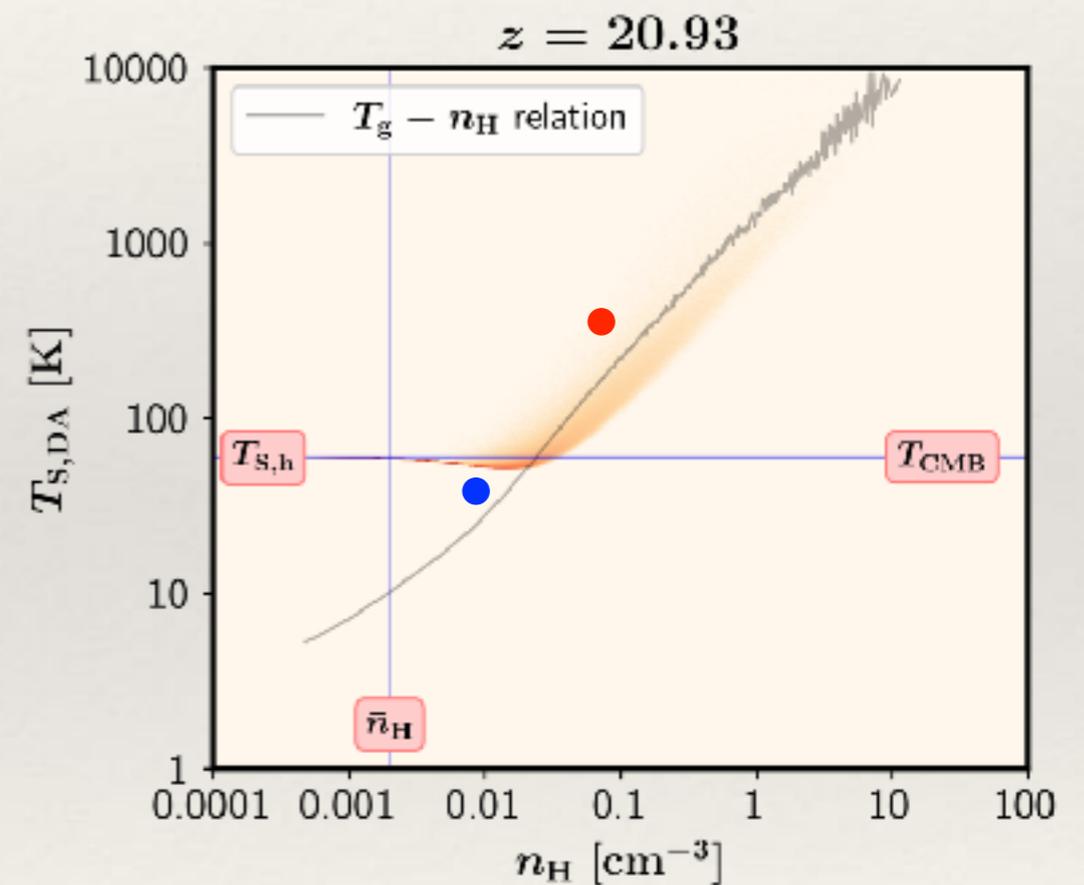


ΔT_S due to Structure Formation ($z < 30$)

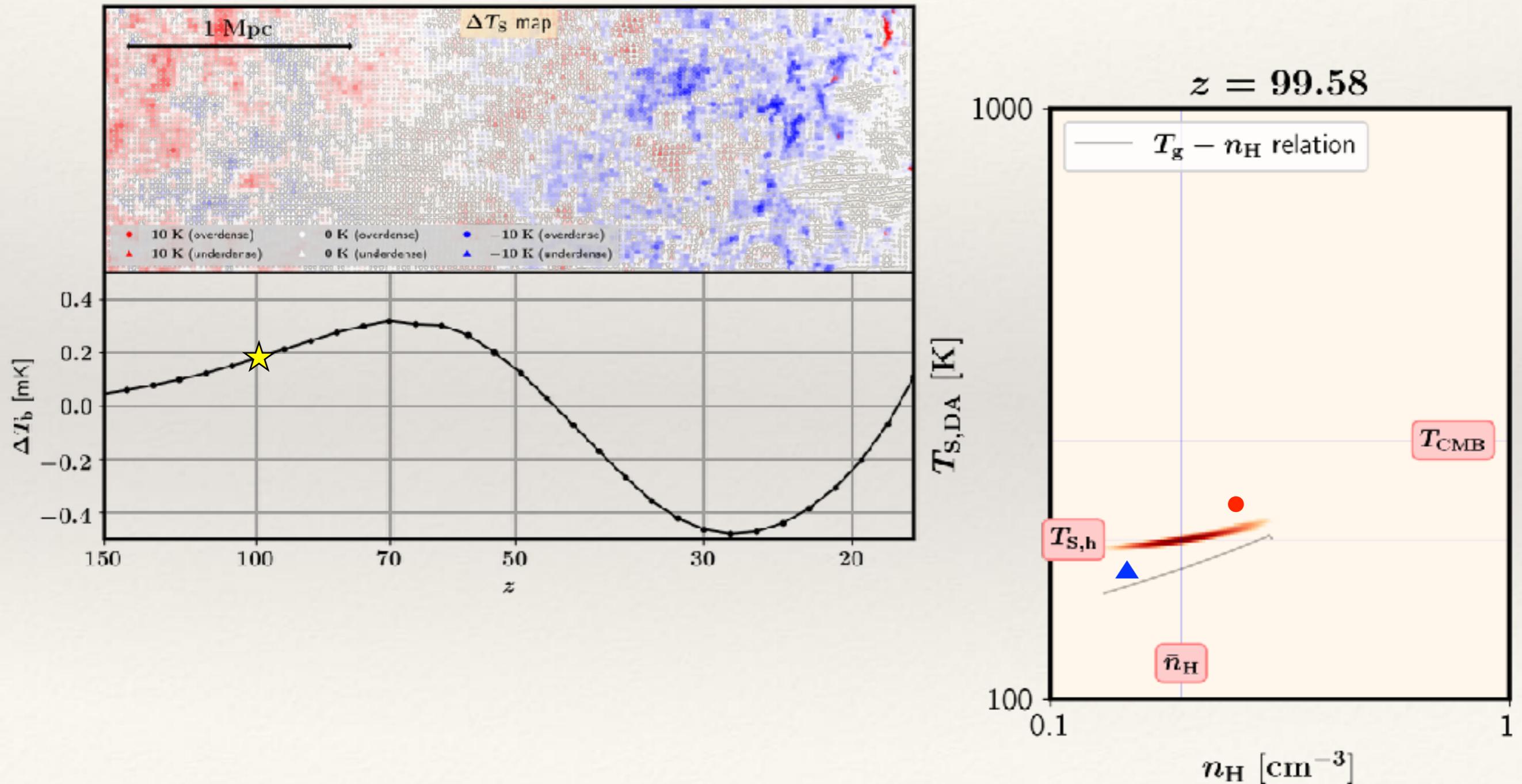


$\Delta T_S < 0$ from $0 < \delta \lesssim 10$ (●) due to $T_{\text{gas}} < T_{\text{CMB}}$

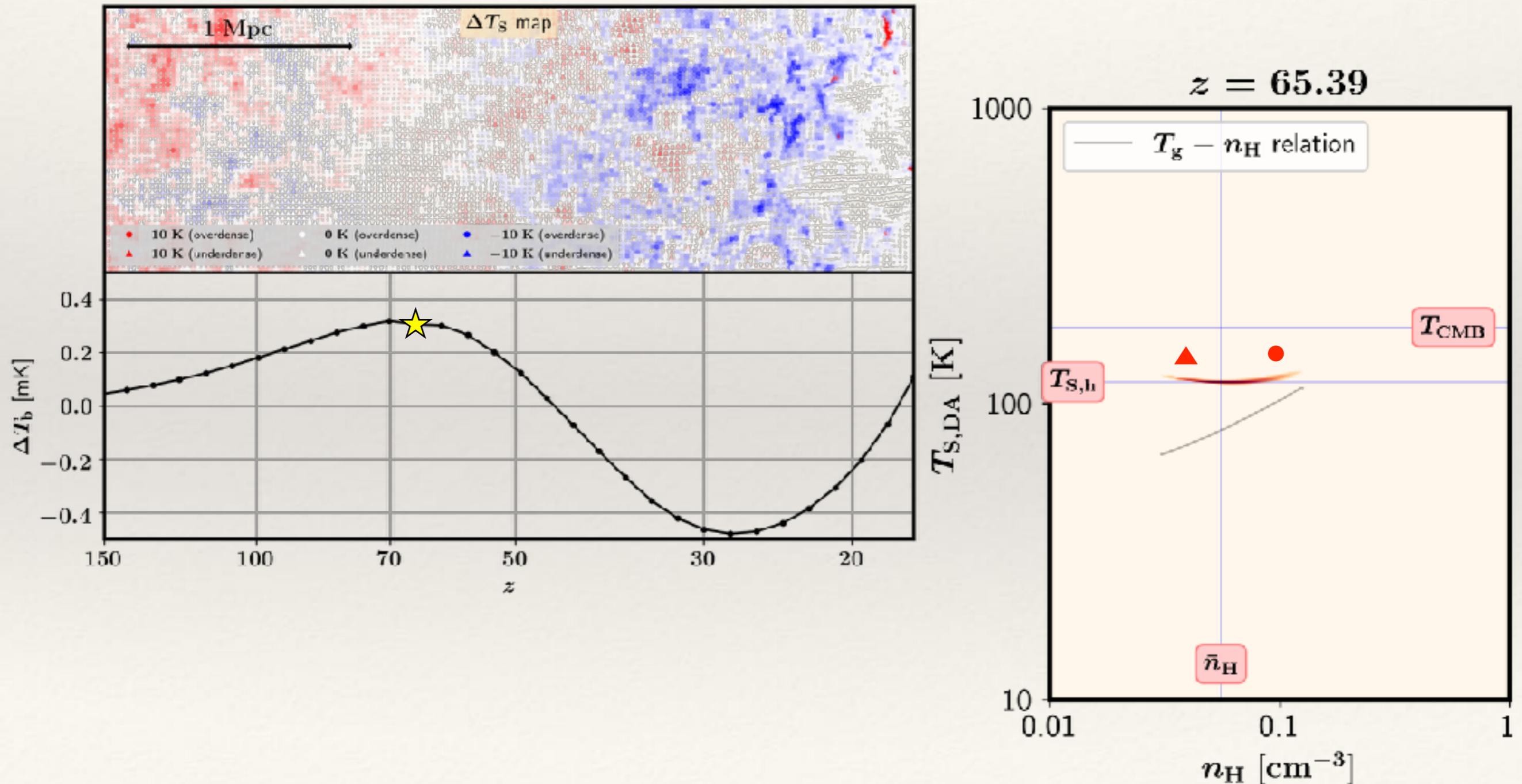
$\Delta T_S > 0$ from $\delta > 10$ (●) due to $T_{\text{gas}} > T_{\text{CMB}}$



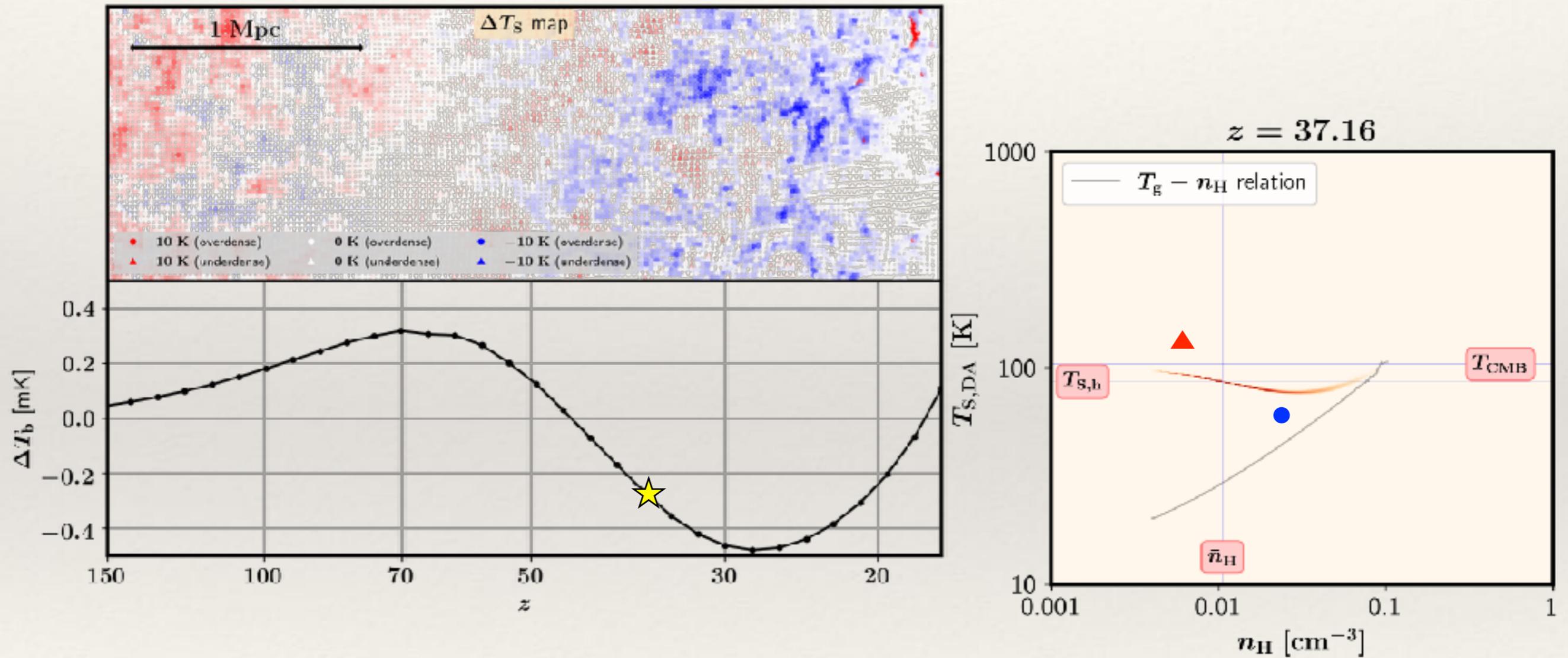
Impact on 21cm Brightness Temperature



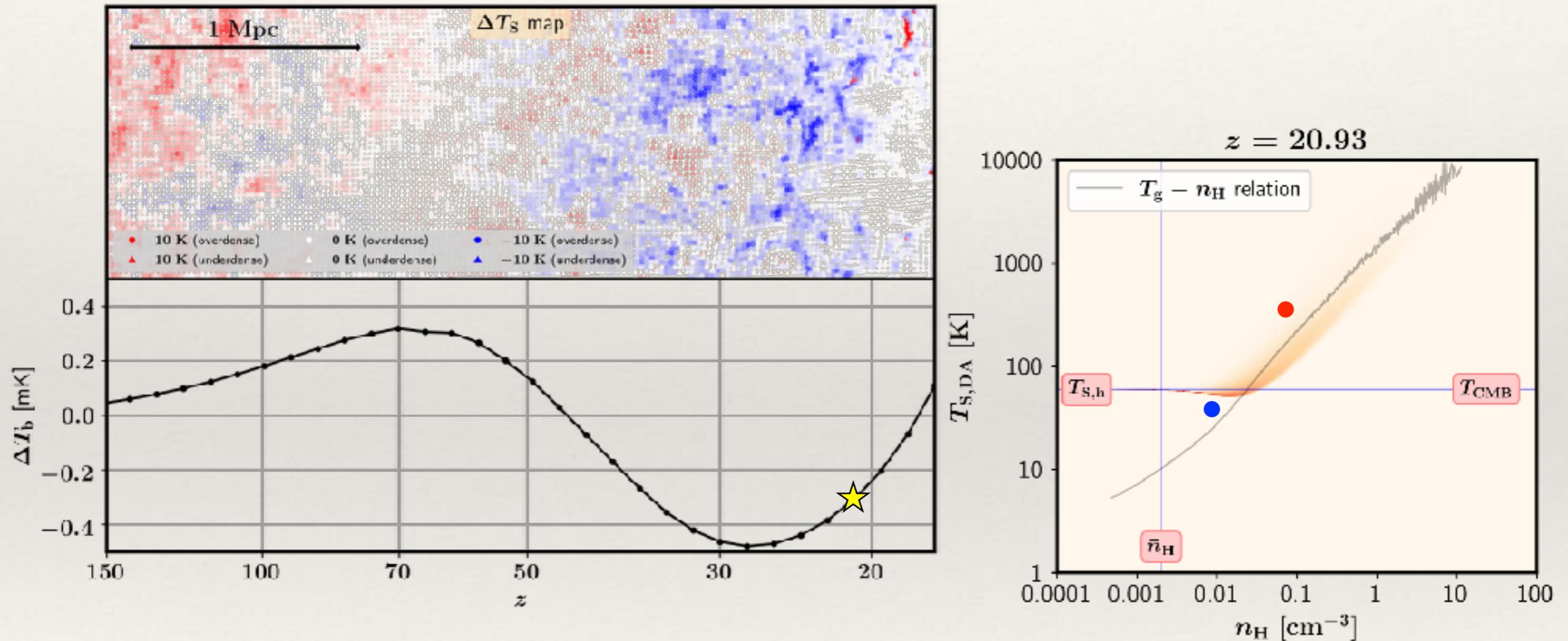
Impact on 21cm Brightness Temperature



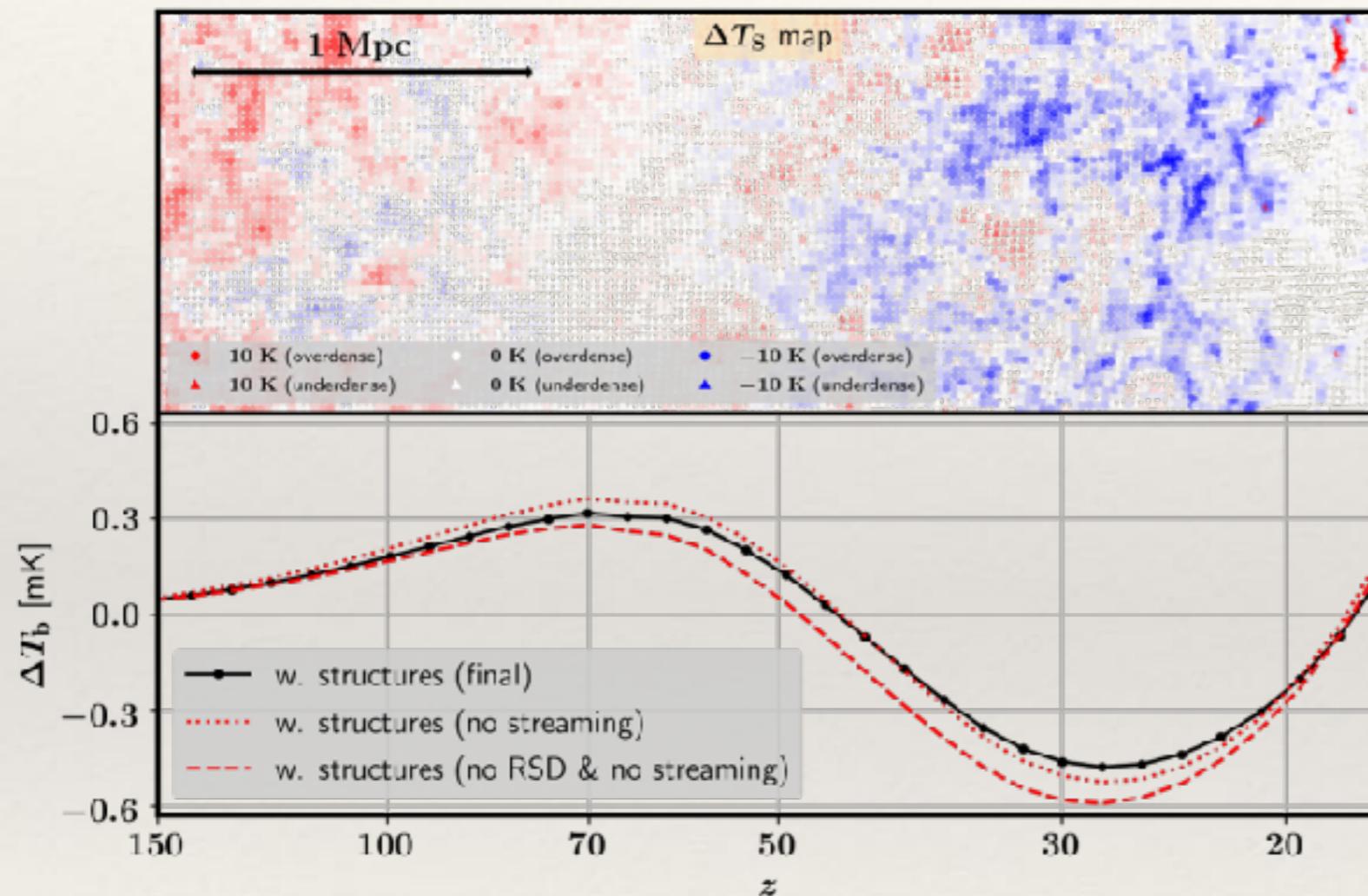
Impact on 21cm Brightness Temperature



Impact on 21cm Brightness Temperature



Impact of Streaming and RSD



$$z = 30$$

$T_b = -4.270$ [mK],
structure formation (X)

$$T_b = -4.270 - 0.583 = -4.853 \text{ [mK]}$$

structure formation (O)

RSD effect (X)

streaming velocity (X)

$$T_b = -4.270 - 0.506 = -4.777 \text{ [mK]}$$

structure formation (O)

RSD effect (O)

streaming velocity (X)

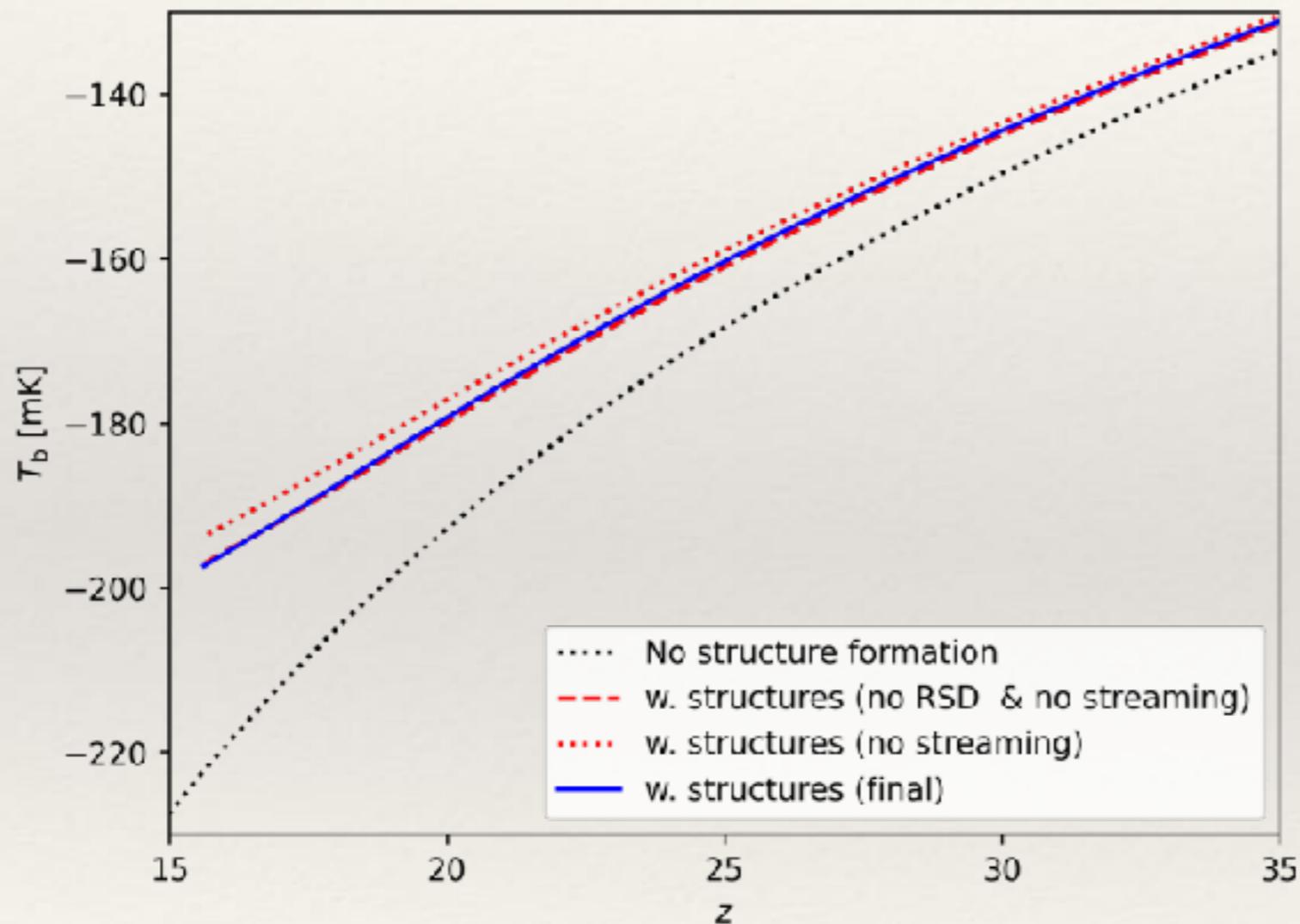
$$T_b = -4.270 - 0.463 = -4.733 \text{ [mK]}$$

structure formation (O)

RSD effect (O)

streaming velocity (O)

Maximal Lya Coupling Case



$z = 16.8$

$T_b = -213.22$ [mK],
structure formation (X)

$T_b = -192.63$ [mK]
structure formation (O)
RSD effect (X)
streaming velocity (X)

$T_b = -189.45$ [mK]
structure formation (O)
RSD effect (O)
streaming velocity (X)

$T_b = -192.56$ [mK]
structure formation (O)
RSD effect (O)
streaming velocity (O)

Did not fully understand the RSD effect yet...

Summary

- ❖ Calculated 21cm signal from the dark ages, accounting for structure formation, RSD effect, and baryon-dark matter streaming velocity.
- ❖ $\sim 10\%$ correction to the homogeneous Universe value at $z \sim 30$
- ❖ Important correction for constraining cosmology parameter with the dark age signal